

Interstellar matter and star formation in W5-E

A *Herschel*★ view

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ABSTRACT

Aims. We identify the young stellar objects (YSOs) present in the vicinity of the W5-E H II region, and study the influence of this H II region on the star formation process in its surrounding molecular material.

Methods. W5-E has been observed with the *Herschel*-PACS and -SPIRE photometers, as part of the HOBYS key program; maps have been obtained at 100 μm , 160 μm , 250 μm , 350 μm , and 500 μm . The dust temperature and column density have been obtained by fitting spectral energy distributions (SEDs). Point sources have been detected and measured using PSF photometry with DAOPHOT.

Results. The dust temperature map shows a rather uniform temperature, in the range 17.5 K – 20 K in the dense condensations or filaments, in the range 21 K – 22 K in the photodissociation regions (PDRs), and in the range 24 K – 31 K in the direction of the ionized regions. The values in the column density map are rather low, everywhere lower than 10^{23} cm^{-2} , and of the order of a few 10^{21} cm^{-2} in the PDRs. About 8000 M_{\odot} of neutral material surrounds the ionized region, which is low with respect to the volume of this H II region; we suggest that the exciting stars of the W5-E, W5-W, Sh 201, A and B H II regions formed along a dense filament or sheet rather than inside a more spherical cloud. Fifty point sources have been detected at 100 μm . Most of them are Class 0/I YSOs. The SEDs of their envelopes have been fitted using a modified blackbody model. These envelopes are cold, with a mean temperature of 15.7 ± 1.8 K. Their masses are in the range $1.3 M_{\odot} - 47 M_{\odot}$. Eleven of these point sources are candidate Class 0 YSOs. Twelve of these point sources are possibly at the origin of bipolar outflows detected in this region. None of the YSOs contain a massive central object, but a few may form a massive star as they have both a massive envelope and also a high envelope accretion rate. Most of the Class 0/I YSOs are observed in the direction of high column density material, for example in the direction of the massive condensations present at the waist of the bipolar Sh 201 H II region or enclosed by the bright-rimmed cloud BRC14. The overdensity of Class 0/I YSOs on the borders of the H II regions present in the field strongly suggests that triggered star formation is at work in this region but, due to insufficient resolution, the exact processes at the origin of the triggering are difficult to determine.

Key words. Stars: formation – Stars: early-type – ISM: dust – ISM: H II regions – ISM: individual objects: W5-E, Sh 201

1. Introduction

W5 is a Galactic H II region located in the Perseus arm, at a distance of about 2 kpc (Sect. 2). It is composed of two adjacent H II regions, W5-E and W5-W, each surrounding an exciting stellar cluster. Due to its simple morphology and its proximity to the Sun, W5 is an excellent laboratory to study the feedback of massive stars on the surrounding material in the context of ongoing star formation.

W5 is a well-studied H II region – what more can we learn from *Herschel* observations? *Herschel* provides the complete, deep image of all the gas available for a new generation of stars. First, the high sensitivity and high resolution far-infrared (FIR) *Herschel* maps trace the emission from cold dust and thus we

can have a better view of the distribution of the dense molecular material surrounding the ionized region. It is precisely this dense material from which new generations of stars may form. Second, the *Herschel* fluxes from 100 μm to 500 μm of the young stellar objects (YSOs) allow us to create well-sampled spectral energy distributions (SEDs). *Herschel* allows us to determine the peak wavelength of the SED, which is especially important for the determination of dust temperature. With a well-sampled SED we may estimate the parameters of YSOs such as the mass of their envelopes, and better constrain their evolutionary stages. Third, with *Herschel* we can detect rare young Class 0 sources and candidate prestellar cores. Such sources are unresolved with *Herschel* and have faint or no *Spitzer* 24 μm counterparts. Thus, with *Herschel* we may create a complete census of the massive YSOs in W5. We wish to learn the prevalence of triggering in the formation of the young sources. We pay special attention to the formation of massive stars.

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This paper is organised as follows: we present the W5 region in Sect. 2; we discuss the *Herschel* observations and data reduction in Sect. 3; we discuss the distribution and characteristics of the neutral material in Sect. 4; we give the method of detection and the estimation of the physical parameters of YSOs in Sect. 5 (we only discuss one zone here and we describe the other star forming zones in the Appendix). We discuss star formation in the W5-E complex in Sect. 6, and conclude in Sect. 7.

2. Presentation of the region

Based on their proximity to one another on the sky and their similar radial velocities, it is often assumed that W5 and the nearby W3 and W4 H II regions are part of the same large star forming complex. A complete description of the W3/W4/W5 complexes is given in Megeath et al. (2008). The distance to W3, 2.00 ± 0.05 kpc, is known accurately from maser parallax measurements (Xu et al. 2006; Hachisuka et al. 2006). W4, which is adjacent to W3, probably lies at the same distance (Megeath et al. 2008). The situation for W5 is more uncertain, as it is separated in angle (by more than one degree) from the two other H II regions. Spectrophotometric observations, however, give very similar distances for the OB clusters IC 1805 and IC 1848 located respectively in W4 and W5 (in the range 2.1 kpc–2.4 kpc; Becker & Fenkart 1971, Moffat 1972, Massey et al. 1995, Chauhan et al. 2009 & 2011a). In the following we adopt a distance of 2.0 kpc for the W5 region; this allows us to compare directly with the results of Koenig et al. (2008; hereafter KOE08), based on *Spitzer* observations.

W5 is an optically visible H II region. Megeath et al. (2008; their fig. 3) show an optical view of the entire region. In their figure the [SII] emission, emitted by the low excitation zones close to the ionization fronts (IFs), clearly shows the presence of two distinct adjacent H II regions; they have been named W5-E and W5-W by Karr and Martin (2003). In the optical W5-E appears almost circular in projection with a diameter of ~ 24 pc, centered near its exciting star HD 18326. At radio wavelengths W5-E has the same aspect and appears almost circular around the exciting star. In Fig. 1 we show the radio-continuum emission at 1.42 GHz (Canadian Galactic Plane Survey, Taylor et al. 2003); the radio-continuum intensity is slightly enhanced in the central regions and on the borders of bright-rimmed clouds (identified on Fig. 2). As shown by Fig. 1 the W5-E and W5-W H II regions are enclosed in two adjacent bubbles, observed by *Spitzer* at $8.0 \mu\text{m}$ (KOE08, their figure 4; the $8.0 \mu\text{m}$ band is dominated by the emission from polycyclic aromatic hydrocarbons - PAHs - from the photo-dissociation regions or PDRs). The W5-E H II region is slightly extended in the north-south direction (diameter ~ 30 pc). This is likely due to a density gradient (the density being higher in the north than in the south). This density gradient would also explain why the $8 \mu\text{m}$ shell surrounding the ionized region opens towards the south. This is discussed in Sect. 4.

The main exciting star of W5-E, HD 18326 (also BD+59°0578), has a spectral type of O7(± 0.5)V (Chauhan et al. 2011a and references therein). Its coordinates are $\alpha(2000) = 02^{\text{h}}59^{\text{m}}23.171^{\text{s}}$, $\delta(2000) = 60^{\circ}33'59.50''$ (or $l = 138^{\circ}02'59.70''$, $b = +1^{\circ}50'02.43''$; Maíz-Apellániz et al. 2004). Its colour $B-V = 0.362$ and its spectral type indicate a colour excess $E(B-V) = 0.632$ or a visual extinction $A_V = 1.96$ mag. HD 18326 is part of a rich cluster studied by Chauhan et al. (2011a). They estimate a mean age of 1.3 Myr for the cluster, and a varying extinction in the range 1.92 mag–2.50 mag. This gives an estimate of the interstellar extinction in the direction of W5-E.

Heyer and Terebey (1998) show large-scale maps of the $^{12}\text{CO}(1-0)$ emission over the entire W3-W4-W5 complex, obtained with a resolution of $45''$. Their figure 2 shows the molecular material associated with W5-E in the channels at -37.5 km s^{-1} and -42.5 km s^{-1} ; this material is mainly located in the northern areas of the bubble (thus the suggested density gradient). (See also figure 2 in Karr and Martin 2003, based on the same observational data.) Detailed maps of CO emission can be found in Carpenter et al. (2000; $^{13}\text{CO}(1-0)$ maps) and in Karr and Martin (2003; $^{12}\text{CO}(1-0)$ maps). These maps show an accumulation of molecular material in thin layers following the ionization fronts, and in some well defined regions like inside the bright-rimmed cloud BRC14, near the H II region Sh 201, and in the region between W5-E and W5-W; these regions are identified in Fig. 2. High resolution observations of W5-E (HPBW = $15.6''$) have been obtained in the $^{13}\text{CO}(1-0)$ and $\text{C}^{18}\text{O}(1-0)$ lines by Niwa et al. (2009). Eight clouds are identified (three of them are associated with the bright-rimmed clouds BRC12, BRC13, and BRC14), with masses in the range $460 M_{\odot}$ to $36,000 M_{\odot}$ ¹.

Three bright-rimmed clouds (BRCs) are located on the border of W5-E, BRC12, BRC13, and BRC14 (Fig. 2). Their morphology and young stellar content have been the subject of numerous studies whose aim was to establish if the star formation observed in these clouds was triggered by the nearby W5-E H II region. Using a combination of radio, IR, submillimeter, and CO observations, Morgan et al. (2009, and references therein) conclude that these clouds are excellent candidates for triggered star formation. A similar conclusion is reached by Chauhan et al. (2011b), based on age determination and on the observed small scale sequential star formation. These regions will be discussed individually in Sect. 5 or in the Appendix.

The W5 region has been observed by *Spitzer* with the IRAC and MIPS instruments. KOE08 used these data to discuss star formation in W5. They produced a catalogue containing 18,518 stellar objects, among which are 2,064 Class I and Class II YSOs. These YSOs are identified and classified using their *Spitzer* colours, after an extinction correction based on 2MASS data. These authors showed that Class I and Class II YSOs are not distributed uniformly, but are primarily in clustered or filamentary structures. They also showed that within the cavity carved by the H II region, the ratio of Class II (older) YSOs to Class I (younger) YSOs is ~ 7 times higher than within the molecular material at the periphery of the ionized gas. KOE08 attributed this result to an age difference between sources in these two locations, the formation of the younger objects possibly being triggered by the W5-E H II region. We discuss these results in Sect. 6.

3. Herschel: observations and reduction

3.1. Observations with Herschel

W5-E was observed with the PACS photometer (Poglitsch et al. 2010) as part of the HOBYS key program (Motte et al. 2010) on 23 February 2010. The PACS photometer was used to make simultaneous observations in two photometric bands ($100 \mu\text{m}$ and $160 \mu\text{m}$). Two cross-scan maps were performed at angles

¹ These high masses are somewhat puzzling: Karr and Martin (2003) estimate the mass of the molecular material in the whole W5 region (including both W5-E and W5-W) to be $44,000 M_{\odot}$, a large fraction of this material being associated with W5-W. In addition, the most massive clouds in Carpenter et al. (2000) are about ten times less massive than those in Niwa et al. (2009).

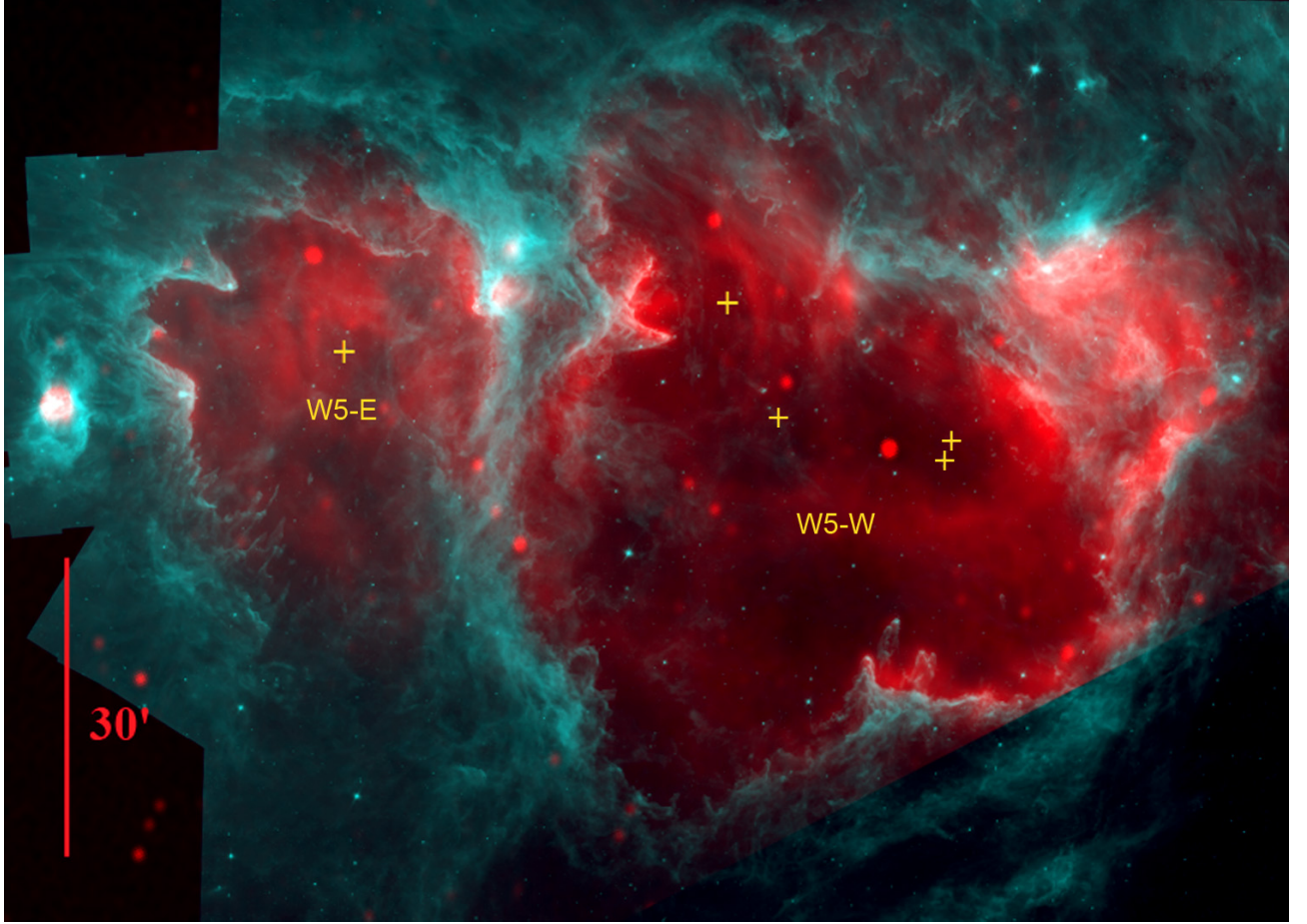


Fig. 1. Composite colour image of the W5 region. Red is the radio-continuum emission of the ionized gas at 1.42 GHz (CGPS map); turquoise is the *Spitzer*-IRAC 8.0 μm PAH emission. The two 8.0 μm bubbles surround the H II regions W5-E and W5-W. The yellow crosses lie at the position of the exciting stars of W5-E and W5-W. Most of the compact radio sources in the field are extragalactic sources. North is up and east is left.

of 45° and 135° with a scanning speed of $20''/\text{second}$. The maps cover an area of $70' \times 70'$ and are centered at $\alpha(2000) = 02^{\text{h}}59^{\text{m}}26^{\text{s}}$, $\delta(2000) = 60^\circ31'05''$. The total observing time was 3.6 hours.

W5-E was observed with the SPIRE photometer (Griffin et al. 2010, Swinyard et al. 2010) as part of the key program “Evolution of Interstellar Dust” (Abergel et al. 2010) on 11 March 2010. The SPIRE photometer was used to make simultaneous photometric observations in the three photometric bands (250 μm , 350 μm and 500 μm). The scan speed was $30''/\text{second}$. Cross-linked scanning is achieved by scanning at angles of 42° and -42° (one map is obtained for each scanning angle). The maps cover an area of $65' \times 60'$ and are centered at $\alpha(2000) = 02^{\text{h}}59^{\text{m}}25^{\text{s}}.71$, $\delta(2000) = 60^\circ32'32''.1$. The total observing time is 1.2 hours.

These data are described and used in Anderson et al. (2012; hereafter AND12). AND12 (their Sect. 3) describes more fully these observations and their subsequent reduction using the HIPE (version 7.1; Ott 2010) and *Scanamorphos* (version 9; Roussel 2012) softwares. We use the AND12 maps.

The *Herschel* maps contain a number of point sources that have counterparts in *Spitzer* maps. The comparison of the positions of the 100 μm point sources with that of their 24 μm counterparts has shown a shift of a few arcsec ($\sim 9.4''$ in right ascension and $0.8''$ in declination) between the two sets of coor-

dinates. A smaller shift also affects the SPIRE maps. These shifts have been corrected, so that our coordinates are now compatible with the *Spitzer* data (see Sect. 3.1 in AND12).

3.2. Photometry of the YSOs

In what follows we measure the flux of YSO candidates detected by *Herschel*. We consider all sources that are not resolved in the PACS 100 μm band as YSO candidates, and analyse these sources at all *Herschel* wavelengths. We pay special attention to the YSOs classified as Class I by KOE08, to see if they have *Herschel* counterparts.

We measure *Herschel* fluxes using the DAOPHOT stellar photometry package with PSF fitting (Stetson 1987), in addition to aperture photometry on bright isolated sources. DAOPHOT, which was designed for crowded fields, is over-performing for the *Herschel* images containing a few tens of sources. These fields, however, are very difficult to reduce because we often find sources located in the photodissociation regions surrounding the ionized ones, in regions of bright and highly spatially variable background emission. We use DAOPHOT in an iterative process. In the first step we use bright isolated sources to construct a PSF at each photometric band. We use this PSF to obtain a first estimate of the source positions and magnitudes. Second, we subtract the identified sources from the original frame (# 1). We smooth the resulting image with a median filter (with a filter

window of 15 pixels×15 pixels, larger than the defects resulting from the first imperfect reduction; this gives an image, at lower resolution, of the background emission alone. We subtract this image from the original one. The resulting image (# 2) shows the point sources superimposed on a more uniform background (the only remaining low-brightness features are smaller than the filter window). In our third step we run DAOPHOT on the improved image # 2. We subtract the identified sources from the original #1 image to determine if the reduction has been improved. We check the results by eye: if the reduction is good no brightness holes or peaks are seen at the position of the sources on the subtracted image. We could in principle iterate this process, but we find that it is not necessary.

The frames on which we have performed the DAOPHOT reduction on have a pixel size of 1''.7 (PACS 100 μm), 2''.85 (PACS 160 μm), 4''.5 (SPIRE 250 μm), 6''.25 (SPIRE 350 μm), and 9''.0 (SPIRE 500 μm). The HPBW (of the PSF) given by DAOPHOT is always slightly larger than the resolution (FWHM) given in the Observers' Manuals: $\sim 7''.5$ at 100 μm , $\sim 12''.5$ at 160 μm , $\sim 25''$ at 250 μm , $\sim 34''$ at 350 μm , and $\sim 48''$ at 500 μm . This is probably the signature of non-Gaussian PSFs and of our choice for the value of the parameter "FITTING RADIUS" in DAOPHOT.

We also measure the *Herschel* fluxes of the isolated stars used to construct the PSF using aperture photometry, employing a circular aperture of radius 8 pixels (at all wavelengths the PSF radius is slightly larger than 2 pixels). We estimate the underlying background using an annular aperture centered on the source of inner radius 8 pixels and outer radius 10 pixels. We apply an aperture correction to take into account the fact that: i) all the flux is not enclosed in an aperture of radius 8 pixels; ii) a few percent of the source flux is present in the annular zone used to estimate the background. These corrections are of the order of 18% at 100 μm , 15% at 160 μm (based on the PACS instrumental manual), 10% at 250 μm , 8% at 350 μm , and 7% at 500 μm (our own estimates). The comparison between the DAOPHOT and the aperture magnitudes allow us to estimate the accuracy of our measurements. The dispersion is ~ 0.1 mag (or 10% of the flux) at 100 μm and 250 μm , 0.15 mag (15% of the flux) at 160 μm , 0.2 mag (20%) at 350 μm , and 0.3 mag (30%) at 500 μm . These numbers are for rather bright sources or for fainter ones superimposed on a flat background. The accuracy is the photometry measurements is probably worse for faint sources superimposed on a bright and variable background, such as those located in the vicinity of PDRs. This shows the difficulty of such measurements.

4. The dense neutral material associated with W5-E

The overall morphology of W5 is more complicated than a simple spherical morphology around the central exciting star. Fig. 2 is a composite colour image of the observed field that shows the overall morphology of this region and allows us to identify its different components. In red we show the cold dust emission at 250 μm from the *Herschel*-SPIRE data; these data trace molecular material associated with the region. In green we show the 100 μm dust emission from the *Herschel*-PACS data; these data trace warmer dust located in the PDRs surrounding the ionized regions or associated with massive YSOs. In blue we show the $\text{H}\alpha$ emission of the ionized gas from the DSS2-red frame. Fig. 2 clearly shows that the dust bubble is opened towards the south. This configuration was already apparent in the CO maps and *Spitzer* images. In the north there are two parallel ionization fronts and associated PDRs, running from east to west, with ionized gas between the two. At the periphery of the bubble we

Table 1. Opacity of the dust at *Herschel*-SPIRE wavelengths

λ (μm)	OH1 ($\text{cm}^2 \text{g}^{-1}$)	OH2 ($\text{cm}^2 \text{g}^{-1}$)	Equation (2) ($\text{cm}^2 \text{g}^{-1}$)
250	21.1	17.5	14.4
350	11.2	10.1	7.3
500	5.5	5.0	3.6
850	1.95	1.8	1.25
1300	1.0	0.9	0.5

find bright-rimmed structures enclosing bright sources such as YSOs, H II regions, and small clusters.

4.1. Determination of the mass and column density

It is generally assumed that the emission from cold dust at *Herschel*-SPIRE wavelengths is optically thin. This property allows us to estimate the distribution of the dense molecular material hosting this cold dust, and, via various assumptions, to determine the H_2 column density, $N(\text{H}_2)$, and the mass of various structures (condensations, filaments, etc...; see AND12). The derivation and the limitations of such determinations can be found in Hildebrand (1983). To summarize, the total (gas+dust) mass of a feature is related to its flux density S_ν by:

$$M_{(\text{gas+dust})} = 100 \frac{S_\nu D^2}{\kappa_\nu B_\nu(T_{\text{dust}})}, \quad (1)$$

where D is the distance to the source, κ_ν is the dust opacity per unit mass of dust at the frequency ν , and $B_\nu(T_{\text{dust}})$ is the Planck function for a temperature T_{dust} . For all estimates in this paper, we assume a gas-to-dust ratio of 100. This value is uncertain and the subject of discussion. The dust opacities however are even more uncertain. As discussed by Henning et al. (1995) the dust opacity depends on the size, shape, chemical composition, physical structure, and temperature of the grains, thus of the dust environment. This has been discussed for example by Martin et al. (2011). Table 1 gives the opacities at the *Herschel*-SPIRE wavelengths estimated from various models. The values obtained by Ossenkopf and Henning (1994; OH1 and OH2) are in columns 2 and 3; they correspond respectively to their models "thick ice mantle" and "thin ice mantle" (density 10^6 cm^{-3} and age $\geq 10^5$ yr). We calculated the values in column 4 using the formula:

$$\kappa_\nu = 10 \times \left(\frac{\nu}{1000 \text{ GHz}} \right)^2, \quad (2)$$

which assumes a spectral index of 2 for the dust emissivity, and gives at 1.3 mm the dust opacity recommended by Preibisch et al. (1993) for cloud envelopes. The thick ice mantle model (OH1) is possibly more appropriate for cold and dense cores (see also Henning et al. 1995)². In the following we use the opacities given in the last column of Table 1 to estimate the mass and column density, but we note that, due to the unknown properties of the dust grains, the dust opacities and thus the derived clouds' masses are uncertain by at least a factor ~ 2 .

Here we use the map of dust temperature over the W5-E region, similar to the map shown in AND12, using PACS and

² A dust opacity of $4.3 \text{ cm}^2 \text{g}^{-1}$ at 250 μm has been estimated for dust in the Galactic ISM at high latitude; two to four times larger values are estimated for dust in the diffuse ISM of the Galactic plane (Martin et al. 2011), showing that the dust properties probably vary with the environment. Little is known about the opacity of dust in dense cores or envelopes.

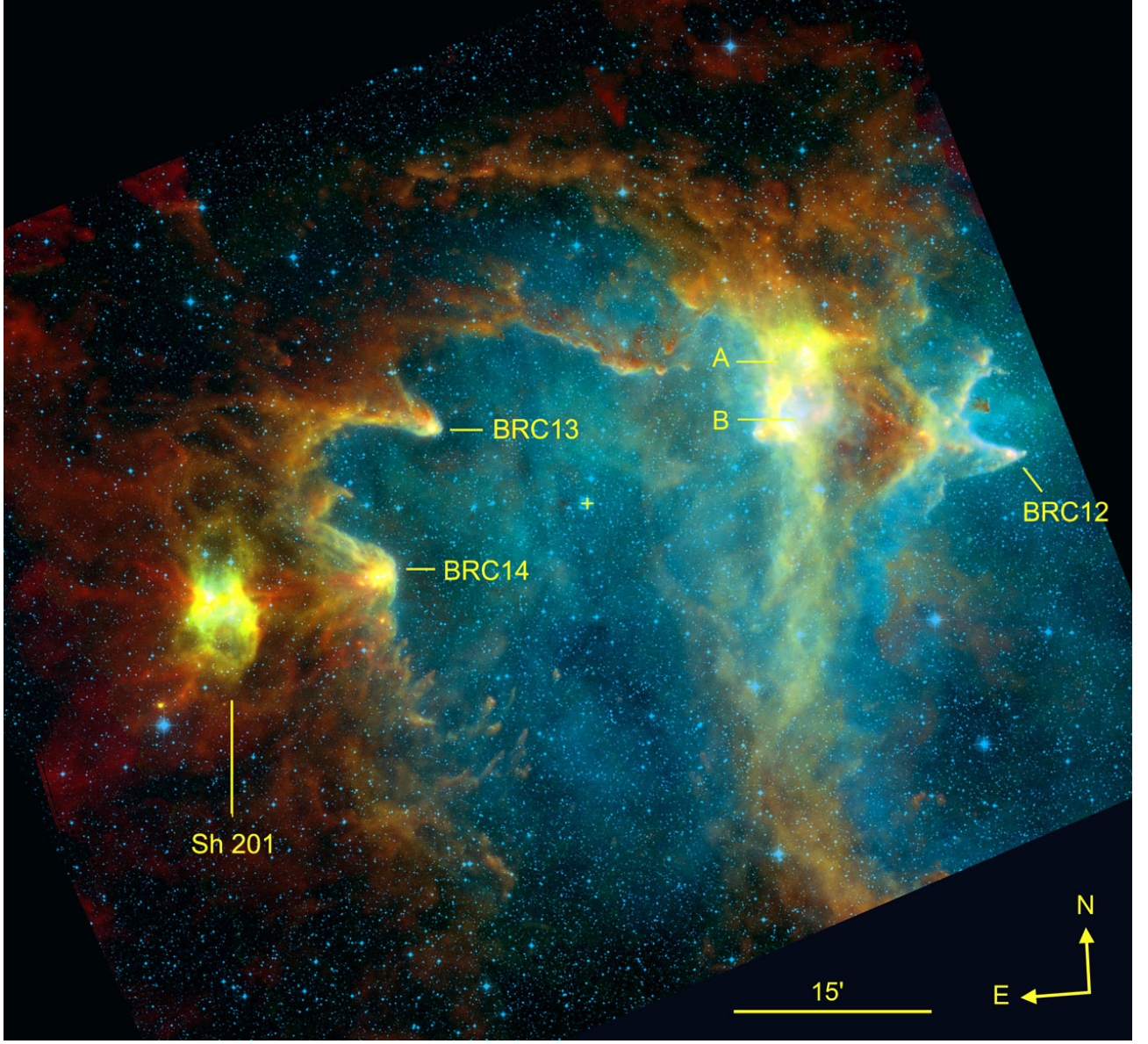


Fig. 2. Composite colour image of the W5-E field observed by *Herschel*. Red is the 250 μm *Herschel*-SPIRE data, which traces cold dust emission, green is the 100 μm *Herschel*-PACS data, which traces emission mainly from dust at higher temperature located in the PDRs, and blue is the DSS2-red survey, which traces the $\text{H}\alpha$ emission of the ionized gas. The yellow cross lies at the position of HD 18326, the exciting star of W5-E.

SPIRE data and a simple model of modified blackbody emission. While AND12 fitted simultaneously for the dust temperature and the column density, we here use the *Herschel* fluxes to derive a temperature at each location, leaving the column density as a free parameter. We smoothed the *Herschel* images to the same resolution using a two-dimensional Gaussian representing the SPIRE beam at 500 μm and reprojected the data to a uniform pixel spacing of the 500 μm data. Fig. 3 shows the temperature map and some iso-temperature contours; cold dust appears in white, hot dust in black. The contours correspond to 18 K, 20 K, 22 K, 24 K and 26 K. The temperature does not vary strongly over the field: T_{dust} lies in the range 17 K to 31 K. Temperatures higher than 26 K are observed in the central cavities of W5-E and W5-W, in the direction of Sh 201, in the direction of the small A and B H II regions lying between W5-E and W5-W, and in the direction of the IR sources inside BRC14. Due to the relatively low fluxes, the temperatures in the direction of the interior of W5-E

are especially uncertain. Temperatures in the range 17 K–20 K are observed in zones of high column density such as the filaments in the vicinity of Sh 201 or inside BRC12, BRC13, and BRC14. The PDR regions, adjacent to the IFs, are at an intermediate temperature of 21 K–22 K.

The knowledge of the dust temperature allows us to estimate the column density of the gas. With the same assumptions, we can estimate the H_2 column density from the surface brightness F_ν , using the formula:

$$N(\text{H}_2) = \frac{100 F_\nu}{\kappa_\nu B_\nu(T_{\text{dust}}) 2.8 m_{\text{H}} \Omega_{\text{beam}}}, \quad (3)$$

where F_ν is expressed in Jy beam^{-1} , B_ν in Jy, the factor of 2.8 is the mean molecular weight, m_{H} is the hydrogen atom mass, and Ω_{beam} is the beam solid angle. The visual extinction can be estimated from the column density. From the classical relations, $N(\text{H} + \text{H}_2)/E(B - V) = 5.8 \times 10^{21} \text{ particles cm}^{-2} \text{ mag}^{-1}$ (Bohlin

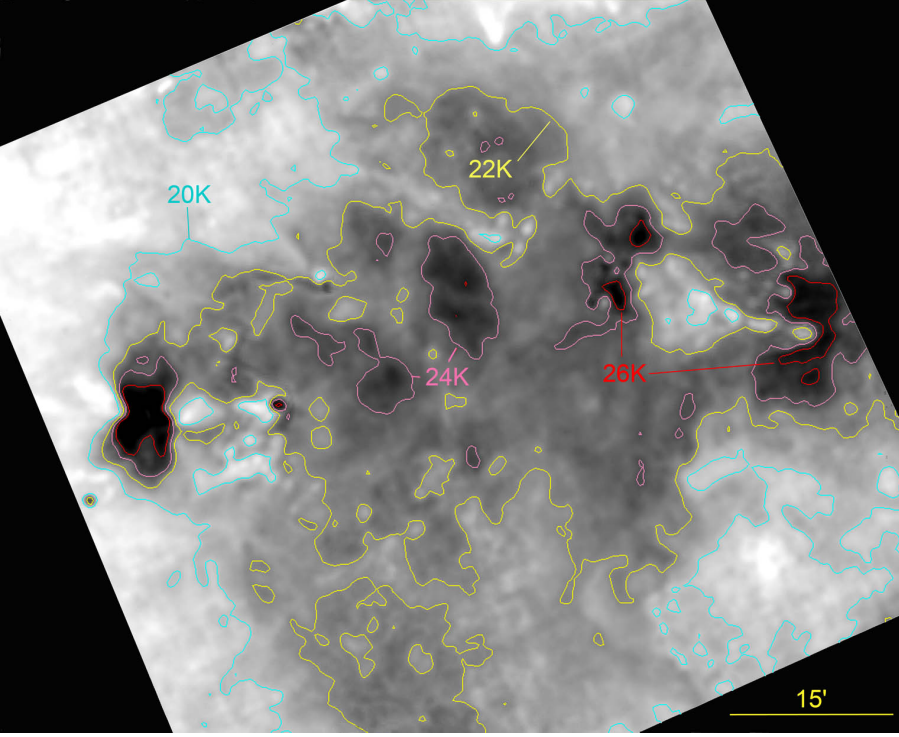


Fig. 3. The temperature map of W5. Iso-temperature contours (20 K in blue, 22 K in yellow, 24 K in pink, and 26 K in red) are superimposed on the grey-scale temperature map. The resolution is that of the 500 μm map, 37". North is up and east is left.

et al. 1978) and $A_V = 3.1 E(B - V)$, we obtain $A_V = 5.34 \times 10^{-22} N(\text{H} + \text{H}_2)$ or $A_V = 1.07 \times 10^{-21} N(\text{H}_2)$ in a molecular medium.

We used the SPIRE 250 μm image, which has the highest angular resolution at SPIRE wavelengths, to create a map of the gas column density³. We reprojected the temperature map previously obtained to the 4.5" pixel grid of the 250 μm image and used a dust opacity of 14.4 $\text{cm}^2 \text{g}^{-1}$ (Table 1). In Figure 4 we show the column density map. The column density does not reach very high values; the highest values are found in the direction of the cloud enclosed by BRC14 ($9.0 \times 10^{22} \text{cm}^{-2}$, or $A_V = 96 \text{mag}$), in the cloud east of Sh 201 ($\sim 6.2 \times 10^{22} \text{cm}^{-2}$, or $A_V = 66 \text{mag}$), inside BRC13 ($1.7 \times 10^{22} \text{cm}^{-2}$, or $A_V = 18.2 \text{mag}$), or in the vicinity of the B H II region ($\sim 9.6 \times 10^{21} \text{cm}^{-2}$). Everywhere else in the field the column density is lower.

Assuming the same kind of dust previously described, we can use the column density to estimate the optical depth of the dust at 100 μm . Our maximum column density $N(\text{H}_2) \sim 10^{23} \text{cm}^{-2}$ corresponds to an optical depth $\tau_{100\mu\text{m}} \sim 0.42$. The column density is everywhere much lower; thus we consider in the following that the dust is optically thin at 100 μm . (This justifies using the 100 μm fluxes to estimate the dust temperature.)

5. Star formation: YSOs and dense condensations

The protostars were originally grouped into three "Classes" according to the spectral index α of their SED, measured in the range 2.2 μm –10 μm (Lada 1987; Greene et al. 1994):

$$\alpha = \frac{d \log(\lambda F_\lambda)}{d \log(\lambda)}, \quad (4)$$

Class I sources have $\alpha > 0.3$; they are highly embedded protostars whose luminosity is dominated by a spherical infalling envelope. Class II YSOs, with $-2 < \alpha < -0.3$, are protostars surrounded by a substantial accreting disk. Class III sources having

³ The column density map in AND12 was created from the 500 μm (lower resolution) map using a slightly different method.

$\alpha < -2$ are young stars having dissipated most of their disk. Sources with $-0.3 \leq \alpha \leq +0.3$ are called flat spectrum sources.

Yet, Robitaille et al. (2006) showed that the same object can be classified differently depending on the viewing angle; for example, a Class II source with an edge-on disk can have a positive spectral index that would make it resemble a Class I object. Thus these authors prefer to discuss the evolutionary stages of their model by adopting a "Stage" classification analogous to the Class scheme, but based on the physical properties of the YSO (e.g., based on disk mass or envelope accretion rate) rather than properties of its SED (e.g., slope). Stage 0 and I objects (which they do not differentiate) have significant infalling envelopes and possibly disks, Stage II objects have optically thick disks (and the possible remains of a tenuous infalling envelope), and Stage III objects have optically thin disks. The boundaries between the different stages are somewhat arbitrary, in the same way as the Class scheme. They define Stage 0/I objects as those that have $\dot{M}_{\text{env}}/M_* > 10^{-6} \text{yr}^{-1}$, Stage II objects as those that have $\dot{M}_{\text{env}}/M_* < 10^{-6} \text{yr}^{-1}$ and $\dot{M}_{\text{disk}}/M_* > 10^{-6}$, and Stage III objects as those that have $\dot{M}_{\text{env}}/M_* < 10^{-6} \text{yr}^{-1}$ and $\dot{M}_{\text{disk}}/M_* < 10^{-6}$.

Class 0 sources were added later to the Class scheme (André et al. 1993, 2000). They are younger than Class I sources (younger than 10^5yr). They differ from the Class I sources by the shape of their SEDs, which resembles a cold single-temperature blackbody. Class 0 protostars have $M_*/M_{\text{env}} < 1$ (or $L_{\text{sub-mm}}/L_{\text{bol}} > 10^{-2}$; see Sect. 6.3.3).

In what follows we comment on the point-like sources detected by *Herschel*, and on the Class I sources identified by KOE08. We discuss separately several well-known regions, some associated with small H II regions and some others associated with bright-rimmed clouds (BRCs). Three well known BRCs are present in the field observed by *Herschel*; they are BRC12, BRC13, and BRC14 (containing the IR source AFGL 4029). The H II regions are Sh 201, region A and region B (that we discuss in "The region between W5-E and W5-W",

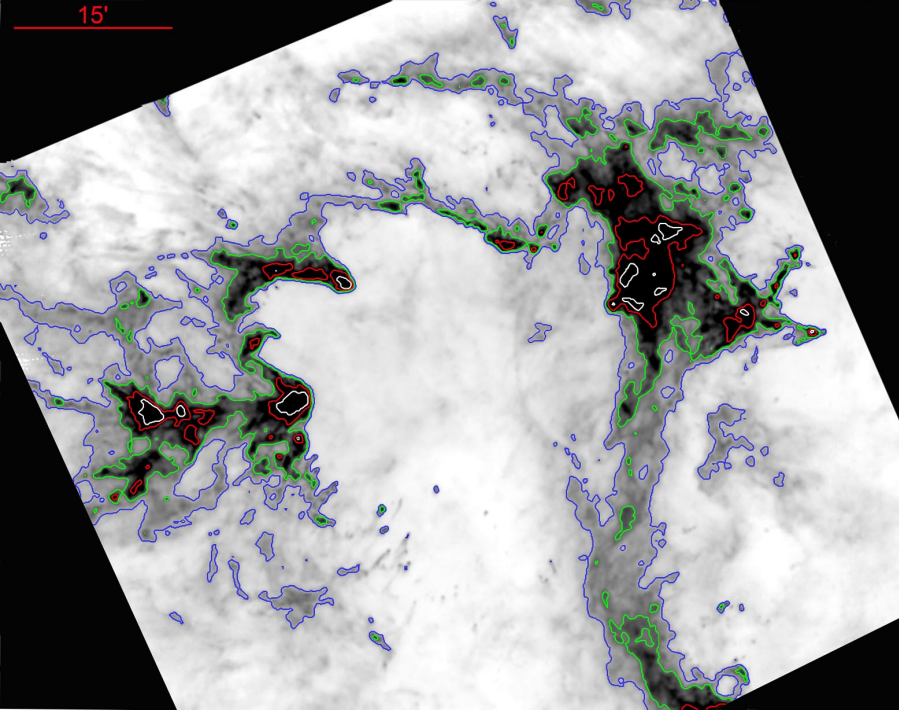


Fig. 4. The column density map of W5. Contours of equal column density ($0.5 \times 10^{21} \text{ cm}^{-2}$ in blue, 10^{21} cm^{-2} in green, $2.5 \times 10^{21} \text{ cm}^{-2}$ in red, and $5 \times 10^{21} \text{ cm}^{-2}$ in white) are superimposed on a grey-scale column density map. The resolution is that of the *Herschel* 250 μm data, $18''$. North is up and east is left.

Appendix D). All these regions are identified on Fig. 2. In this section we present the BRC12 region to illustrate how we conduct the study of the individual regions. We give the same details about all the other zones in Appendixes A to E.

The following section, and the corresponding sections in the Appendix share the same format:

1. We describe each region.
2. We describe and comment on the *Herschel* data. We identify in a figure the *Herschel* point sources discussed in the text. Often the detected point sources have a *Spitzer* counterpart (at least at 24 μm), and are listed in the KOE08 catalogue. For each region a first table (“#1”) gives, according to KOE08, the identification number (column 2), the coordinates (columns 3 and 4), the 2MASS magnitudes (columns 5 to 7), the *Spitzer*-IRAC and MIPS magnitudes (columns 8 to 12), and the classification (column 13) of the *Spitzer* counterpart. Columns 14 to 18 of the same table contain the measured *Herschel* fluxes.
3. Whereas the classifications of KOE08 are based on IR colours and magnitudes up to 24 μm , we try to take advantage of the now more complete SEDs and use the SED fitting-tool of Robitaille et al. (2007; hereafter ROB07) to estimate the evolutionary stage of the sources. We list the parameters necessary to estimate the evolutionary stage (the mass of the central source, the mass and accretion rate of the disk, and the accretion rate of the envelope) in a second table (“#2”); we consider only the models with $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 ; the first numbers in table #2 correspond to the best model, and the range of values in brackets corresponds to the other selected models. Some limitations of the ROB07 SED fitting tool are discussed in Robitaille (2008). Offner et al. (2012) also discuss the accuracy of protostellar properties inferred from the SEDs; they caution us against the use of fits to the Robitaille grid of models for constraining the stellar and disk properties of unresolved protostellar sources. Thus, one must be very cautious when considering the parameters given in tables #2. A large fraction of the *Herschel* point-

sources have a dominant envelope. We show in Sect. 6.3.1 that these envelopes are not well treated in ROB07 models. We prefer to use a modified blackbody model to obtain the temperatures and masses of the envelopes; these are given in Sect. 6.3.1. This situation probably has some repercussions on the disk’s parameters obtained with the SED fitting tool; we often find that the parameters of the disk are not well constrained. In turn, this has some repercussions on the luminosity⁴ of the central object, and thus on its mass⁵.

4. Keeping this in mind, we show the SEDs of some interesting sources, and comment on individual objects.

Some point sources are only observed at *Herschel* wavelengths. These are potentially the most interesting sources as they may be Class 0 YSOs or starless cores (candidate prestellar condensations). Their coordinates and *Herschel* fluxes are also given in our tables (#1). We discuss the candidates Class 0 and starless cores in Sect. 6.3.3 and Sect. 6.6 respectively.

5.1. The bright-rim cloud BRC12 and vicinity

BRC12 (or SFO12) is part of the catalogue of bright-rimmed clouds observed in the direction of northern HII regions by Sugitani et al. (1991). It is associated with IRAS 02511+6023, an IR point source of 640 L_{\odot} according to Sugitani et al. (1991;

⁴ The luminosity of the central object is the difference between the total luminosity of the source and the luminosity of the accreting disk, based on the disk accretion rate; the luminosity of the accreting envelope is not taken into account in ROB07 models.

⁵ The mass of the central object is derived from its luminosity and temperature, placing these on evolutionary tracks in an Hertzsprung-Russell diagram. The mass may be wrong if the evolutionary tracks are wrong. The pre-main sequence tracks used, for example those of Siess et al. (2000) for stars in the mass range 0.1 to 7.0 M_{\odot} , are for non accreting stars. Since then, Hosokawa et al. (2009, 2010) have shown that protostars with high accretion rates may occupy different positions in the HR diagram, depending of the accretion rate and of the geometry of the accretion.

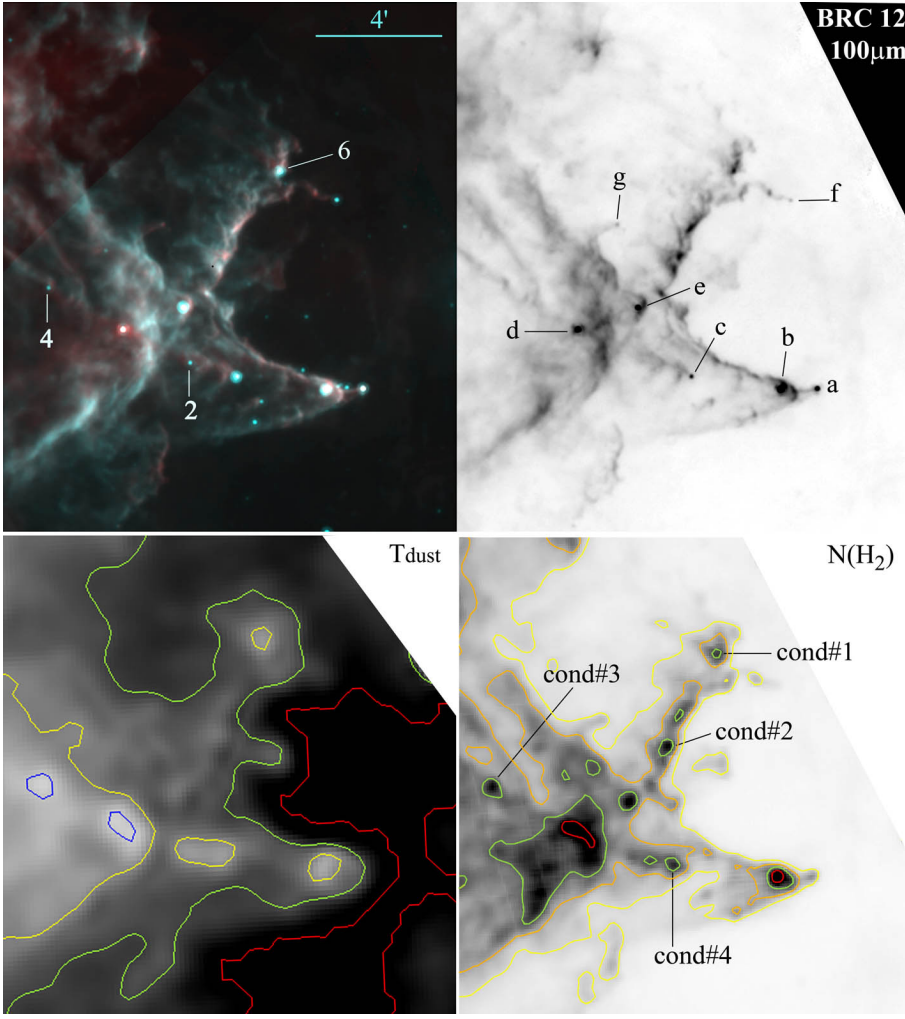


Fig. 5. Field containing the bright-rimmed cloud BRC12. *Top left:* Composite colour image: red is the *Herschel*-PACS emission at 100 μm and turquoise is the *Spitzer*-MIPS emission at 24 μm . *Top right:* The YSOs discussed in the text and present in Table 2 are identified. The underlying grey-scale image is the *Herschel*-PACS image at 100 μm . *Bottom left:* Temperature map of the dust: the red contour is 26 K, the green is 24 K, the yellow is 22 K, and the dark blue is 20 K. *Bottom right:* Column density map: the red contour is $4 \times 10^{21} \text{ cm}^{-2}$, the green is $2 \times 10^{21} \text{ cm}^{-2}$, the orange is $1 \times 10^{21} \text{ cm}^{-2}$, and the yellow is $0.5 \times 10^{21} \text{ cm}^{-2}$. The column density is high mostly in the direction of the point sources (it is the column density of their envelopes, smoothed by the beam). Some other condensations appear, which may be prestellar (see Sect. 6.6). North is up and east is left.

from IRAS fluxes and a distance $d=2$ kpc) or $220 L_{\odot}$ according to Morgan et al. (2008; from SCUBA fluxes and $d=2$ kpc). BRC12 is associated with W5-W and not W5-E. It is part of the bubble surrounding W5-W, and is facing BD+60°0586, an O8III star that contributes to the ionization of W5-W. Morgan et al. (2004) hypothesize that this star is the main source of ionizing photons reaching the rim (it is the closest O star in projection). The NVSS radio-continuum map at 20 cm (Condon et al. 1998) shows strong emission following the bright rim BRC12 and the nearby ionization front (IF). The emission at 20 cm traces the dense ionized boundary layer (IBL) bordering the enclosed molecular material (see fig. 1 in Morgan et al. 2004; online material). From the integrated radio flux of this feature Morgan et al. (2004) calculate an electron density in the IBL of 340 cm^{-3} . BRC12 has been observed at 450 μm and 850 μm with SCUBA by Morgan et al. (2008). A dense core is detected at the tip of the bright rim (their figure A.9). This core corresponds to the source BRC12-b discussed below (Table 2 and Fig. 5).

Fig. 5 shows a colour image of this region with *Herschel* 100 μm emission in red, and *Spitzer* 24 μm emission in turquoise (*Top left*). We have detected seven 100 μm point sources in this region; they are identified by letters in Fig. 5 (*Top right*). Six of them have *Spitzer* counterparts. Table 2 gives their parameters, according to KOE08, and the measured *Herschel* fluxes. Five more Class I and one bright Class II YSO, identified by KOE08

but without 100 μm counterparts, are present in the field covered by Fig. 5.

We fit the SEDs of the YSO candidates in Table 2 using the SED fitting-tool of ROB07. We have assumed a distance of 2 kpc and an interstellar extinction of 1.96 mag (the extinction of HD 18326, that we consider of interstellar origin). We show several SEDs in Fig. 6. In the following we comment on individual sources:

- BRC12-a: this YSO lies at the extreme tip of BRC12, linked to the main BRC by a very thin filament. It is a Stage I YSO with, possibly, a central object of intermediate mass. The disk's mass and accretion rate are not well constrained.

- BRC12-b: this is the brightest mid-IR source in the field, located at the tip of BRC12. It is probably the counterpart of IRAS 02511+6023; it is however less bright than estimated for the IRAS source (due to the low resolution of IRAS, the IRAS fluxes probably include some of the PDR emission). The two possible models reported in Table 3 are somewhat different, but both point to a rather massive Stage I YSO. The dense core observed by Morgan et al. (2008) at 450 μm et 850 μm with SCUBA is co-spatial with this YSO. Morgan et al. (2008) obtain for the temperature of the cold component of the source 21 K, which differs for the temperature of 17.9 K that we estimate from the four *Herschel* fluxes (Table 5). Using their integrated flux at 850 μm , an opacity of $2.0 \text{ cm}^2 \text{ g}^{-1}$ at this wavelength, and a distance of 1.9 kpc they estimate a mass of $18.9 M_{\odot}$ for this source. Their flux, assuming our opacity law (Table 1), and a distance of

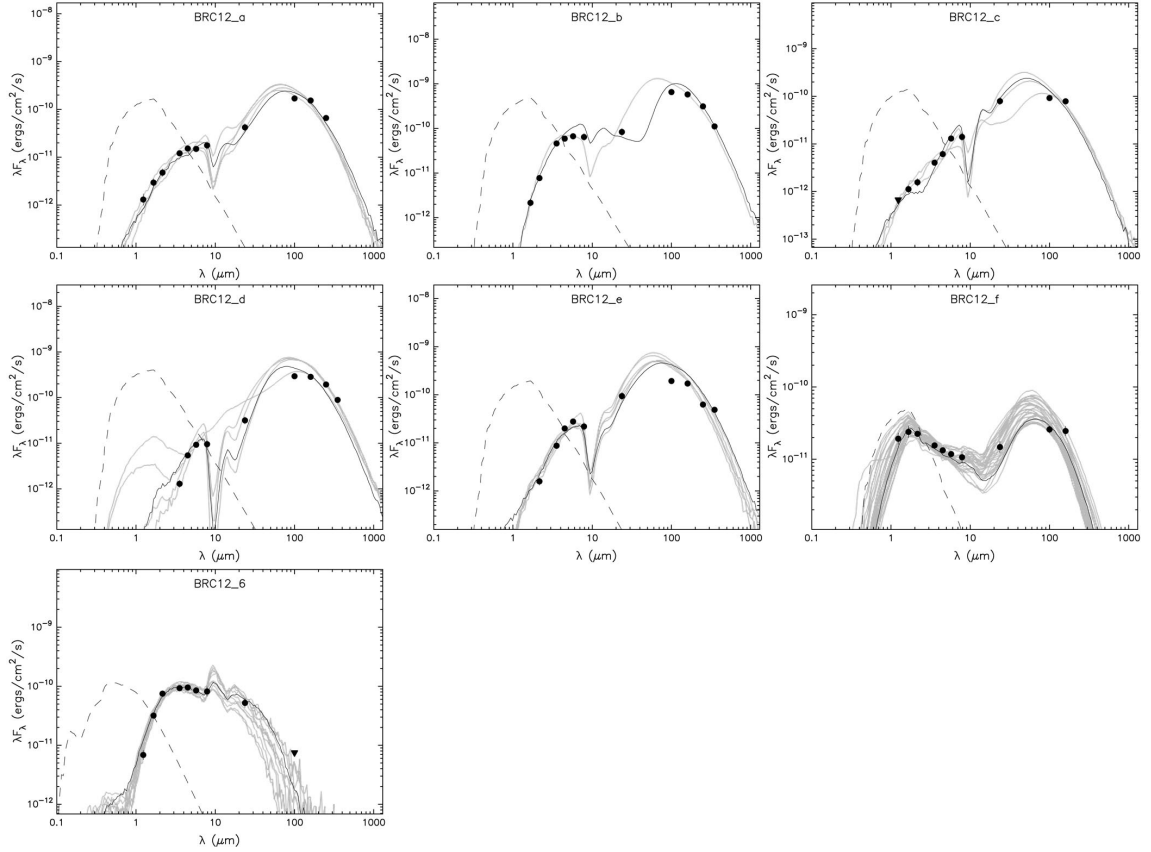


Fig. 6. Spectral energy distributions of point sources in the BRC12 field. We used the SED fitting-tool of ROB07. The filled circles represent the input fluxes and the black line shows the best fit. The grey lines show subsequent fits, all corresponding to $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 . The dashed line shows the stellar photosphere corresponding to the central source of the best fitting model, as it would look in the absence of circumstellar dust, but including interstellar extinction.

2.0 kpc gives a mass of $32 M_{\odot}$, very similar to our estimate (Table 5). Thus, BRC12 harbors a rather massive Stage I YSO, embedded in a high mass envelope with, possibly, a high envelope accretion rate. This suggests that a high mass star is forming there.

- BRC12-c: the three models with $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 that fit this source's SED are rather similar. They all point to a Stage I YSO of intermediate mass.

- BRC12-d: this is also a bright mid-IR source. Five models, somewhat different, fit the SED. All correspond to a Stage I YSO, possibly of intermediate mass; the parameters of the disk are not well constrained.

- BRC12-e: this YSO seems to be at the tip of another bright rim cloud. It resembles BRC12-b, and is a Stage I YSO. But its SED is not well fitted.

- BRC12-f: as for BRC12-a, this source lies at the extremity of a very narrow pillar. It is, however, less bright and less massive than BRC12-a. The SED is not well-constrained, as shown by Fig. 6, especially the disk parameters. The best model corresponds to an uncertain Stage I/II YSO.

- BRC12-g: this source is situated at the extreme tip of a narrow pillar. It has a very faint $24 \mu\text{m}$ counterpart that possibly is extended (its flux was not measured by KOE08).

Five other sources in the field have been classified as Class I YSOs by KOE08 (Table 2). None of these have counterparts in the *Herschel*-PACS image at $100 \mu\text{m}$. Their SEDs are not well-constrained. YSO #1 is observed only in the *Spitzer* bands and lies very close to the bright YSO #6; thus its classification is very

uncertain. YSO #5 lies $\sim 8''$ away from the very bright BRC12-b, and the two sources cannot be separated at $24 \mu\text{m}$ and at longer wavelengths; therefore, its classification is also uncertain. YSOs #2 and #4 lie in the vicinity of condensations #4 and #3 respectively (Fig. 5). Their SEDs are not well constrained; however \dot{M}_{env} is never equal to zero; thus they could be Stage I or Stage I/II YSOs. YSO #6 lies very close to the IF, in the direction of condensation #1; its SED is not well constrained, especially the disk's parameters. But it is clearly a Stage II YSO as its envelope accretion rate is null in all models.

6. Discussion

6.1. The overall morphology of W5-E and the long-distance influence of the ionizing radiation

W5-E appears to have the simple morphology of a large bubble H II region of radius ~ 12 pc (the radius of the circular H α H II region). Its morphology is more complicated though, as shown by Fig. 7 and discussed below:

- In the north two PDRs are present, separated by some 7.5 pc. PDR 1 is ~ 7 pc from the HD 18326 exciting star and PDR 2 is ~ 14.5 pc from it (these are projected distances, measured in the tangential plane). H α and radio-continuum emissions are observed between these two PDRs. This morphology probably results from a projection effect. PDR 1 is seen in absorption in the optical (on the DSS-red or H α images), and thus lies in front of the northern part of the ionized region.

Table 2. Photometry of point sources: the bright-rim cloud BRC12 and its vicinity

Name	#	RA(2000) (°)	Dec(2000) (°)	<i>J</i> mag	<i>H</i> mag	<i>K</i> mag	[3.6] mag	[4.5] mag	[5.8] mag	[8.0] mag	[24] mag	Class	S(100) Jy	S(160) Jy	S(250) Jy	S(350) Jy	S(500) Jy
a	10337	43.718074	60.595320	16.18	14.49	13.22	10.73	9.74	9.01	7.85	3.33	I	5.68	8.24	5.55		
b	10566	43.758246	60.594901		14.83	12.70	9.28	8.27	7.39	6.45	2.59	I	21.86	31.00	26.14	13.03	
c	11055	43.854618	60.602491		15.54	14.43	11.92	10.73	9.16	8.10	2.64	I	3.09	4.22			
d	11604	43.976307	60.628341				13.16	10.86	9.53	8.52	3.64	I	9.87	15.29	16.19	10.42	
e	11311	43.911143	60.639847			14.42	11.09	9.45	8.34	7.62	2.46	I	6.47	9.17	5.25	5.71	
f	10484	43.743111	60.695488	13.26	12.22	11.54	10.46	9.89	9.27	8.40	4.47	II	0.87	1.33			
g													0.98	2.50	2.28		
1	10810	43.806725	60.707375				13.87	13.50	11.58	9.93		I					
2	11275	43.903961	60.610319	15.36	14.43	13.76	12.67	12.04	10.85	9.14	5.01	I					
3	11886	44.053573	60.603121				15.31	14.10	12.87	11.54	7.26	I					
4	11895	44.056178	60.650891	16.55	14.51	13.43	11.98	11.06	10.35	9.42	5.24	I					
5	10592	43.763089	60.594435	16.52	14.82	13.85	12.26	11.49	10.56	9.70		I					
6	10797	43.803910	60.711229	14.38	11.91	10.23	8.52	7.75	7.13	6.19	3.10	II					

Table 3. Characteristics of the YSOs: the bright-rim cloud BRC12 and its vicinity

Name	χ^2	M_* (M_\odot)	T_* (K)	M_{disk} (M_\odot)	\dot{M}_{disk} ($M_\odot \text{ yr}^{-1}$)	\dot{M}_{env} ($M_\odot \text{ yr}^{-1}$)	L (L_\odot)	i (°)	Stage
a	39 (−72)	2.3 (1.9–3.0)	4334	3.3e-2 (4.0e-4–1.3e-1)	3.3e-8 (2.4e-9–1.4e-6)	2.5e-4 (−1.1e-4)	51 (44–77)	41 (31–56)	I
b	58 (−91)	2.9 (6.9)	4175	8.4e-3 (6.4e-3)	3.8e-7 (2.3e-7)	1.3e-3 (4.0e-4)	139 (364)	18 (56)	I
c	49 (−76)	1.6 (−2.0)	4130	3.2e-2 (−7.8e-3)	7.3e-7 (1.2e-8–1.0e-6)	4.8e-5 (3.6e-5–9.1e-5)	43 (36–66)	31 (−63)	I
d	125 (−152)	4.1 (1.6–5.1)	4420	4.5e-2 (1.6e-3–1.5e-1)	5.3e-7 (1.1e-8–4.5e-6)	3.1e-4 (−9.7e-4)	133 (52–183)	69 (−18)	I
e	126 (−159)	2.5 (−5.3)	4329	5.4e-2 (2.5e-3–2.1e-1)	4.3e-6 (−1.8e-10)	4.1e-4 (−1.4e-4)	80 (−184)	31 (−63)	I
f	37 (−70)	1.2 (0.7–4.7)	4228	4.0e-4 (2.0e-4–7.9e-2)	2.5e-9 (−7e-7)	1.0e-5 (3.3e-7–9.1e-5)	14 (10–25)	56 (18–87)	I/II ?
6	13 (−37)	4.8 (−9.1)	16000 (−24000)	4.0e-4 (2.9e-5–1.7e-2)	1.7e-8 (2.3e-11–6.0e-7)	0	400 (−4400)	81 (−87)	II

• In the south the PDRs lie far away from the exciting star: PDR 3 is ~ 16 pc from HD 18326, PDR 4 is ~ 18 pc from it. That the region has expanded more in the south than in the north shows that during the evolution of the H II region its IF has reached regions of lower density in the south than in the north. The evolution of an H II region in a medium with a density gradient has been studied by Franco et al. (1990). If the density gradient is steep (steeper than in $r^{-1.5}$ where r is the distance to the exciting star), the ionization occurs very quickly and the H II region opens to the outside. In this case, no neutral material is collected. This is probably what happened in the southern parts of W5-E. W5-E is mostly ionization-bounded in the east, north, and west, and the ionization radiation escapes freely in the south. The fact that the ionized gas escapes freely in the south probably slows down the expansion of the H II region in the other directions.

• The tips of the bright-rimmed clouds BRC13 and BRC14 are located close to the exciting star (respectively at 7.3 pc and 8.9 pc from it in projection). In these directions the ionizing radiation has reached, somewhere along its path, a relatively high density medium. Also, in these directions, the column density is higher (by a factor 2–3) than in the directions of PDR 1 or PDR 2: more neutral material has been collected in the PDR associated with BRC13 and BRC14. Like PDR 1, BRC 13 and BRC 14 are seen in absorption in the optical; thus, the enclosed condensations lie in front of the ionized region.

• Some structures resembling fingers pointing toward the exciting star are observed at large distances from this star (especially in the north and east). They are observed at 250 μm , 350 μm , and 500 μm . The tips of fingers #1, #2, and #3 (Fig. 7) are respectively at 17 pc, 21.5 pc, and 25 pc from the exciting star (in projection). These fingers point very clearly to this star. This suggests that the main IF has holes through which the ionization radiation escapes, shaping the neutral material far away

from the star. Finger #1 is especially interesting, as it is seen in absorption on the DSS-red image (right insert in Fig. 7); it lies in front of the stellar field, and is related to W5-E as it clearly points to its exciting star. Molecular condensations are always present at the head of these fingers; the masses of these condensations are in the range $1.4 M_\odot - 5.5 M_\odot$ (Sect. 6.6, Table 8). They are susceptible to form, by gravitational collapse later-on, low- or intermediate-mass stars, but not massive stars.

• There is a large quantity of material between W5-E and W5-W. This material is ionized in the A and B H II regions, and is molecular in the long ridge extending north-south, south of the A and B H II regions. This material is possibly compressed by the two expanding W5-E and W5-W H II regions; however its temperature and column density are similar to those of PDR 1 and PDR 2.

6.2. The pillars at the south-east border of W5-E

“Pillars,” also called “elephant trunks”, are small structures located at the interface between an H II region and the adjacent molecular cloud that look like “fingers”. They have a “head”, that is nearest the ionizing star, and a filamentary “tail” that leads from the head away from the star.

The W5-E H II bubble is open in the south, a direction of low density, and many pillars are observed there, especially on the south-east border. These pillars are especially conspicuous at 8.0 μm as shown in Fig. 8. They have counterparts in the *Herschel*-SPIRE images. What we see at 8.0 μm is mainly the emission of PAHs from the “skin” of these dense molecular structures. All such pillars have the same basic morphology, with a rounded head at the tip of a thin filament. Most of these structures point toward the exciting star HD 18326. The column density is the highest in the head (Fig. 8; Table 4) and it decreases

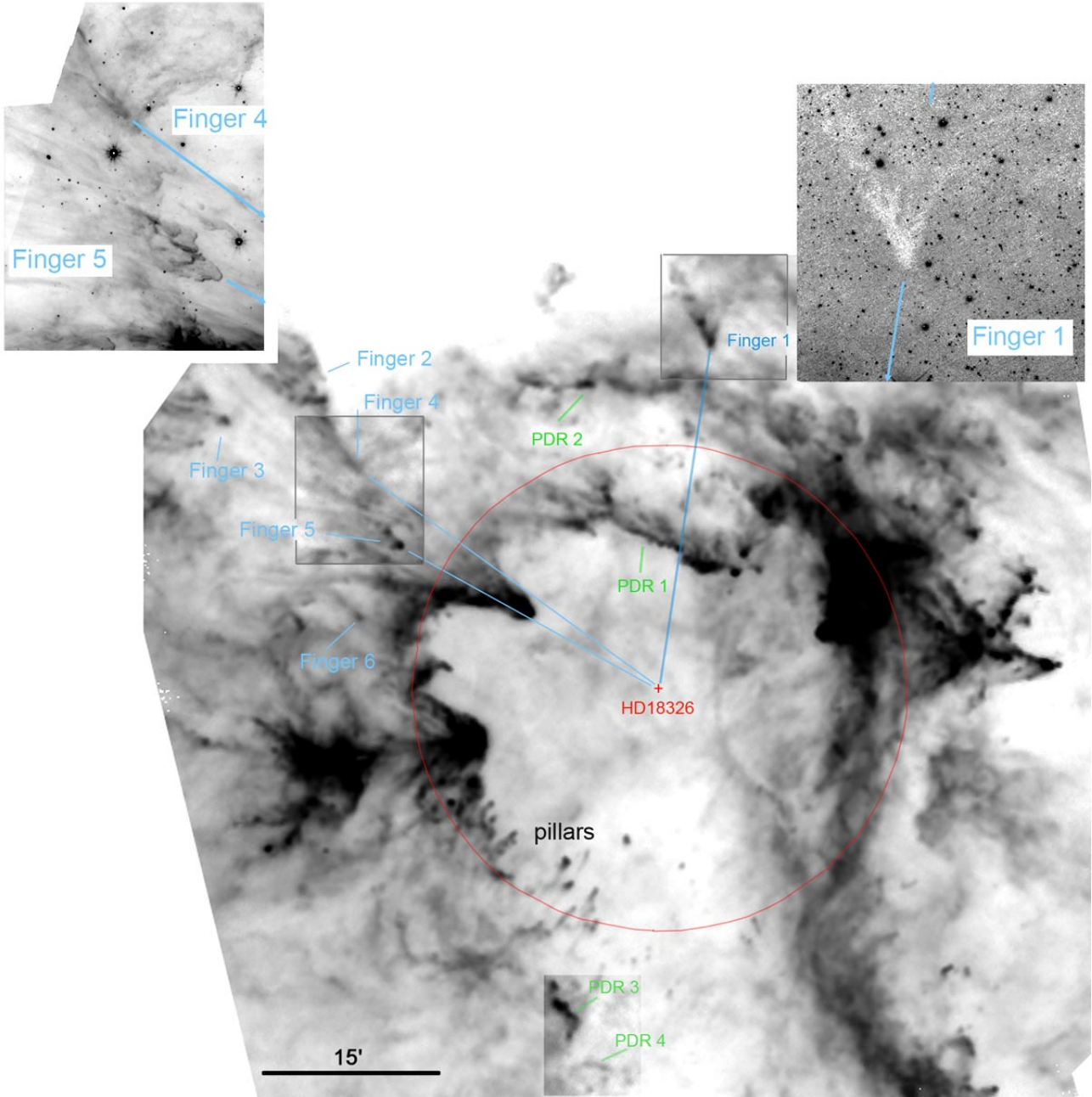


Fig. 7. The morphology of W5-E. Peculiar structures discussed in the text (PDRs, fingers, pillars) are identified on the $350\ \mu\text{m}$ image (logarithmic scale). The red circle shows the approximate extent of the optical H II region (radius 12 pc). The right insert shows finger #1 in absorption, on a DSS-red image. The left insert shows fingers #4 and #5 pointing towards the exciting star, as they appear on the *Spitzer* $8.0\ \mu\text{m}$ image (fingers #2 and #3 lie outside the *Spitzer* field). North is up and east is left.

from the head to the tail. Some heads are separated from the parental molecular cloud, but most of them are not (this is especially evident in the $8.0\ \mu\text{m}$ map). These structures are not fully resolved at $250\ \mu\text{m}$ since the diameters of the heads and the widths of the filaments are of the order of the PSF.

We list some parameters associated with these structures in Table 4. To estimate their mass and column density we need their temperature. We could use the dust temperatures from the temperature map (Sect. 4.1 and Fig. 3), but the size of the pillars is smaller than the resolution of the temperature map (which is the resolution of the $500\ \mu\text{m}$ map), and thus the temperature map values are not necessarily accurate for these structures. We also measured the fluxes of isolated structures

at all *Herschel* wavelengths, using an aperture of radius $45''$ (thus much larger than the beam at $500\ \mu\text{m}$), and fitted their SEDs with a modified blackbody model. (This is not necessarily better than the temperature map, as what we get is an average temperature within the aperture.) The first three columns of Table 4 give the identification of these structures according to Fig. 8 and the coordinates of the condensations at the heads of the pillars. Column four gives the dust temperature obtained from the temperature map. These temperatures are very similar: for the 15 structures in Table 8 we obtain $T_{\text{dust}} = 22.7 \pm 0.5\ \text{K}$. This temperature is characteristic of the PDRs in W5 (Fig. 3). The temperatures obtained by fitting the SEDs (aperture photometry) are also given in column four (with an asterisk). Again these

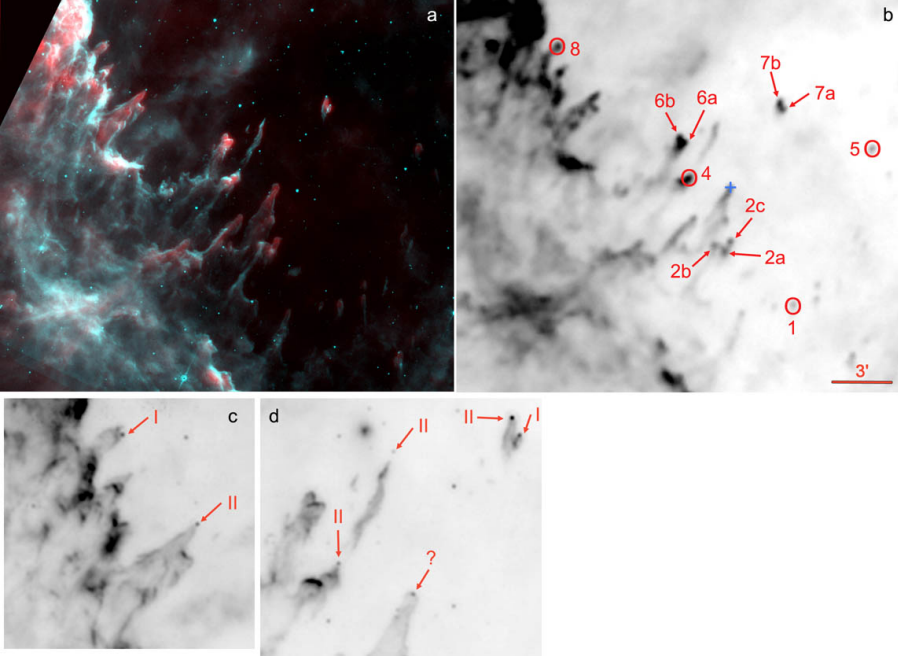


Fig. 8. The IR emission from the “pillars” located south-east of W5-E. *a*: Colour composite image showing in red the 250 μm cold dust emission and in turquoise the 8.0 μm PAH emission. *b*: Identification of a few pillars on the grey-scale *Herschel* 250 μm image. *c*, *d*: Details of the pillars: the grey-scale underlying *Spitzer* 24 μm image shows Class I/II sources lying at the tip of the pillars (from KOE08). North is up and east is left.

temperatures ($T_{\text{dust}} = 21.5 \pm 1.9 \text{ K}$) are compatible with the temperature of a PDR. In the following, we use a mean temperature of 22 K to estimate the mass and column density of all these structures. Column five gives the flux density at 250 μm of the heads of the pillars, calculated by integrating over an aperture of radius 22.5'' (much larger than the beam at 250 μm). Columns six and seven give the mass of the heads and their column density at the peak emission. The masses are all small, in the range $0.3 M_{\odot}$ – $1.4 M_{\odot}$ (except for pillar #10). These masses are uncertain mostly because of the uncertainty on the opacity. We find, however, that no massive stars can form at the head of these pillars. These structures are not well-resolved at 250 μm but we may estimate their sizes using the 8.0 μm emission, which traces the ionization fronts bordering these structures. For example, the column density in the direction $\alpha(2000) = 03^{\text{h}}00^{\text{m}}08.46^{\text{s}}$ $\delta(2000) = 60^{\circ}17'02''$ (the head of the pillar marked with a blue cross in Fig. 8b) is $7.15 \times 10^{20} \text{ cm}^{-2}$. The diameter of the head given by the 8.0 μm image is $\sim 0.145 \text{ pc}$, and thus we find a mean density $\sim 1600 \text{ molecules cm}^{-3}$ in the head of this standard pillar (estimated assuming a uniform density and temperature).

All the pillars at the south-east border of W5-E point to the exciting star HD 18326. A group of pillars on the east side of W5-E between BRC13 and BRC14 (including pillars #9, #10, and #11 in Table 4 and Finger #6 in Table 8), however, seem to point towards a different direction. They point to a zone where extended 24 μm emission is observed (see Fig. 9). There is another massive star V1018 Cas in this direction (green cross in Fig. 9). This star is an eclipsing binary (Bulut & Demircan 2007). Its spectral type is O7V (Otero et al. 2005) or B1 (Reed 2003). This star is possibly responsible for the surrounding extended 24 μm emission but its radial velocity ($\sim 10 \text{ km s}^{-1}$, Duflo et al. 1995) indicates that it likely does not belong to the W5 complex (the ionized gas has a velocity in the range -34 to -39 km s^{-1}). This situation is difficult to understand, and the velocity of V1018 Cas needs to be confirmed.

What is the origin of the pillars? There are several models that attempt to explain these structures. All these models are

Table 4. Physical parameters of the pillars

Name	RA(2000)	Dec(2000)	T_{dust} (K)	$S_{\nu}(250\mu\text{m})$ (Jy)	M (M_{\odot})	$N(\text{H}_2)$ (cm^{-2})
1	02:59:42.18	+60:11:09	23.7	0.50	0.33	3.7×10^{20}
1*			20.8*			
2a	03:00:10.34	+60:13:53	22.8	1.75	1.16	6.9×10^{20}
2b	03:00:14.58	+60:14:10	22.95	0.89	0.59	6.1×10^{20}
2c	03:00:07.93	+60:14:21	22.75	1.25	0.83	6.7×10^{20}
3	02:58:21.41	+60:14:32	22.3	1.46	0.97	7.0×10^{20}
4	03:00:24.93	+60:17:35	22.6	2.25	1.49	1.2×10^{20}
4*			24.3*			
5	02:59:09.89	+60:19:07	23.5	0.71	0.47	4.5×10^{20}
5*			19.1*			
6a	03:00:26.96	+60:19:24	22.0	2.08	1.38	1.5×10^{20}
6b	03:00:28.86	+60:19:36	22.0	2.08	1.38	1.5×10^{20}
6a*+6b*			21.2*			
7a	02:59:46.24	+60:21:08	22.75	1.45	0.96	8.8×10^{20}
7b	02:59:47.53	+60:21:27	22.75	1.64	1.09	9.6×10^{20}
7a*+7b*			22.3*			
8	03:01:19.28	+60:24:13	22.5	1.72	1.14	1.1×10^{20}
9	03:01:48.07	+60:35:31	22.6	3.71	2.46	1.5×10^{21}
10	03:01:49.13	+60:40:01	23.0	1.45	0.96	7.6×10^{20}
11	03:01:58.76	+60:38:29	23.2	0.66	0.44	4.9×10^{20}

based on the influence of the UV radiation field of the central exciting star upon an inhomogeneous surrounding medium. They differ in the origin of the inhomogeneities. In the well-known radiation-driven implosion models (RDI), pre-existing dense condensations are present. The RDI models describe the compression of a spherical isothermal isolated globule (Bertoldi 1989; Lefloch & Lazareff 1994; Kessel-Deynet & Burkert 2003; Miao et al. 2009; Bisbas et al. 2009, 2011; Haworth & Harries 2011; Tremblin et al. 2011). In another model, the inhomogeneities are due to turbulence (Gritschneder et al. 2009a, 2010). Both models are able to produce structures similar to the observed pillars with a dense core (the head) at the tip of a filament. As discussed in Gritschneder et al. (2010) the velocity field in principle allows us to distinguish one process from the other. Unfortunately, we do not have the velocity field to discriminate easily between the two processes.

Both models predict star formation in the pillars, under specific conditions. In Gritschneder et al.’s simulations, gravitational collapse occurs at the tip of the pillars, leading to the for-

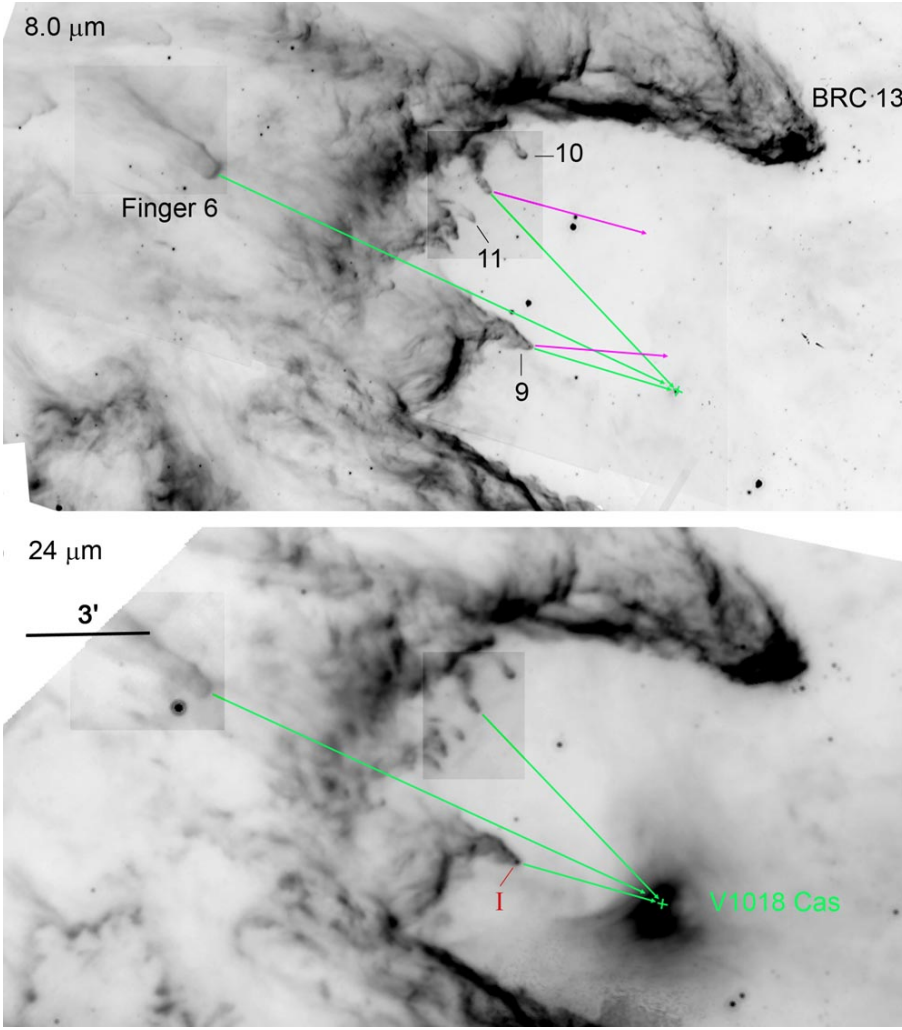


Fig. 9. Details of the pillars situated between BRC13 and BRC14. These pillars seem to point towards the massive star V1018 Cas. *Top:* The pillars are identified on the *Spitzer* 8.0 μm image, according to their number in Tables 4 and Table 8. The green arrows indicate the direction of V1018 Cas (green cross), the pink arrows that of HD 18326. *Bottom:* The *Spitzer* 24 μm image shows the bright extended emission possibly associated with V1018 Cas. A Class I YSO has been detected at the head of pillar #9 by KOE08. North is up and east is left.

mation of cores and then stars. This occurs after a few 10^5 yr, depending of the initial conditions of the simulations (mean density and Mach number). The formed cores are not very massive, with masses $\leq 1 M_{\odot}$. In the simulations of Bisbas et al. (2011), the critical parameter is the incident ionizing flux reaching the globule Φ_{LyC} . Star formation occurs for Φ_{LyC} over a particular range. We estimate that $\Phi_{\text{LyC}} \sim 8 \times 10^8 \text{ cm}^{-2} \text{ s}^{-1}$ at the head of pillar #4 or at the position of the blue cross in Fig. 8; this value allows star formation according to Bisbas et al. (2011); in these conditions the first star forms rather quickly (in less than 0.1 Myr) once the pre-existing globule is reached by the ionizing photons, and the new stars have low masses. YSOs are indeed present at the tip of some pillars, as shown by Fig. 8 (see also Fig. 1 in Chauhan et al. 2011a). According to KOE08 they are Class I or Class II YSOs. Chauhan et al. (2011b) have used V versus $V - I$ diagrams and YSO evolutionary tracks to estimate the age of a few sources present at the tip of the pillars. They estimate that they are pre-main sequence low-mass stars having ages of one Myr or less. All this is consistent with all the simulations.

The RDI simulations of Haworth & Harries (2011) suggest a somewhat different origin for the pillars: they could be due to small-scale instabilities of the IF at the periphery of a dense globule, and not to the compression of the globule. The ionizing photon flux is also a critical parameter in their simulations. The ionizing photon flux is low at the position of the pillars (position of the blue cross, Fig. 8). As shown by the simulations this configuration exhibits the largest susceptibility to develop insta-

bilities (but the simulations also show that the resulting pillars are smoothed out by the diffuse field radiation). Chauhan et al. (2011b) favour this explanation for the origin of the pillars.

That there are numerous pillars observed at the south-west border of W5-E, a region of low ionizing flux, favours the explanation put forth by Gritschneider et al. (2010) involving turbulence, or the explanation of Haworth & Harries (2011) involving instabilities; however, following the model of Haworth & Harries, we would expect the formation of pillars at the periphery of dense condensations; these are not observed with *Herschel*. Thus, the formation of the pillars, the cores at their heads, and the subsequent low-mass YSOs are likely triggered by the W5-E H II region, resulting from the interaction of the UV radiation of its exciting star with a low density inhomogeneous medium, probably of turbulent origin.

6.3. The properties of the young stellar objects

We detect 50 100 μm point-like sources in the field covered by *Herschel*. The number of detected sources decreases at longer wavelengths (respectively 43, 37, 23, and 11 at 160 μm , 250 μm , 350 μm , and 500 μm). Due to the temperature of their envelopes (Sect. 6.3.1) their fluxes are maximum in the 100 μm –160 μm range, and decrease at longer wavelengths; also, due to the decreasing angular resolution, it becomes more and more difficult to separate point sources from their bright background.

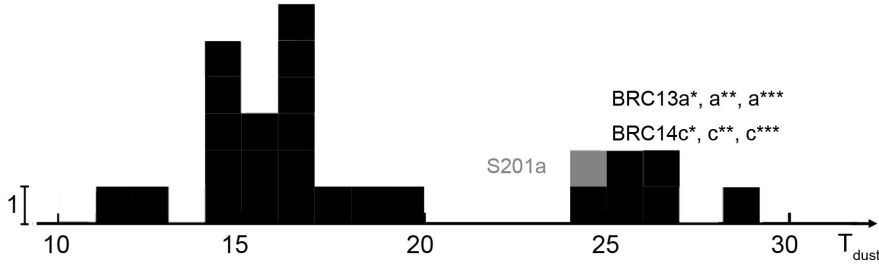


Fig. 10. The dust temperature of YSO envelopes or of compact condensations, estimated by fitting the *Herschel* fluxes using a modified blackbody model.

6.3.1. Temperatures and masses of the envelopes

The *Herschel* emission from YSOs comes mostly from their envelopes. We can use their fluxes and a modified blackbody model (see Sect. 4) to estimate the envelope temperatures and masses. We use the fluxes given in Tables 2, A.1, B.1, C.1, D.1, and E.1. We apply a colour correction, which takes into account the varying spectral response with dust temperature (see Sect. 4.1 in AND12). We compute the masses from the 350 μm fluxes, using a spectral index $\beta=2$ and the temperatures previously estimated. We list the temperatures and masses in Table 5, in columns 3 and 4 respectively. This Table contains only the sources that are detected in at least four *Herschel* bands. The first column gives the identification of the sources, the second one the number of flux measurements (n equal to 4 or 5).

Fig. 10 displays the temperatures of the detected sources. The histogram shows two maxima: one around 15.7 K, containing the majority of the point sources, and one around 25.8 K. Thus we find that the envelopes of YSOs are generally cold, with a characteristic temperature of 15.7 ± 1.8 K. The second peak contains extended sources associated with clusters: the extended BRC13-a source around the luminous BRC13-a Class II YSO (three apertures, the smallest one is the hottest), the extended BRC14-c1+c2 sources associated with IRS1 and IRS2 in AFGL 4029 (again three apertures and the smallest is the hottest), and the YSO S201-a. The high temperature of S201-a is also apparent in the temperature map in Fig. 3 and therefore this source differs from the other YSOs that have cold envelopes. S201-a seems to be extended at 8.0 μm , and unresolved at the other *Spitzer* wavelengths; it is possibly surrounded by PAH emission. Its SED points to a luminous Class I YSO, observed with the disk face-on (through the cavity, at least for the best model). As mentioned in Sect. 6.4, this YSO is probably not associated with W5-E, and instead lies in the far background; it displays a massive envelope.

Table 5 allows us to compare the masses of the envelopes estimated with two different methods. Column 4 gives the masses obtained using the modified blackbody model and the flux at 350 μm . Column 5 gives the masses estimated using the SED fitting-tool of ROB07, fit using the *Herschel* fluxes alone. We also give in column 5 the range of masses corresponding to the ten best models. Fig. 11 shows that the masses obtained using the SED fitting-tool of ROB07 are always much larger than the masses estimated with the modified blackbody model. This results mainly from the different opacities used in the two models. At 350 μm , ROB07 use an opacity of $2 \text{ cm}^2 \text{ g}^{-1}$ whereas we use an opacity of $7.3 \text{ cm}^2 \text{ g}^{-1}$. The masses derived with the two methods therefore should be in the ratio $7.3/2=3.65$ (the red line in Fig. 11). The discrepancy however is even greater: for the 25 regions in Table 10 we obtain a mass ratio of 5.58 ± 2.54 . In Fig. 11, we show the masses estimated using fluxes in five *Herschel* bands in colour (yellow and green), whereas the masses

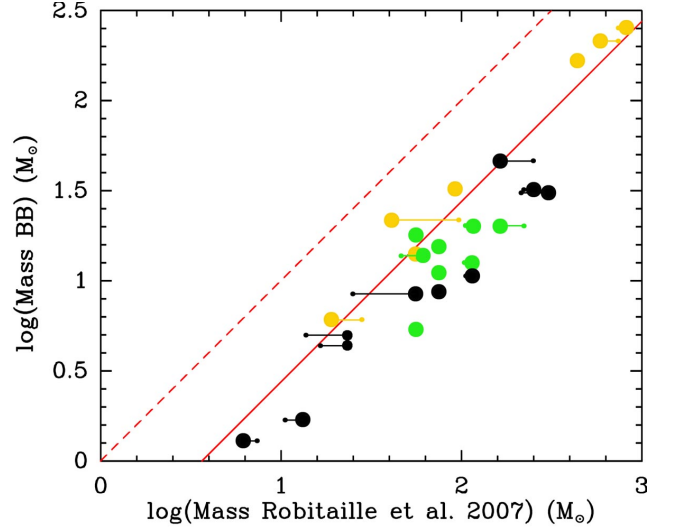


Fig. 11. The masses of the YSO envelopes or of compact condensations estimated from the *Herschel* fluxes. Here we use only sources that have four or five measured *Herschel* fluxes. We estimate the masses $M(\text{BB})$ using a modified blackbody model. We estimate the masses $M(\text{Robitaille et al.})$ using ROB07 SED fitting-tool. The red dashed line shows equality between the two masses and we find no sources located along it. The red solid line shows where the points should lie, taking into account the different dust opacities used in the mass estimations. The black and coloured dots are respectively for sources with 4 and 5 measured fluxes. Yellow and green dots are respectively for hot ($T \sim 25.8$ K in Fig. 10) and cold ($T \sim 15.7$ K in Fig. 10) sources. The horizontal bars indicate the range of values obtained using ROB07's SED fitting-tool.

estimated using only four *Herschel* fluxes are in black. The discrepancy is worse for the sources with only 4 flux measurements (ratio equal to 6.60 ± 2.26 and 4.78 ± 2.53 for respectively 11 sources with $n = 4$ and 14 sources with $n = 5$); but the main difference is between sources surrounded by hot dust (second maxima in the histogram of Fig. 10; yellow dots) or cold dust (first peak in the histogram of Fig. 10; green and black dots): ratio equal to 2.96 ± 0.79 and 6.59 ± 2.23 for respectively 7 sources surrounded by hot dust or 18 sources with cold envelopes. The largest discrepancies are for the cold envelopes.

The transfer models presently used in the SED-fitting tool of ROB07 do not include cold dust (the external radii of the envelopes correspond to a dust temperature of 30 K; T. Robitaille, private communication). To account for the *Herschel* measured fluxes - due to cold dust, the SED-fitting tool overestimates the mass of the envelope. Additionally, the emission peak of the envelope is shifted toward shorter wavelengths (as the SED-fitting tool overestimates the temperature of the envelope). These two

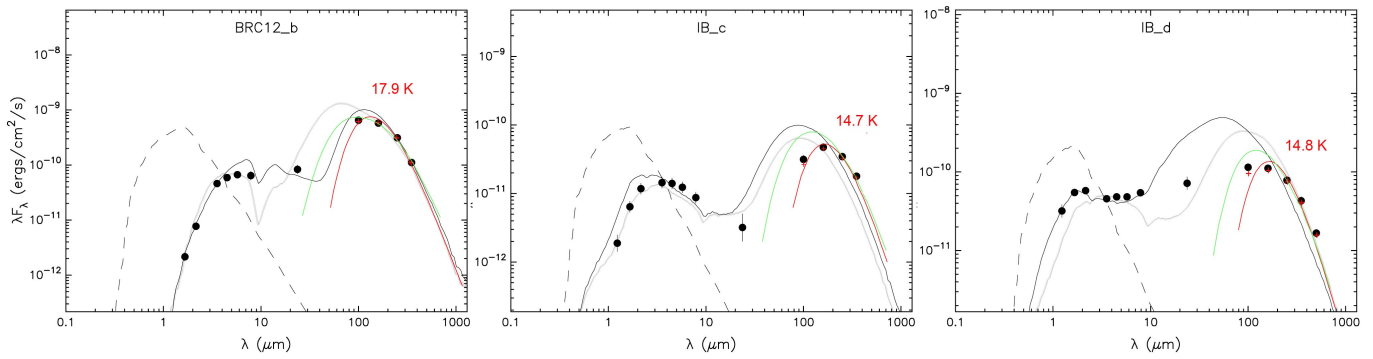


Fig. 12. Fitting the SEDs of the envelopes of three YSOs. The black curves show the fits of the whole SEDs, obtained using the SED fitting tool of ROB07; the green curves show the fits obtained for the envelopes, using the same models, but considering only the *Herschel* fluxes. The red curves are the fits obtained using the blackbody modified model.

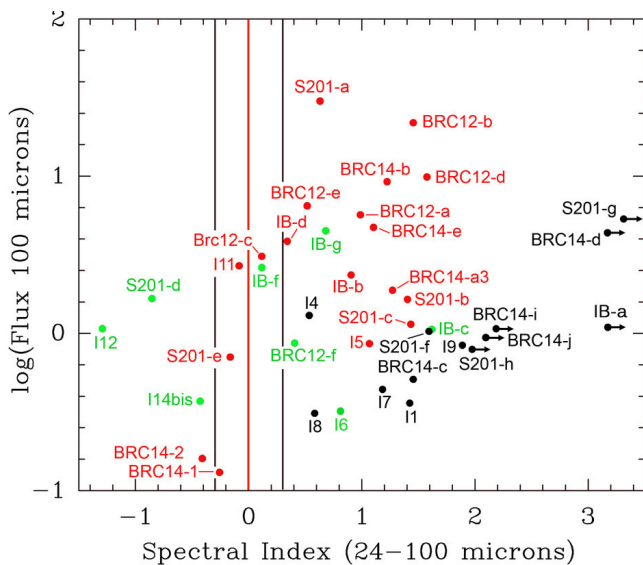


Fig. 13. The integrated flux at $100\,\mu\text{m}$ of YSOs versus their spectral index between $24\,\mu\text{m}$ and $100\,\mu\text{m}$. The sources classified as Class I and Class II YSOs by KOE08 are respectively represented by red and green dots. The black dots show sources with no counterparts in KOE08.

effects (overestimation of the envelope mass and temperature) are conspicuous in Fig. 12 which shows the fits obtained by the two methods. Obviously, the modified blackbody model (red curves) fits the *Herschel* data points better than ROB07's models (green curves). New dust radiative transfer models are in preparation, including PAH emission, heating by external radiation, and larger envelope radii (Sewilo et al. 2010). The first tests of these models are discussed by these authors. Presently it is probably better to use the modified blackbody model to estimate the characteristics of the YSOs' envelopes.

The masses of the envelopes detected by *Herschel* around 100 μm point sources, estimated using the modified blackbody model, lie in the range $1.3\ M_{\odot} - 47\ M_{\odot}$.

6.3.2. Evolutionary stages - Comparison with Koenig et al.'s (2008) classification

The field covered by *Herschel* contains 50 point sources detected at 100 μm (we are considering here the point sources and not

Table 5. Temperatures and masses of the envelopes of YSOs and compact molecular condensations

Identification	n	$T_d(\text{BB})$ (K)	$M_{\text{env}}(\text{BB})$ (M_{\odot})	$M_{\text{env}}(\text{Rob.})$ (M_{\odot})
BRC12-b	4	17.9	30.9	306.0 (−215)
BRC12-d	4	16.3	32.1	251.0 (−223)
BRC12-e	4	18.7	12.3	115.0 (−106)
BRC13-a*	5	28.3	6.1	19.2 (−28)
BRC13-a**	5	26.0	13.8	61.5 (−46)
BRC13-a***	5	24.8	21.8	41.2 (−97)
BRC14-a3	5	14.4	15.6	75.4 (−79)
BRC14-b	5	16.7	20.2	165.0 (223)
BRC14-c	5	11.8	22.5	
BRC14-c*	5	26.4	166.6	440.0 (−460)
BRC14-c**	5	25.5	214.7	591 (−741)
BRC14-c***	5	25.0	254.9	823 (−741)
S201-a ^a	5	24.9	32.6	92.8
S201-b	4	19.4	1.3	6.2 (−7.4)
S201-c	5	12.9	18.0	56.0
S201-d	5	15.9	5.4	56.0
S201-g	4	14.5	46.5	165 (−251)
S201-h	4	16.1	4.4	23.4 (−16.6)
IB-a	4	14.1	14.1	56.0 (−54.7)
IB-b	4	16.4	8.7	75.4 (−54.7)
IB-c	4	14.7	8.5	56.0 (−25.1)
IB-d	5	14.8	20.3	115.2 (−106.0)
IB-e	5	16.0	10.7	115.2 (−106.0)
Isolated-5	4	16.6	1.7	13.6 (−10.6)
Isolated-11	5	15.5	11.1	75.4 (−78.7)
Isolated-13	4	15.6	5.0	23.4 (−13.8)

Notes. ^(a) For a distance of 5.45 kpc (Sect. 6.5).

the condensations). Twenty eight of these 100 μm point sources (56%) have a counterpart in the KOE08 catalogue: 19 are Class I sources, and 9 are Class II; the remaining eleven sources are candidate Class 0 sources. We fit the SEDs of these sources using the SED fitting-tool of ROB07 and 2MASS, *Spitzer*, and *Herschel* fluxes. Among the 19 Class I YSOs of KOE08, we confirm that 16 are Stage I sources (84% of the cases); among the 9 Class II YSOs we confirm with certainty that 3 are in Stage II (30% of the cases). Thus the agreement with KOE08 is rather good, especially for the Class I sources that, due to their cold envelopes, we expect *Herschel* to detect. We comment on these results below.

Fig 13 and Fig. 14 show the flux versus spectral index diagrams for these sources (the equivalent of the more classical

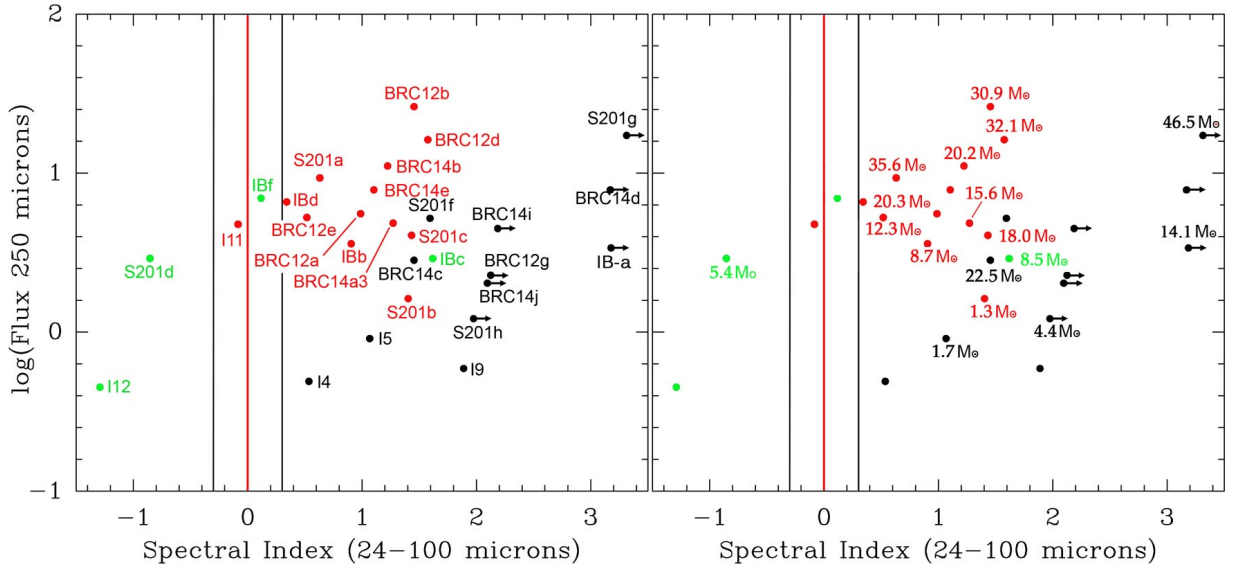


Fig. 14. The integrated flux at 250 μm of YSOs versus their spectral index between 24 μm and 100 μm . The symbols are the same symbols as in Fig. 13. *Left*: identification of the YSOs. *Right*: the masses of the envelopes are indicated, estimated using the modified blackbody model (we require at least four *Herschel* fluxes for a mass estimate).

magnitude – colour diagrams) using *Herschel* measurements. We estimate the spectral index between 24 μm and 100 μm . The spectral index indicates whether the flux is increasing between these wavelengths (a possible signature of the presence of an envelope). We have used colours to indicate the classification of KOE08, red for their Class I and green for their Class II sources. The black symbols correspond to sources that are not in KOE08 catalogue.

Fig. 13 shows that most Class I YSOs in KOE08 have a positive spectral index between 24 μm and 100 μm ; however a few have a negative one. KOE08 Class II YSOs have positive as well as negative spectral indices. One explanation for this large range of spectral indices may be, as Robitaille et al. (2006) demonstrate, the influence of the inclination angle (between the line of sight and the plane of the disk) upon the shape of the SED. Their fig. 7 shows that: i) Stage 0/I YSOs, dominated by their envelopes, have a positive spectral index if the central source is of intermediate or low mass and a negative index if the central source is massive; ii) Stage II YSOs, dominated by their disks, have a spectral index that depends of the inclination angle. For example, Stage II YSOs with a central object of intermediate mass, seen edge-on, have a double-peaked SED and a positive spectral index; and again if the central object is massive, the spectral index is negative. Several of the Class II YSOs of KOE08 that have a positive spectral index present a double-peaked SED (BRC12-f, IB-c, IB-f, IB-g). It is difficult to classify these sources, they are possibly intermediate mass Class II YSOs viewed edge-on. We do not confirm the classification of BRC14-1 and -2 as Class I YSOs, as proposed by KOE08. The classification of S201-e, I-11, and I-6, which have flat SEDs, is also uncertain. We conclude that the spectral index is not a great indicator of the stage of the YSO. Sources with a highly positive spectral index, however, are Class 0/I sources (or Stage 0/I); among them, as is discussed in Sect. 6.3.3, BRC14-c, BRC14-d, BRC14-i, BRC14-j, S201-g, S201-h, and IB-a are candidate Class 0 YSOs; they have no 24 μm counterparts or very faint ones.

Fig. 14 uses the flux at 250 μm in ordinate and the same spectral index from Fig. 13 in abscissa; the flux at 250 μm can be

considered as an indicator of the mass of the envelope. Indeed Fig. 14 (*Right*) shows a decrease in envelope masses from top ($\sim 47 M_{\odot}$) to bottom ($\sim 1 M_{\odot}$). The mass depends on the dust temperature of the envelope, which is sometimes difficult to estimate accurately (see the dispersion of the temperatures in Fig. 10). This explains why a few nearby sources in this diagram can have rather different masses. (This is the case for IB-c and BRC14-c. BRC14-c has the coldest envelope temperature of the whole sample; its temperature is possibly underestimated and thus its mass overestimated). Fig. 14 shows that several YSOs have an accreting envelope mass higher than 20 M_{\odot} ; these YSOs will possibly form massive stars (for example, BRC12-b, with an envelope accretion rate of $\sim 1 \times 10^{-3} M_{\odot} \text{ yr}^{-1}$ or BRC12-d, with an accretion rate of $\sim 3 \times 10^{-4} M_{\odot} \text{ yr}^{-1}$, or BRC14-a3, BRC14-b, BRC14-c, BRC14-e, or S201-c, which have accretion rates $\geq 4 \times 10^{-4} M_{\odot} \text{ yr}^{-1}$; we again caution that the envelope accretion rate, which is estimated using the ROB07 SED fitting-tool, is very uncertain). Among the candidate Class 0 sources, S201-g has a high mass envelope; it is probably also the case of BRC14-d (15.6 M_{\odot} assuming a temperature of 15.7 K for the envelope).

The field covered by *Herschel* contains 55 Class I sources according to KOE08. Only 19 of these have a *Herschel* counterpart at 100 μm , 16 of which are confirmed as Stage I sources. Thus, most KOE08 Class I sources (65%) are not detected by *Herschel*. Most of these sources are also faint at 24 μm ; we cannot say anything new about them. A few of the KOE08 Class I sources, however, are bright at 24 μm ; their non-detection at 100 μm makes their identification as Class I sources doubtful (see for example the case of the 3 bright sources in the Sh 201 region, Appendix C). Also, a few KOE08 Class I sources belong to small stellar groups (for example in the BRC14 field, Appendix B), or lie close to another bright YSO (for example in the BRC12 field, Sect. 5.1); the resolution of *Herschel* does not allow us to separate these sources.

Several *Spitzer* sources (KOE08 Class I and Class II) are sometimes observed in the direction of *Herschel* point sources. Sometimes the resolution at 100 μm is sufficient to associate one *Spitzer* source with the *Herschel* source (for example BRC14-a3, Appendix B), but not always. In these cases *Herschel* may be

detecting one compact condensation inside which a small stellar group is forming (for example BRC14-g+h, Appendix B, or IB-e, Appendix D). Higher resolution observations are needed to better understand these situations.

6.3.3. Candidate Class 0 sources

In the preceding sections we have identified Class I or Stage I sources as those that are dominated by their envelopes. We would like now to identify the youngest of them, the Class 0 sources. The lifetime of these YSOs is $\leq 10^5$ yr, according to Maury et al. (2011; assuming that the lifetime of Class I YSOs is a few 10^5 yr). (These lifetimes are still very uncertain and depend possibly of the final mass of the central source.) Different indicators can be used to identify Class 0 sources: i) Class 0 sources have M_{env}/M_* larger than 1 (André et al. 1993 and references therein; M_* is the mass of the central object); however if the source is not observed at near- and mid-IR wavelengths it is difficult to estimate M_* . It has been shown, in nearby star forming regions, that the envelope of Class 0 sources is extended (FWHM in the range $10''$ – $14''$ for two Class 0 in the ρ Oph region at 160 pc, André & Montmerle 1994); however the resolution of *Herschel*-SPIRE does not allow us to resolve these sources in W5-E at 2 kpc; ii) for sources observed at sub-mm or mm wavelengths another condition is $L_{\text{bol}}/L_{\lambda \geq 1.3\text{mm}} \leq 2 \times 10^4$, or $L_{\lambda \geq 350\mu\text{m}}/L_{\text{bol}} \geq 10^{-2}$ (André et al. 1993); a roughly equivalent criterium is $M_{\text{env}}/L_{\text{bol}} \geq 0.1 M_{\odot}/L_{\odot}$ (André et al. 2000). One can also use evolutionary diagrams like M_{env} versus L_{bol} where it has been shown that Class 0 and Class I sources occupy different locations (Motte & André 2001; André et al. 2008; Hennemann et al. 2010); iii) some authors use T_{bol} , the bolometric temperature (which is the temperature of a blackbody having the same mean frequency as the observed YSO SED) to estimate the importance of the envelope. Enoch et al. (2009 and references therein) define Class 0, and Class I sources as sources with respectively $T_{\text{bol}} < 70$ K, and $70 \text{ K} \leq T_{\text{bol}} \leq 650$ K.

Table 6 gives these different indicators estimated for the *Herschel* point sources, when possible. The first column is the identification of the source and the second and third columns are the mass of the envelope using the modified blackbody fit (as in Table 5) and the mass of the central object estimated using the SED fitting-tool of ROB07. The resulting ratio M_{env}/M_* is given in column 4. Columns 5 and 8 give respectively the bolometric luminosity and temperature (according to Enoch et al. 2009), obtained by a simple integration of the observed SED:

$$L_{\text{bol}} = 4\pi d^2 \int S_{\nu} d\nu \quad (5)$$

and

$$T_{\text{bol}} = 1.25 \times 10^{-11} \frac{\int \nu S_{\nu} d\nu}{\int S_{\nu} d\nu} \quad (6)$$

Columns 6 and 7 give respectively the ratios $M_{\text{env}}/L_{\text{bol}}$ and $L_{\lambda \geq 350\mu\text{m}}/L_{\text{bol}}$. Column 9 gives the spectral index $\alpha(24 \mu\text{m} - 100 \mu\text{m})$. In column 10 we list the presence or absence of a near- or mid-IR component and of a CO outflow (see Sect. 6.4).

Table 6 shows that most of these indicators are in rather good agreement, pointing to the same sources for the candidate Class 0 objects. We think that the most reliable indicators are these estimated directly from the observations, such as $M_{\text{env}}/L_{\text{bol}}$ or T_{bol} . Good candidate Class 0 sources are BRC14-c and S201-c which have $M_{\text{env}}/L_{\text{bol}} \geq 1$ and $T_{\text{bol}} \leq 70$ K; S201-g, S201-h and IB-a, which have $M_{\text{env}}/L_{\text{bol}} \geq 1$ and a large value of

$\alpha(24 \mu\text{m} - 100 \mu\text{m})$, are also good candidate Class 0 sources. All of these candidate Class 0 sources display high values for the ratio $L_{\lambda \geq 350\mu\text{m}}/L_{\text{bol}}$. BRC14-a3 appears to be at the limit between Class 0 and Class I. And all the indicators confirm that S201-d is a Class II YSO. The indicator M_{env}/M_* is less reliable as a large range of values are often proposed for M_* by the SED fitting-tool of ROB07. (Table 6 shows that all sources, except S201-d, should be Class 0 sources according to this indicator; see also our reservation concerning M_* in Sect. 5.) Enoch et al. (2009) find that the spectral index $\alpha(2 \mu\text{m} - 10 \mu\text{m})$ shows a large dispersion for the Class I sources; we observe the same trend for the index $\alpha(24 \mu\text{m} - 100 \mu\text{m})$. However this index is high, ≥ 1.2 , for all our candidate Class 0 sources.

Several other sources for which we have less measurements are also possibly Class 0 sources. They are BRC14-d, BRC14-i, and BRC14-j, which have no *Spitzer* counterpart (even at $24 \mu\text{m}$) and thus present a very high spectral index (≥ 2); their masses, estimated from their $250 \mu\text{m}$ integrated flux assuming a cold envelope of temperature 15.7 K are respectively $16.8 M_{\odot}$, $10.6 M_{\odot}$, and $11.1 M_{\odot}$. Also the two bright $100 \mu\text{m}$ sources s1 and s2 at the waist of Sh 201 could be Class 0 sources.

Most of these candidate Class 0 sources have massive envelopes and, by accretion, may become massive stars (S201-h is an exception).

S201-a, a distant background source (Sect. 6.4; distance of 5.45 kpc), appears as a Class I object, with $T_{\text{bol}}=113$ K, $M_{\text{env}}=32.6 M_{\odot}$, $L_{\text{bol}}=2460 L_{\odot}$. Its high luminosity suggests that it harbors a cluster containing one or several class I YSOs. It is similar to the condensation at the head of BRC13 which harbors a cluster with Class I and Class II sources: same T_{bol} (112 K for BRC13-a***), same spectral index (0.30 for S201-a, 0.37 for BRC13-a***), and high luminosity ($1400 L_{\odot}$ for BRC13-a***).

6.4. The CO outflows

CO outflows are a good tracer of ongoing embedded star formation. Class 0/I objects are very often associated with CO outflows in star forming regions (for example Curtis et al. 2010 for the Perseus region). Using JMCT HARP CO(3-2) observations, Ginsburg et al. (2011) detected 18 candidate CO outflows in the field mapped by *Herschel*. Table 7 lists some parameters of these outflows and their possible *Herschel* counterparts. Columns 1 to 6 give the identification, the center coordinates, the momentum, the energy, and the dynamical age of the components of the outflows. In column 7 we identify the central sources at the origin of the outflows as one of our point sources; this identification is based on the proximity between the center of the flow and this source; the distance separating these positions is given in column 8. (We consider that the association is probable if the distance YSO–center of the outflow is $\leq 12''$). Columns 9 and 10 give the evolutionary stage of these sources. Eleven of the 18 CO outflows are confidently bipolar; the BRC14 region has many blue and red lobes, and according to Ginsburg et al. (2011) confusion prevents pairing. Two or possibly three of the bipolar outflows are probably associated with a candidate Class 0 YSOs (S201-g, BRC14-a3, S201-s1?), six with a Class I YSO (BRC12-b, BRC12-d, IB-b, IB-d, BRC14-c1, S201-a), and one with a cluster (BRC13-cluster; this cluster contains several Class I YSOs). Furthermore, we propose to pair outflows #31 and #32 (respectively blue and red components) and to associate them with BRC14-g or -h, as they are situated close to and on each side of these condensations associated with possible Class 0/I sources. We do not identify clear counterparts for out-

Table 6. Selection of Class 0 candidates

Identification	$M_{\text{env}}(\text{BB})$ (M_{\odot})	M_{*} (M_{\odot})	M_{env}/M_{*}	$L_{\text{bol}}(\text{BB})$ (L_{\odot})	$M_{\text{env}}/L_{\text{bol}}$	$L_{\lambda \geq 350 \mu\text{m}}/L_{\text{bol}}$ (%)	$T_{\text{bol}}(\text{BB})$ (K)	$\alpha(24-100)$	Counterpart
BRC12-b	30.9	2.8–6.9	11.0–4.5	210	0.15	1.6	82	1.45	2MASS+ <i>Spitzer</i> , CO outflow
BRC12-d	32.1	4.1–5.1	7.8–6.3	92	0.35	2.6	46	1.58	<i>Spitzer</i> , CO outflow
BRC12-e	12.3	2.5–5.3	4.9–2.3	74	0.17	1.1	93	0.52	2MASS+ <i>Spitzer</i>
BRC14-a3	15.6	1.4	11.3	21	0.74	4.5	45	1.27	<i>Spitzer</i>
BRC14-b	20.2	2.2–4.15	9.1–4.9	93	0.22	2.1	78	1.22	2MASS+ <i>Spitzer</i>
BRC14-c	22.5	1.2	18.8	7	3.21	11.1	36	1.46	<i>Spitzer</i>
S201-b	1.3	0.7	1.9	18	0.07	1.1	129	1.40	2MASS+ <i>Spitzer</i>
S201-c	18.0	1.4	12.9	13	1.38	6.9	50	1.43	<i>Spitzer</i>
S201-d	5.4	4.7	1.15	118	0.05	0.4	789	−0.85	2MASS+ <i>Spitzer</i>
S201-g	46.5			40	1.16	7.5		≥ 3.31	no counterpart, CO outflow
S201-h	4.4			4	1.10	5.2		≥ 1.98	no counterpart
IB-a	14.1			7.5	1.88	8.4		≥ 3.19	<i>Spitzer</i> -MIPS
IB-b	8.7	0.7–1.4	12.4–6.2	26	0.33	2.3	78	0.9	2MASS+ <i>Spitzer</i>
IB-c	8.5			15	0.57	3.4	286	1.6	2MASS+ <i>Spitzer</i>
IB-d	20.3	1.9–3.25	10.7–6.2	63	0.32	2.0	383	0.3	2MASS+ <i>Spitzer</i>
I-5	1.7	0.45–0.9	3.8–1.9	7.5	0.23	1.9	161	1.1	2MASS+ <i>Spitzer</i>
I-11	11.1	4.6	2.4	51	0.22	1.7	246	−0.1	2MASS+ <i>Spitzer</i>

flows #22, #23, #27, #28, #29, #30, and #34. Thus we propose a *Herschel* counterpart for 12 over 13 bipolar outflows (the one exception is outflow #34).

Ginsburg et al. (2011) have detected a total of 40 candidate CO outflows associated with W5-E and W5-W. Outflow #26, associated with BRC14-c1 (also AFGL 4029-IRS1) is the most powerful of their sample (it has the highest momentum and highest energy, and a high momentum flux). Outflows #20 (BRC12-b) and #21 (BRC12-d) are also powerful.

Some other CO outflows detected by Ginsburg et al. (2011) have velocities inconsistent with the W5 complex velocities. Outflow #53 is one of them; its central velocity of -59.7 km s^{-1} indicates, according to the authors, a kinematic distance of 5.45 kpc. This velocity placed it in the outer arm, far away from W5-E. The center of this outflow is offset from the S201-a *Herschel* point source by $9.2''$. They are probably associated, and this association can possibly explain the peculiar properties of S201-a (its relatively hot dust temperature and low envelope mass - if placed at a distance of 2.0 kpc).

6.5. The YSOs color-selected IRAS sources of Karr & Martin (2003)

IRAS sources with colors of YSOs were selected by Karr & Martin (2003) to discuss star formation in W5. Twenty six YSOs were proposed by these authors in the field covered by the *Herschel* observations (see their table 4). We find that most of these IRAS sources are not YSOs but extended structures, bright fragments of PDRs or filaments. Only 3 of their sources correspond to YSOs detected by *Herschel*: IRAS02511+6023, IRAS02551+6042, and IRAS 02598+6008 which lie respectively in the direction of BRC12-b, I12, and S201-a. The sources IRAS02570+6028, IRAS02576+6017, and IRAS02593+6016 correspond to the extended sources associated with BRC13, BRC14, and Sh 201. IRAS02531+6032 lies in the direction of a bright $100 \mu\text{m}$ structure present in the center of the A H II region, just above the exciting star ex1 (see Appendix C).

6.6. Candidate prestellar condensations

Some compact condensations appear at *Herschel*-SPIRE wavelengths. They are resolved at $250 \mu\text{m}$, with diameters in the range

$0.2 \text{ pc} - 0.4 \text{ pc}$. They have no *Herschel*-PACS counterparts at $100 \mu\text{m}$, and thus they may be prestellar. Some of these sources are located at the tip (or head) of fingers pointing toward exciting stars, some other are located along filaments. We identify some of these sources in Fig. 15, and the others in Fig. 5 and Fig. 7. In Table 8 we give their coordinates, temperatures (from the temperature map), peak column densities, fluxes at $350 \mu\text{m}$ integrated in a circular aperture of radius $25''$ or $30''$ (column 6), and masses estimated from these temperatures and fluxes. These compact condensations have masses of a few solar masses. We do not know if they are gravitationally stable or not; they are potential sites of low-mass star formation.

6.7. The history of star formation in W5

The radio-continuum flux at 1420 MHz of the W5-E H II region has been measured by Karr & Martin (2003) and found to be 30 Jy. Using Equation (1) in Simpson and Rubin (1990) we estimate the ionizing photon flux to be $\sim 1.0 \times 10^{49} \text{ s}^{-1}$ (for an electron temperature of 7000 K as assumed by Karr & Martin). According to Martins et al. (2005) the ionizing flux of an O7V star, the exciting star of W5-E, is $4.27 \times 10^{48} \text{ s}^{-1}$, which is too faint by a factor of two to account for the ionization of the gas. Additionally, the observed extended $24 \mu\text{m}$ emission inside the central cavity may represent hot dust that absorbs some of the ionizing photons. Thus we are probably missing some exciting stars⁶. In the following we adopt a photon flux of 10^{49} s^{-1} for the ionizing flux of W5-E. Assuming that W5-E is a spherical H II region of radius 12 pc, we estimate its electron density to be 13.5 cm^{-3} and the ionized gas mass to be $3100 M_{\odot}$.

The sound crossing time, assuming a sound speed $\sim 10 \text{ km s}^{-1}$ in the ionized gas, is $\sim 1.2 \text{ Myr}$, which gives a

⁶ Karr & Martin (2003) used the calibration of Vacca et al. (1996) for massive stars, which gives an ionizing flux of $1.3 \times 10^{49} \text{ s}^{-1}$ for an O7V star. Thus they concluded that HD 18326 was the main exciting star of the W5-E H II region. With the calibration of Martins et al. (2005) a O7V star does not provide enough ionizing photons. HD 18326 belongs to a cluster studied by Chauhan et al. (2011a). It is possible that this cluster, containing one O7V star, also contains several late O or early B stars contributing to the ionization of W5-E. For example BD+60°0606 located 1.1 pc away from HD 18326 has been classified as a B0e star (Voroshilov et al. 1985).

Table 7. CO outflows and their Herschel counterparts

CO name	RA	Dec	Momentum ($M_{\odot} \text{ km s}^{-1}$)	Energy (10^{42} ergs)	Age (10^4 yr)	IR counterpart	d ^a ($''$)	Stage
20r	43.755678	60.597782	0.50	46.3	0.5	BRC12-b	8	Class I
20b	43.759476	60.596599	0.33	26.6	0.5	idem	7 (b)	
21b	43.967801	60.628583	0.58	41.4	1.7	BRC12-d	12	Class I
21r	43.973007	60.631990	0.08	4.3	1.7	idem		
23b	44.171755	60.769192	0.03	0.9	4.5	IB-b?	20	Class I
23r	44.177242	60.761364	0.03	1.0	4.5	idem	7 (r)	
24r	44.259049	60.698702	0.30	26.1	1.7	IB-d	9.5	Class I
24b	44.267036	60.694962	0.28	18.3	1.7	idem	6.5 (b)	
25b	45.229989	60.675872	0.09	6.8	1.0	BRC13		? (cluster)
25r	45.236467	60.677109	0.35	23.0	1.0	idem		
26b	45.372510	60.487571	0.24	26.1	1.1	BRC14-c1	3.5	Class I
26r	45.384479	60.487318	0.85	106.0	1.1	idem		
31	45.423500	60.477057	0.03	0.7		BRC14-g or h		? (cluster)
32	45.437614	60.483742	0.18	14.5		BRC14-g or h		
33b	45.452775	60.412993	0.14	10.1	3.0	BRC14-a3	11	Class 0/I
33r	45.445968	60.401839	0.20	16.1	3.0	idem	10.5 (r)	
35r	45.836843	60.471159	0.12	11.0	1.3	S201-s1	8.5	? (Class 0)
35b	45.837186	60.466721	0.12	11.0	1.3	idem	5 (b)	
36b	45.858773	60.483234	0.19	15.8	2.7	S201-g	7.5	Class 0
36r	45.855888	60.475103	0.10	6.8	2.7	idem		
37r	45.886094	60.471649	0.04	1.6	2.1	S201-e	12(r)	?
37b	45.896772	60.474332	0.06	4.2	2.1	idem		
53b	45.952487	60.338230	3.16	194.0	4.7	S201-a	9.2	Class I
53r	45.957001	60.339161	0.31	8.5	4.7	idem		

Notes. ^(a) Distance from the YSO to the center of the CO outflow, or to the blue (b) or red (r) CO component.

Table 8. Parameters of some compact condensations, candidate prestellar condensations. R is the radius of the aperture used to measure the flux.

Name	RA(2000) ($^{\circ}$)	Dec(2000) ($^{\circ}$)	T_{dust} (K)	$N(\text{H}_2)$ (cm^{-2})	R ($''$)	Flux(350 μm) (Jy)	M (M_{\odot})
Finger #1	44.7074	+61.0495	20.3	1.6×10^{21}	30	2.21	4.0
Finger #1bis	44.7348	+61.0753	20.0	1.5×10^{21}	30	2.3	4.1
Finger #2	45.8231	+60.9639	20.3	2.3×10^{21}	30	3.00	5.5
Finger #5	45.5943	+60.7624	21.0	2.0×10^{21}	30	1.48	2.7
Finger #6	45.7089	+60.6599	21.5	1.5×10^{21}	30	0.85	1.4
Finger #7	44.6145	+60.7526	19.8	3.2×10^{21}	30	2.43	4.8
Filament #1	44.1416	+60.4838	19.7	1.9×10^{21}	30	2.03	4.1
Filament #2	45.9890	+60.3777	19.7	2.6×10^{21}	30	1.84	3.7
filament #3	45.9444	+60.4745	18.3	2.2×10^{21}	25	1.05	2.5
BRC12-cond#1	43.8066	+60.7135	21.1	3.3×10^{21}	25	3.08	5.2
BRC12-cond#2	43.8665	+60.6674	22.4	3.7×10^{21}	25	2.68	4.0
BRC12-cond#3	44.0554	+60.6516	18.9	3.8×10^{21}	25	2.14	4.7
BRC12-cond#4	43.8695	+60.6050	20.1	3.5×10^{21}	25	3.10	5.9

minimum limit for the age of W5-E. The lifetime of an O7V star on the main-sequence gives the maximum limit, ~ 8 Myr, according to the models of rotating stars by Meynet & Maeder (2003, 2005). Estimating the age of a cluster is a difficult task because it requires accurate photometry and isochrones. As noted by Chauhan et al. (2011a) the use of different isochrones can introduce a systematic shift in the age. Nakano et al. (2008), using g', i' photometry and the isochrones of Palla and Stahler (1999), estimate that the mean age of the young H α emission stars located near the exciting star is ~ 4 Myr. Chauhan et al. (2011a), using V, I_c photometry and the PMS isochrones of Siess et al. (2000), conclude that star formation in the cluster is non-coeval, extending over 3.5 Myr, with a mean age of 1.3 Myr. Thus no straightforward conclusions are obtained

about the age of the exciting cluster of W5-E; it probably lies in the range 1 Myr – 4 Myr.

AND12 estimate that about 8000 M_{\odot} of molecular material are associated with the W5-E H II region. They estimate this mass from the 350 μm flux density integrated in an irregular aperture of area 2515 ($''$)², not very different from a circular aperture of radius 16.5 pc (the Sh 201 H II region lies outside this aperture). We assume that this molecular material was collected around the ionized region during its expansion. If W5-E has evolved in a homogeneous medium of density n_o atoms cm^{-3} , we can estimate n_o by equating the mass initially in the sphere of radius 16.5 pc to the mass now present in both the ionized region and its collected shell (3100 M_{\odot} + 8000 M_{\odot}). We obtain

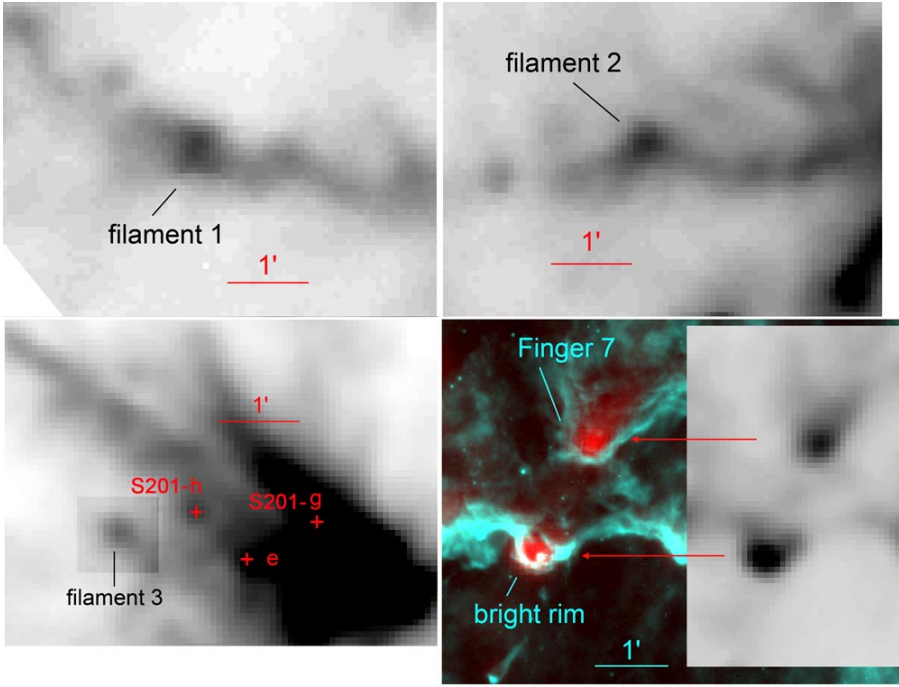


Fig. 15. Identification of some of the candidate prestellar condensations discussed in the text. The grey images show the column density. In the colour image (bottom right), red is the column density map and turquoise is the *Spitzer* 8.0 μm image. The “bright rim” feature is associated with the *Herschel* source I-5 discussed in Appendix E.

$n_0 \sim 20 \text{ atoms cm}^{-3}$. This very low value is unrealistic: such a large H II region cannot have formed and evolved during all its life into such a low density medium. (See also fig. 14 in AND12 which shows that W5-E differs from the other regions of their sample, as it has collected little material.)

An alternative situation is suggested by the distribution of the Class II YSOs, as observed by KOE08. These authors found that the Class II sources are clustered along filamentary structures. Their figures 7 and 12, and our Fig. 16 show clusters of Class II YSOs forming an east-west filament going through Sh 201, BRC14, HD 18326, the A and B H II regions between W5-E and W5-W, BRC12, to BD+60°0586, and even possibly to HD 17505 and HD 17520 (also BD+59°0552 and BD+59°0553, other exciting stars of W5-W). This suggests that a massive and probably dense filament or sheet was originally present, and that the exciting clusters of W5-E and W5-W formed inside, a few megayears ago (see Fig. 17). Their associated H II regions subsequently grew in size, first inside this dense filament, then opening toward the outside, on the low density southern side (figure 7 in KOE08 and our Fig. 1 show that the general morphology of W5-E and W5-W are similar). During their expansion these H II regions accumulated neutral material at their borders, remaining mostly ionization bounded on the north side (and density bounded on the south side for W5-E).

The H II region Sh 201 fits well with this scenario (Fig. 17). Its morphology suggests that its exciting star formed near the axis of the filament. During the expansion of the H II region the IF reached almost simultaneously the north and south low-density borders of the filament. The ionized gas expanding away from the dense regions, in two directions perpendicular to the filament, is at the origin of the bipolar structure of Sh 201 in agreement with the simulations of Bodenheimer et al. (1979). The two ionized lobes are surrounded by neutral material collected during the expansion. The dust located in this collected material, close to the IF, heated by the radiation leaking from the H II region, is observed at 24 μm and 100 μm . Colder dust, located in the collected material but farther away

from the IF, is observed at all *Herschel*-SPIRE wavelengths.

It is not clear presently if the late OB stars of the field, those exciting Sh 201, the AFGL 4029 cluster, and the A and B H II regions, formed at the same time as the exciting clusters of W5-E and W5-W or later. Are these massive stars first- or second-generation stars? The associated H II regions are much smaller than W5-E and W5-W, and thus are possibly younger, but their exciting stars are also less massive than those of W5-E (O9-B1V versus O7V). Also, they possibly formed at the center of the dense filament and/or inside dense condensations, which could explain their smaller sizes. The fact that the AFGL 4029 cluster lies in the core enclosed by BRC14, and that the A and B H II regions lie in the region of possible compression between W5-E and W5-W could indicate that they are triggered second-generation massive stars. Only accurate age determinations could answer this question, but the required accuracy is presently far from being achieved⁷.

Fig. 16 shows the distribution of the youngest sources: Class I YSOs according to KOE08 (small red dots), 100 μm point-like sources, most of them Stage I YSOs (small orange circles), our candidate Class 0 sources (large and full orange circles); it shows also the location of the CO outflows (Ginsburg et al. 2011; blue circles). These sources are all young, much younger than the exciting star of W5-E, and younger than the Class II YSOs. A timelife of about $1\text{--}2 \times 10^5$ yr is generally

⁷ It is almost impossible to estimate the dynamical age of an H II region. For this it is necessary to know at each time the physical conditions (especially the density) of the medium into which the H II region evolves (and H II regions never evolve into an homogeneous medium). The medium that we see presently has been highly modified by ionization. If we want to know the age of an H II region we must estimate the age of its exciting star(s). With the presently available tools we need accurate photometry or deep spectra at UV or optical wavelengths; an accuracy better than 1 Myr is very difficult to achieve. It is why we are unable to determine if AFGL 4029 or the A and B H II regions are younger than HD 18326, and thus if they are first- or second-generation H II regions.

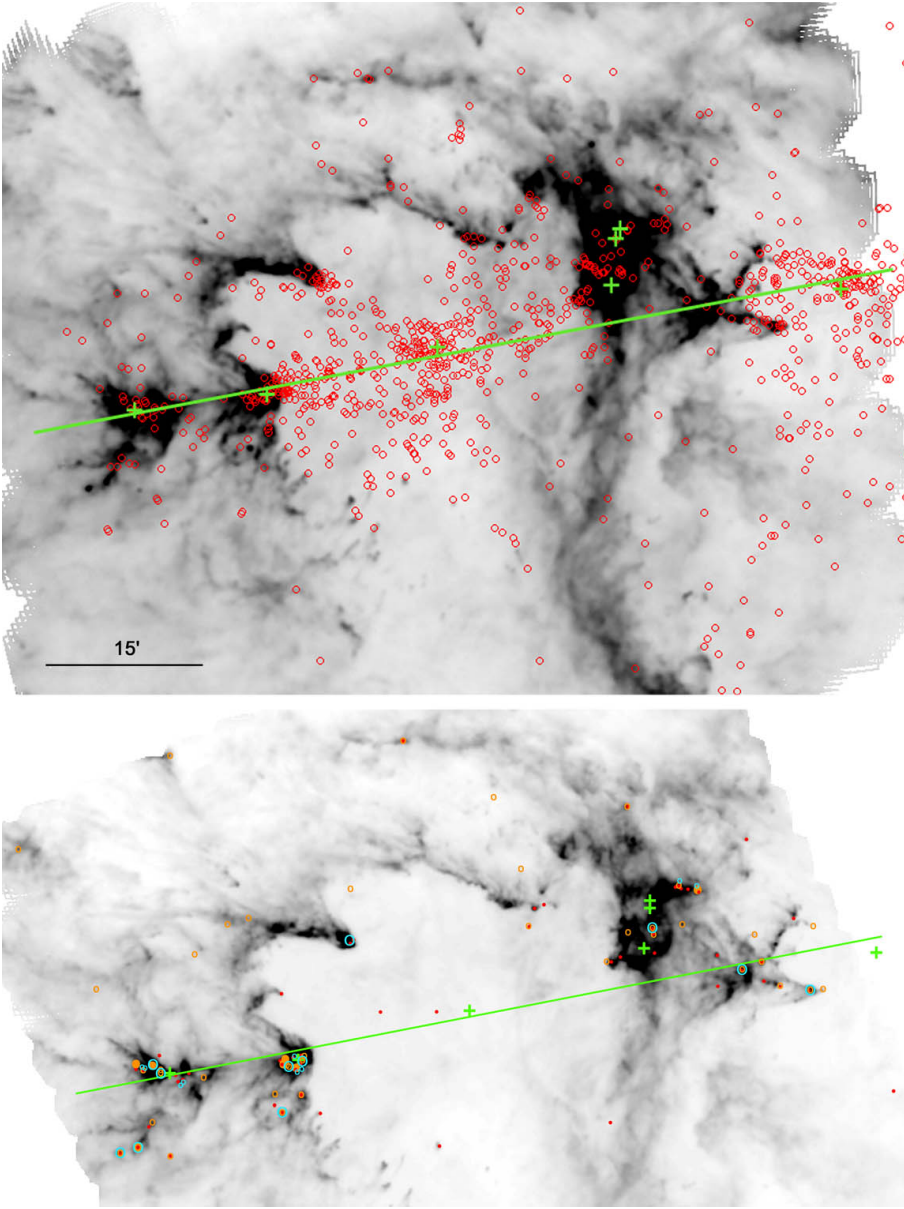


Fig. 16. Distribution of YSOs in W5-E. *Top:* Class II YSOs and exciting stars. The red circles are the YSOs identified as Class II sources by KOE08. The green crosses are the massive exciting stars of the H II regions (from east to west) Sh 201, AFGL 4029, W5-E, A & B between W5-E and W5-W, and W5-W. The green line shows the east-west filament inside which we suggest that the exciting stars of W5-E and W5-W formed first. The underlying grey-scale image is the *Herschel*-SPIRE 350 μm map. *Bottom:* The youngest sources, Class I and candidate Class 0 YSOs. The red dots are the YSOs identified as Class I sources by KOE08. The orange circles are the 100 μm point sources, mostly Stage I YSOs. The large full orange circles are our candidate Class 0 sources. The blue circles are for the CO outflows detected by Ginsburg et al. (2011). The underlying grey-scale image is the column density map.

estimated for low-mass Class I YSOs (André et al. 2000 and references therein); Class 0 are younger, with a lifetime of a few 10^4 yr (Maury et al. 2011). Fig. 16 shows that very few Class I sources are observed in the direction of the ionized region ($\leq 10\%$)⁸. Our study confirms the finding of KOE08 that the ratio of Class II to Class I YSOs is much higher in the direction of the ionized cavities than in the direction of the surrounding molecular material. We find that the youngest sources are mostly observed in the regions of higher column density, in condensations or filaments; they are preferentially located along the previously mentioned east-west filament, in the condensations adjacent to the H II region Sh 201, in the condensations enclosed by the bright-rimmed clouds, and in the region of compression nearby the A and B H II regions.

⁸ Two Class I sources, #16221 and #15522, lie in the center of W5-E. They have no 100 μm counterparts. Source #16221 has no 24 μm emission; thus, its classification is doubtful. Three Class I sources lie in the center of Sh 201, sources #5, #6, and #7; we have seen in Appendix C that their classification is doubtful.

Another characteristic of the young population is the lack of luminous sources. All the confirmed Stage I sources have luminosities smaller than a few hundred solar luminosities, and central objects of only a few solar masses. The only locations of high-mass star formation are the massive condensation enclosed by BRC14 (with the AFGL4029 cluster already formed and c1), and the massive condensation at the east waist of Sh 201 (with the source s1, if confirmed). Presently we are in a phase of low-to intermediate-mass star formation. Is it the end of massive-star formation in this region? We cannot yet answer this question, as we note that there are two massive condensations that contain enough mass for further high-mass star formation.

6.7.1. Processes of triggered star formation

Several processes can trigger star formation on the borders of H II regions. Four of them are described and illustrated in Deharveng et al. (2010; their fig. 4). They are: 1) small-scale instabilities of various origins in the collected layer; 2) large-scale gravitational instabilities along the collected layer leading to the formation of massive fragments (the collect and collapse process); 3) ioniza-

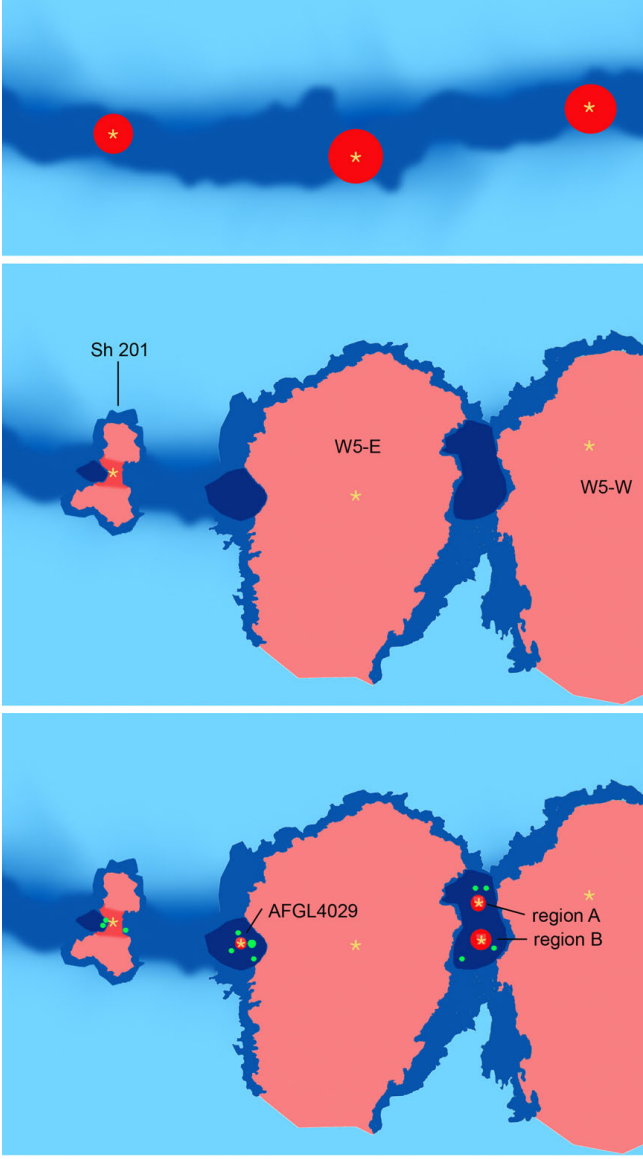


Fig. 17. Scenario of star formation in the W5 region. *Top:* The exciting stars of Sh 201, W5-E and W5-W formed in a dense filament. The exciting star of Sh 201 formed in the middle of the filament whereas those of W5-E and W5-W formed on its border. *Middle:* The expanding H II regions have burst out of the filament. In Sh 201 the ionized gas flows away of the filament in two opposite directions, forming a bipolar H II region. *Bottom:* A new generation of stars is forming in the densest condensations, adjacent to the first generation H II regions.

tion radiation acting on a turbulent medium; 4) radiation-driven compression of pre-existing dense clumps (the RDI process). It is often difficult to prove that star formation has been triggered because the age difference between two stellar populations is difficult to estimate with accuracy (for example Martins et al. 2010), and we are often left with only morphological arguments. Furthermore it may be difficult to identify the processes at the origin of the triggering. What can we learn from the *Herschel* observations of W5-E?

Neutral material has been collected around the W5-E H II region, about $8000 M_{\odot}$ according to AND12. It is observed for example in the direction of PDR1 and PDR2, at the south border of BRC13, and at the periphery of BRC12. The column den-

sity in these structures ($\leq 3 \times 10^{21} \text{ cm}^{-2}$), however, is not high enough to allow triggered star formation via the collect and collapse process (according to Whitworth et al. 1994 who estimate that the collapse occurs when the column density is higher than about $6 \times 10^{21} \text{ cm}^{-2}$). Collected neutral material is also observed around the bipolar nebula Sh 201. Around the bipolar lobes the column density is rather low, $\leq 10^{21} \text{ cm}^{-2}$, and no star formation is observed. The column density is much higher at the waist of the nebula. There, we hypothesize that it is the material of the dense parental filament that has been collected. Several candidate Class 0 and Class I YSOs are detected in this dense collected material (S201-f, -g, -h, -s1, -s2; Fig. C.5, Fig. C.7); CO outflows are present as well. These YSOs are second-generation sources, probably younger than 10^5 yr. Some of these Class 0/I or candidate starless cores are massive (S201-g and possibly S201-s1-s2) and possibly susceptible to transform via accretion into massive stars. We are dealing with triggered star formation in the act, in the dense layer present at the waist of the Sh 201 H II region. But, due to the poor resolution, the exact process operating is difficult to identify.

Triggered star formation is possibly observed in other locations:

- In the pillars, on the east and south-east border of W5-E. As discussed in Sect. 6.2 the best explanation for these numerous structures seems to be the interaction of the ionizing radiation of the central exciting star with a low-density turbulent medium. Some low-mass YSOs have already formed at the tip of these structures, which are observed as Class I or Class II sources. Low-mass condensations are also detected at the head of the pillars, which are potential sites of future low-mass star formation. We suggest that this process may explain the dispersed population of low-mass Class II YSOs detected by KOE08. This process possibly leads to a continuous star formation: as shown by Gritschneider et al. (2009a) the ionizing radiation of an O star provides a good mechanism to sustain and even drive turbulence in the parental molecular cloud (the turbulence does not decay with time), and thus, this star formation process may remain at work during a long period.

- In the vicinity and possibly inside the bright-rimmed clouds. BRC13 and BRC14 enclose massive condensations that harbor clusters and several Class I YSOs. BRC12 is probably different: it is bordered by collected material, but does not enclose a massive condensation; it harbours several individual Class I YSOs but no cluster. Star formation in the vicinity of bright rims has possibly been triggered by the ionizing radiation (i.e. by the expanding H II region), as shown by several indicators:

1. The morphology of these BRCs, pointing towards the exciting star
2. The comparison of the pressure of the gas inside and outside the rims. Morgan et al. (2004), using radio-continuum observations, estimated the density and the pressure in the dense ionized layer bordering the BRCs. Morgan et al. (2008), using SCUBA observations, estimated the temperature and the density of the molecular condensations enclosed by the BRCs. They showed that the pressure of the ionized gas is higher than that of the molecular gas in BRC12, BRC13, and BRC14. As a consequence, an ionization front of D type and its associated shock front progress inside these structures, compressing material and thus providing favorable conditions for star formation.
3. The observed “small-scale sequential star formation” in the vicinity of the rims. This was first advocated by Sugitani et

al. (1995) to explain the distribution of near-IR sources associated with BRCs. They observed that the bluer sources, thus the older, were located closer to the exciting stars, sometimes outside the BRCs, while the redder sources, the younger, were located inside the BRC close to the enclosed IRAS source. This suggests a propagation of star formation from the side of the exciting star to the IRAS position. BRC12, BRC13, and BRC14 are among the BRCs that have this configuration (Sugitani et al. 1995). The same effect was observed with the $H\alpha$ emission stars, which are pre-main-sequence stars of T Tauri type, and thus are low-mass Stage II YSOs. $H\alpha$ emission stars have been identified in the vicinity of BRC12, BRC13, and BRC14 (Ogura et al. 2002). These stars are concentrated near the tip of the BRCs, most often just outside the rims, on the side of the exciting stars (see for example Appendix A and Fig A.1). Several photometric studies of these near-IR stars and/or these $H\alpha$ emission stars aim at estimating their ages (Ogura et al. 2002; Chauhan et al. 2009; Chauhan et al. 2011a). The recent study by Chauhan et al. (2011a), which also concerns mid-IR *Spitzer* sources, shows that in all the BRCs the YSOs lying on or inside the rim are younger (by several 10^5 yr; Appendix A & B) than those located outside it. This sequential small-scale star formation suggests that star formation occurs in the vicinity of the IF bordering the BRCs, and moves with it away from the central exciting star.

The bright-rimmed clouds differ from the pillars in terms of size, mass enclosed, and stellar content (clusters versus isolated low-mass YSOs). Thus, they probably differ in origin, being due not to turbulence but more probably to pre-existing dense structures, condensations or filaments affected by the stellar radiation⁹. The RDI process is generally proposed to explain star formation in their vicinity. But the whole picture remains unclear. Several simulations of the RDI process have recently been presented and discussed (references in Sect. 6.2). They all include self-gravity, but only the Haworth & Harries' simulations take into account the diffuse radiation field. They also differ in their initial conditions: the mass ($5 M_{\odot}$ to $96 M_{\odot}$), size (0.3 pc to 1.6 pc), density distribution in the pre-existing globule (uniform or centrally peaked), and the ionizing photon flux reaching the globule. Thus it is difficult to draw straightforward conclusions. All simulations (except those of Miao et al. 2009) point to the importance of the ionizing photon flux. We estimate that these fluxes are $\leq 1.5 \times 10^9 \text{ s}^{-1} \text{ cm}^{-2}$ and $\leq 0.9 \times 10^9 \text{ s}^{-1} \text{ cm}^{-2}$ respectively at the center of the condensations enclosed by BRC13 and BRC14 (they may even be smaller if there is a projection effect). For all simulations this is a case of low ionizing photon flux, and the consequence for all simulations is that the evolution of the globule is slow, and that star formation is delayed (it occurs after ~ 0.6 Myr in Gritschneder et al. 2009b; this is compatible with the ages estimated by Chauhan et al. 2011a for the enclosed clusters). The BRCs' morphology is more difficult to account for. BRC13 and BRC14 enclose massive and dense condensations (respectively $\sim 22 M_{\odot}$ in a radius of 0.29 pc and more than $200 M_{\odot}$ in a radius of 0.48 pc); the associated clusters lie in the very centers of the condensations, far from the bordering IF (respectively at 0.29 pc and 0.81 pc from the IF in BRC13 and BRC14; see Fig. A.1 and B.4; this distance may be smaller if there is a projection effect). This is not predicted by the simu-

lations which show star formation at the tip of a filamentary feature (the ultimate fate of a globule before dispersion). Also, no simulations consider a globule massive enough to allow the formation of a whole cluster, as is observed in BRC14. Simulations considering the action of a low ionizing photon flux falling upon a massive condensation or filament are missing. On the other hand, the *Herschel*-SPIRE images lack the resolution necessary to show the exact morphology of the enclosed condensations. As a consequence, if it is clear that the BRC13 and BRC14 structures are presently shaped by the UV radiation of HD 18326, it is difficult to determine if the formation of the enclosed clusters has been triggered or not by this radiation.

The configuration of the east condensation at the waist of of Sh 201 has similarities and differences with the condensation enclosed by BRC14. Here again the condensation is massive and dense ($\sim 230 M_{\odot}$ in a radius of 0.48 pc), and is bordered and compressed by ionized gas on one side. But the similarities stop there: the ionizing photon flux is much higher, of the order of $7 \times 10^{10} \text{ s}^{-1} \text{ cm}^{-2}$, because the star lies much closer to the IF (if there is no strong projection effect). This condensation is experiencing high ionizing flux, leading to a rapid star formation (on a time scale of 0.1 – 0.2 Myr according to Gritschneder et al. 2009b or Bisbas et al. 2009). The Class 0/I YSOs associated with this condensation (S201-g, S201-s1-s2) lie in the direction of the filaments separating the ionized and molecular material, and thus are very near to the IF (Fig. C.7). They formed not at the center of the condensation but more probably in the compressed layer bordering the condensation. Triggering is more obvious than in the case of BRC14, possibly via the RDI process. But, here again, we do not have the resolution necessary to unravel the small scale morphology of the condensation and the exact location of the YSOs.

7. Conclusions

We have taken advantage of the very simple morphology of the W5-E H_{II} region to study star formation, and the possible impact of triggering. This region has various features interesting in terms of star formation: bright rims, pillars, and a nearby bipolar H_{II} region. For the first time, thanks to *Herschel*-PACS and SPIRE observations, we have access to the distribution of the cold dust over the whole region, tracing the dense molecular material inside which stars form. Also we have identified the youngest protostellar objects, the Class 0/I YSOs, via the emission of their cold envelopes.

Some of our conclusions are still uncertain, for various reasons: i) the masses of the envelopes and condensations are uncertain because of uncertainties in the opacity of the dust. The nature of the dust varies with its location (or environment). How this affects its opacity is not yet well established; ii) we have shown that the SED fitting tool of Robitaille et al. (2007), which presently does not include dust colder than 30 K, does not allow us to accurately estimate the properties of the deeply embedded Class I YSOs (especially those concerning their cold envelopes). New models will soon be available. Meanwhile it seems far better to use a modified blackbody model to fit the envelopes of Class 0/I YSOs; iii) some sources, such as S201-a, may not be part of the W5-E complex; only the knowledge of the source velocity (which requires high resolution molecular observations) could confirm the association. Furthermore a few associations between *Spitzer* and *Herschel* point sources are problematic due to the resolution of the *Herschel*-SPIRE observations; in some cases we may see with *Herschel* a compact condensation adjacent to a Class I YSO rather than the YSO's envelope.

⁹ To summarize we suggest that pillars are due to stellar radiation acting on a low-density turbulent medium, whereas bright-rimmed clouds result from the action of stellar radiation on pre-existing dense condensations

This could explain why a number of SEDs are not well fitted (for example IB-c and IB-f in Appendix D, I-11 in Appendix E).

Remembering these uncertainties our conclusions are the following:

- W5-E is a bubble H II region: the ionized region is surrounded by a ring of PAHs emitting at $8.0 \mu\text{m}$, observed by MSX and *Spitzer*. Several bubble H II regions have been observed with *Herschel* (Anderson et al. 2012). W5-E differs from most of these regions, possibly because of its large size. As a whole its dust temperature is more uniform than that of the regions studied in Anderson et al. (2012), everywhere higher than 17.5 K, and its column density is rather low, everywhere lower than 10^{23} cm^{-2} . Neutral material has been collected around the ionized region during its expansion. We observe this material at SPIRE wavelengths as a layer of cold dust adjacent to the IFs. The temperature in these PDRs is of the order of 21 K, and the column density is low, of the order of a few 10^{21} cm^{-2} . According to the existing models, it is too low for star formation to proceed; very few YSOs are observed in these directions.

The total mass of collected material is $\sim 8000 M_{\odot}$ (Anderson et al. 2012), which is low with respect to the volume of the H II region, and differs from the other studied regions. We suggest that W5-E has not formed and evolved in a spherical homogeneous cloud, but in a dense filament (or sheet). One signature of this scenario is the distribution of Class II sources (along a filamentary structure) according to Koenig et al. (2008). Another signature is the detection by *Herschel*-SPIRE of the remaining parental filament extending east-west over more than 15 pc, from BRC14 to regions east of Sh 201.

- We have detected and measured 50 point-like sources and a few compact condensations. Their SEDs have been fitted using a modified blackbody model. The envelopes of the YSOs are cold, with a mean temperature of $15.7 \pm 1.8 \text{ K}$. The more extended condensations, harboring young clusters (enclosed by BRC13 and BRC14), are hotter, with a temperature around 26 K. The far away point source, S201-a, lying in the background of W5-E at 5.45 kpc, belongs to this last group.

- The masses of the envelopes of the YSOs detected by *Herschel* lie in the range $1.3 M_{\odot} - 47 M_{\odot}$.

- The spectral index between $24 \mu\text{m}$ and $100 \mu\text{m}$ cannot, alone, be used to separate infallibly Class I and Class II sources.

- We have used several indicators to identify the youngest YSOs, the Class 0 sources. The purely observational indicators, $M_{\text{env}}/L_{\text{bol}}$ and T_{bol} , are probably the best ones in distant regions such as W5-E. Eight of the fifty $100 \mu\text{m}$ point sources are good candidate Class 0 YSOs; three more are possible candidate Class 0. They are BRC14-c, BRC14-d, BRC14-i, BRC14-j, S201-c, S201-g, S201-i, IB-a (also possibly BRC14-a3, S201-s1, -s2).

- Eighteen CO outflows have been detected in the field observed by *Herschel* (Ginsburg et al. 2011). Thirteen of them are bipolar. We found a *Herschel* point source counterpart for twelve of them (3 are possibly associated with a Class 0 YSO, 6 with a Class I, 2 with clusters containing Class I YSOs, 1 with an YSO of uncertain evolutionary stage).

- Most of the Class 0/I YSOs are located in regions of high column density: near condensations enclosed by BRCs, near condensations at the waist of Sh 201, or in the vicinity of the compact A and B H II regions. These YSOs are not very massive. We have used the SED fitting-tool of Robitaille et al. (2007) to estimate the parameters of the central object and of the disk in Class I and Class II sources (their masses, the disk and envelope accretion rates). We find only central sources of low-

intermediate-mass (except for c1 and c2 in BRC14). But about fifteen sources have massive envelopes ($10 M_{\odot} - 47 M_{\odot}$); some of these may have high envelope accretion rates and thus may become massive stars.

- We observe probable triggered star formation in several locations:

1. Near the numerous pillars present south-east of W5-E, in a region of low density. Some Class I and Class II low-mass YSOs have been detected there by *Spitzer* (Koenig et al. 2008). The *Herschel* $250 \mu\text{m}$ map shows the presence of low-mass condensations at the head of the pillars. We propose that the UV radiation of the central exciting star, irradiating a turbulent low-density medium, is responsible for these structures and the YSOs formation, in good agreement with the simulations of Gritschneider et al. (2010).
2. Triggered star formation was already identified in the well known bright-rimmed clouds BRC12, BRC13, BRC14. *Herschel* images confirm the presence of condensations enclosed by BRC13 and BRC14, with masses respectively of a few tens to a few hundreds of solar masses. The massive BRC14 condensation harbors a cluster exciting a compact H II region, several Class I YSOs, at least four good candidate Class 0 sources, and several CO outflows. The RDI process, in the scope of a low ionizing photon flux, is possibly at work there. But *Herschel*-SPIRE images lack the resolution necessary to display the morphology of the condensations and allow a comparison with the RDI simulations.
3. At the waist of the bipolar H II region Sh 201.

- The *Spitzer* images show many “fingers” (structures intermediate in size between the large BRCs and the narrow pillars) pointing to the central exciting star, and located in the outskirts of the W5-E region, far away from the main PDRs. *Herschel*-SPIRE images show the presence of condensations of a few solar masses (diameters in the range $0.2 \text{ pc} - 0.4 \text{ pc}$) at the head of these fingers; these condensations are good candidate prestellar cores as they are bordered by high pressure ionized gas. They are shaped by the UV radiation of the exciting star. This suggests that the main IF limiting the H II region presents holes, allowing the ionizing radiation to leak and interact with the surrounding neutral medium, at large distances from the exciting star.

- The bipolar nature of the Sh 201 H II region, well demonstrated by the *Spitzer* images (Koenig et al. 2008), is confirmed by the *Herschel* observations. The $100 \mu\text{m}$ image shows the warm dust surrounding the two ionized lobes. The SPIRE images show a layer of cold dust of low column density surrounding the two lobes; for the first time we see the material collected during the expansion of this H II region. The SPIRE images also show a cold filament of high column density on each side of the H II region (the parental filament). Two massive cold condensations are observed at the waist of the nebula, on each side of the ionized region. For the first time star formation is observed in the direction of the most massive one (mass $\geq 200 M_{\odot}$); six $100 \mu\text{m}$ point sources are detected there, two of them are good candidate Class 0 sources; three more are possible candidate Class 0 YSOs. The fact that they are observed in the direction of filaments, in the zone of interaction between the ionized gas and the molecular condensation, strongly suggests that their formation has been triggered by the expanding Sh 201 H II region.

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Appendix A: The bright-rim cloud BRC13 and vicinity

BRC13 (Sugitani et al. 1991) is the textbook example of a bright-rimmed cloud: it resembles a finger pointing to HD 18326, the exciting star of W5-E. An IR source IRAS 02570+6028 lies at the tip of the finger, inside the rim; its luminosity is $1440 L_{\odot}$ according to Sugitani et al. (1991; from IRAS fluxes) or $490 L_{\odot}$ according to Morgan et al. (2008; from SCUBA fluxes); both estimates are for a distance of 2 kpc.

Fig. A.1 shows the morphology of BRC13. The head of BRC13 is symmetric (over a length of 1.7 pc) with respect to the direction of HD 18326, the exciting star of W5-E. This is strong evidence that the ionizing radiation of HD 18326 has shaped this structure. Fig. A.1 shows that a small IR cluster lies inside BRC13, at its tip. This cluster is likely the counterpart of IRAS 02570+6028. According to KOE08 it contains several Class II and Class I YSOs. Sources #1 and #2 in Table A.1 have been identified by KOE08 as Class I YSOs. YSO #2 has no detectable counterpart at *Herschel* wavelengths. Source #3 (identified in Fig. A.1) is not present in the catalogue of KOE08; however it is probably also a Class I YSO: it strongly resembles YSO #1 as both have $24 \mu\text{m}$ and $100 \mu\text{m}$ counterparts. The brightest source at $24 \mu\text{m}$ is a Class II source, BRC13-a, located at the center of a region of extended emission visible in $\text{H}\alpha$ and at all *Spitzer* wavelengths. It is this YSO and the extended emission region (plus probably the nearby Class I YSOs #1 and #3) that are seen in the PACS and SPIRE images. At these wavelengths we are unable to separate the central object from its associated extended emission zone. We have measured the fluxes of this extended source using apertures of diameter $30''$, $45''$, and $60''$ (respectively one, two, and three asterisks in Table A.1); the first aperture is that used by Morgan et al. (2008) to measure the SCUBA $450 \mu\text{m}$ and $850 \mu\text{m}$ fluxes, but only the last aperture covers the entire extended emission zone. We have used the same apertures to measure the fluxes in the *Spitzer* bands (asterisks). We list these measurements in Table A.1.

Fig. A.1 shows the dense molecular material present at the head of BRC13, around BRC13-a. The diameter at half intensity of the molecular condensation is 0.31 pc. Its peak column density is $N(\text{H}_2)=1.7\times 10^{22} \text{ cm}^{-2}$, thus a density $n(\text{H}_2)\sim 1.7\times 10^4 \text{ cm}^{-3}$. Its temperature, estimated from the *Herschel* fluxes, is in the range 24.4 K to 27.8 K (the smallest aperture giving the hotter temperature; Table 5). This differs slightly from the temperature of 23 K obtained by Morgan et al. (2008). The mass of this condensation is estimated to be $\sim 6.1 M_{\odot}$ in the $30''$ diameter aperture and increases to $21.8 M_{\odot}$ in the $60''$ diameter aperture (considering only the *Herschel* data; Table 5). (from SCUBA data Morgan et al. estimate a mass of $35 M_{\odot}$ in the $30''$ central aperture, corrected for a distance of 2 kpc and a dust opacity at $850 \mu\text{m}$ of $1.25 \text{ cm}^2 \text{ g}^{-1}$). We suggest that this condensation is heated from the inside by the nearby Class II object BRC13-a; this YSO is possibly partly embedded and appears to have carved a cavity open in our direction. This has created a path of low extinction allowing us to see BRC13-a and the light reflected by the borders of the cavity, even at optical wavelengths.

Another fainter condensation (called b) is also present along the axis of BRC13, at the back of the brightest condensation, and thus farther away from the exciting star (Fig. A.1). We measured its *Herschel* fluxes using an aperture of diameter $90''$ (measurements in Table A.1). The mass of this condensation is estimated to be $55 M_{\odot}$ for a temperature of 19.5 K found in the temperature map (Fig. 3). The column density decreases along the axis when

going away from the tip of the rim. Dense material is also present on the south border of BRC13 (and not at its north border, a deviation from axial symmetry). The observed column density is not very high, of the order of $3\times 10^{21} \text{ cm}^{-2}$. This material is probably made of material collected during the expansion of W5-E. Alternatively it may be part of a east-west pre-existing filament as shown by Fig. 4. The column density is not high enough for star formation to proceed via the collect and collapse process, and indeed no YSOs are detected there.

At $\text{H}\alpha$ wavelength BRC13 appears as an extinction region surrounded by a bright rim. Therefore the molecular condensation inside the rim, responsible for the extinction, is located slightly in front of the ionized region. The NVSS image shown by Morgan et al. (2004) shows faint radio-continuum emission from the very head of BRC13 (the ionized boundary layer seen also in $\text{H}\alpha$); they estimated the electron density there to be $\sim 250 \text{ cm}^{-3}$. Morgan et al. (2004) have shown that the pressure in the ionized boundary layer is much higher than the pressure in the enclosed molecular material. (This is confirmed by the *Herschel* observations.) Thus a D type ionization front progresses inside the BRC, compressing the molecular material and potentially triggering star formation. The observed IR cluster possibly results from this triggering.

Another signature of triggered star formation in BRC13 is the small-scale sequential star formation observed in this object. Numerous $\text{H}\alpha$ emission line stars have been detected in the vicinity of BRC13 by Ogura et al. (2002). These stars are YSOs, pre-main sequence stars in the T Tauri/Herbig AeBe phase. They are identified on Fig. A.1. They are located near the tip of the bright rim, at the head of the rim, or just outside the rim. Most of these emission line stars are also Class II YSOs in KOE08. The youngest Class I YSOs are located inside the bright rim. This suggests small-scale sequential star formation, with star formation proceeding from the side of the exciting star outward away from the central H II region. Chauhan et al. (2011a) used deep V, I photometry to estimate the age of the $\text{H}\alpha$ emission stars; they obtained different ages, $2.44\pm 1.37 \text{ Myr}$ and $1.61\pm 1.41 \text{ Myr}$ respectively for 24 stars outside the rim and 10 stars inside the rim. This confirms the small-scale sequential star formation (however the large scatter in the stellar ages is worrisome).

Appendix B: The bright-rim cloud BRC14 and vicinity

BRC14 is another bright-rimmed cloud catalogued by Sugitani et al. (1991). It harbours the bright IR source IRAS 02575+6017 (also AFGL 4029, Price and Walker, 1976). The luminosity of this IRAS source is $\sim 1.9\times 10^4 L_{\odot}$ (Snell et al. 1988; corrected for a distance of 2 kpc). Beichman (1979) has shown that AFGL 4029 is composed of two mid-IR sources, IRS1 and IRS2, separated by $22''$.

IRS1 and IRS2 are radio sources, respectively G138.295+1.555 and G138.300+1.558 (Kurtz et al. 1994, Zapata et al. 2001). Both are also small $\text{H}\alpha$ emission regions (see Figure 1 in Deharveng et al. 1997). G138.300+1.558 is a thermal H II region harbouring a small cluster dominated by a B1V star (Deharveng et al. 1997; their star #26, affected by a visual extinction $\sim 8 \text{ mag}$). IRS1 is associated with a YSO and its reflection nebula (Deharveng et al.; their star #25, affected by a visual extinction $\sim 30 \text{ mag}$). A high velocity optical jet originates from IRS1 (Ray et al. 1990). The *Spitzer* images confirm the presence of two bright IR components, a cluster around a bright star in the direction of IRS2, and a bright YSO

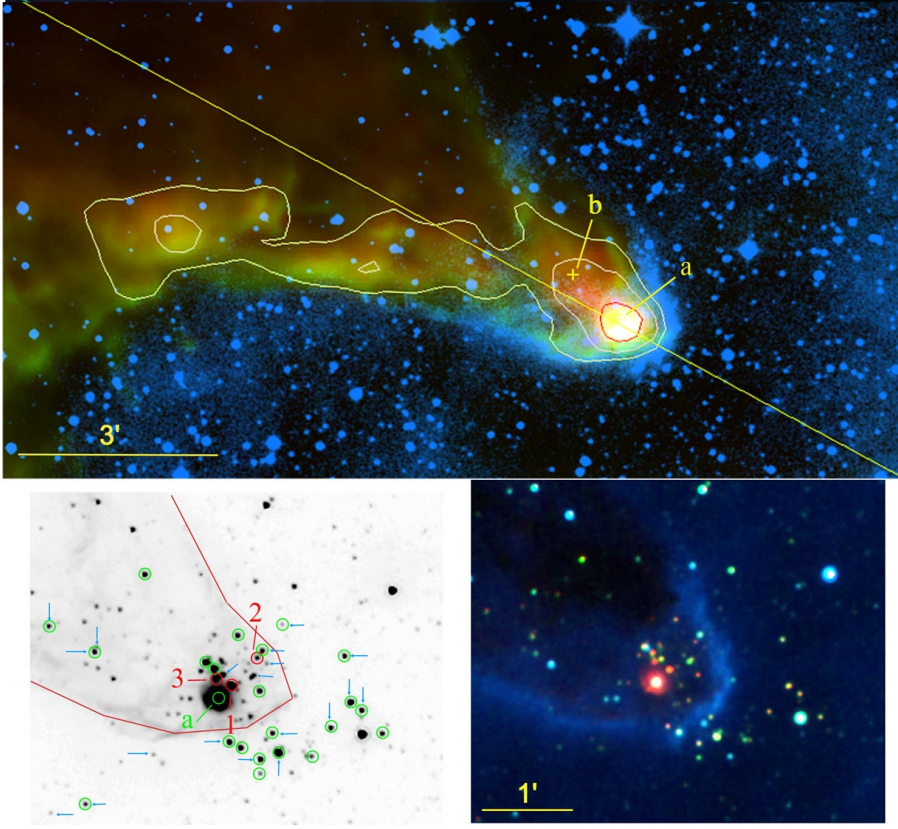


Fig. A.1. The bright-rimmed cloud BRC13. *Top:* Composite colour image: red is the column density map from *Herschel* data, green is the *Herschel*-PACS image at 100 μm , and blue is the DSS2-red image. The peak column density is $1.7 \times 10^{22} \text{ cm}^{-2}$. The contours correspond to column densities of 1/2, 1/4, and 1/8 of the peak value. The yellow line shows the direction of the exciting star HD 18326. The field size is $14.7' \times 7.5'$. North is up and east is left. *Bottom right:* The cluster inside BRC13: red is the *Spitzer* emission at 4.5 μm , green is the 2MASS *K* stellar emission, and blue is the $\text{H}\alpha$ emission from the ionized boundary layer (image obtained at the 1.2-m telescope of the Observatoire de Haute-Provence). *Bottom left:* Identification of the YSOs discussed in the text and present in the KOE08 catalogue. The grey-scale underlying image is the *Spitzer* 4.5 μm data. The red and green circles are respectively the proposed Class I and Class II YSOs. The blue arrows point to the $\text{H}\alpha$ emission stars (classical T Tauri-like stars) identified by Ogura et al. (2002).

Table A.1. Photometry of point sources: the bright-rim cloud BRC13 and vicinity

Name	#	RA(2000) ($^{\circ}$)	Dec(2000) ($^{\circ}$)	<i>J</i> mag	<i>H</i> mag	<i>K</i> mag	[3.6] mag	[4.5] mag	[5.8] mag	[8.0] mag	[24] mag	Class	S(100) Jy	S(160) Jy	S(250) Jy	S(350) Jy	S(500) Jy
a	16614	45.235304	60.669934	12.36	11.68	11.23	10.50	9.99			0.96	II					
a*													78.95	67.42	20.07	6.99	2.55
a**							7.57**	7.36**	4.65**	2.96**	0.18**		112.08	113.54	35.83	13.36	5.02
a***							7.33***	7.16***	4.48***	2.80***	0.033***		137.22	152.19 *	49.55	19.25	7.48
b		45.2615	60.6843												56.2	27.1	12.8
1	16594	45.230365	60.672069		15.08	13.51	11.18	10.05	8.82	7.57	3.32	I					
2	16555	45.221034	60.677145		15.54	14.62	12.89	12.30	11.29	9.63	5.46	I					
3*		45.2359*	60.6735*														

associated with a reflection nebosity in the direction of IRS1. IRS1 has been observed at high resolution in the 8 μm –13 μm range by Zavagno et al. (1999) and at 24.5 μm by de Wit et al. (2009). De Wit et al. estimate a flux density of 680 Jy for IRS1 from the *Spitzer*-MIPS observations at 24 μm .

The bright-rimmed cloud BRC14 borders a massive molecular cloud. This cloud has been mapped in $^{12}\text{CO}(1-0)$ by Loren and Wootten (1978), Snell et al. (1988), Carpenter et al. (1990; 2000), Karr and Martin (2003). High resolution maps have been obtained by Niwa et al. (2009; HPBW=15.6'') in $^{13}\text{CO}(1-0)$ and $\text{C}^{18}\text{O}(1-0)$. The condensation enclosed by BRC14 is their clump 7, the brightest one, with a peak column density in the range $5.8\text{--}9.4 \times 10^{22} \text{ cm}^{-2}$ (depending on the transition), and a mass of 740 M_{\odot} .

A CO outflow was detected by Snell et al. (1988) in the direction of AFGL 4029 (size $\sim 0.55 \text{ pc}$; age $\sim 1.7 \times 10^5 \text{ yr}$). Their resolution (HPBW $\sim 45''$) is not sufficient to associate it definitively with IRS1 or IRS2 (see figure 2 in Deharveng et al. 1997). An H_2O maser has been detected in the direction of IRAS 02575+6017 by Churchwell et al. (1990); its velocity, $V_{\text{LSR}} = -33.3 \text{ km s}^{-1}$, indicates that it is associated with the

region. Once again the resolution of the observations do not allow a direct association with IRS1. Methanol maser emission at 6.7 GHz has been searched for but not detected in the direction of the IRAS source (Slysh et al. 1999).

Small-scale sequential star formation is also observed in the direction of BRC14 (Matsuyanagi et al. 2006, Chauhan et al. 2011a and references therein). Using deep photometry Chauhan et al. (2011a) estimated a mean age of $1.01 \pm 0.73 \text{ Myr}$ for 18 YSOs on or inside the rim, and an age of $2.32 \pm 1.22 \text{ Myr}$ for 58 sources outside the rim.

B.1. YSOs in the field of BRC14

The *Herschel* observations show a bright massive condensation enclosed by BRC14, and numerous point-like sources. These objects are identified in Fig B.1, listed in Table B.1, and commented on below. We show a few SEDs in Fig. B.2. The SEDs have been fitted allowing the interstellar (external) extinction to range between 1.96 mag and 8 mag.

- BRC14-a: This source is observed in the direction of three nearby stellar sources (separated by some 6''), all classified as

Table B.1. Photometry of point sources: the bright-rim cloud BRC14 and vicinity

Name	#	RA(2000) (°)	Dec(2000) (°)	<i>J</i> mag	<i>H</i> mag	<i>K</i> mag	[3.6] mag	[4.5] mag	[5.8] mag	[8.0] mag	[24] mag	Class	S(100) Jy	S(160) Jy	S(250) Jy	S(350) Jy	S(500) Jy
a1	17120	45.443647	60.406051				13.91	13.20	12.39	11.40		I					
a2	17121	45.444053	60.404731				14.48	13.16	12.09	11.00		I					
a3	17125	45.447332	60.405626				14.30	12.72	11.37	10.14	4.97	I	1.88	4.22	4.84	3.59	1.63
b	17033	45.384975	60.431850	14.95	13.32		10.72	9.44	8.37	7.33	3.17	I	9.21	14.82	11.08	7.05	4.86
c	17157	45.468562	60.434342				14.96	13.64	12.77	11.44	6.67	I	0.51	1.66	2.83	2.71	2.77
d		45.40186	60.47630								no		4.36	8.38	7.84:		
e	17052	45.400109	60.477348	15.08	13.63	12.09	10.94	10.02	9.19	3.71		I	4.72	13.14	7.84:		
f	17012	45.375738	60.493892	14.06	12.42	10.87	10.20	9.85	9.27			II	4.8:				
g		45.43024	60.47904										1.74:	7.71	7.15:		
g1	17094	45.430795	60.480435				11.37	10.63	9.96	9.26	4.67	II					
g2	17091	45.429855	60.478460	15.88	14.02	12.99	12.35	11.88	11.47	11.01		II					
h		45.43635	60.48052										2.22:	6.81	5.43:	9.61 (g+h)	
h1	17111	45.436898	60.480450	15.88	13.86	12.48	10.76	10.12	9.58	8.98		II					
i		45.44198	60.48221								no		1.07	4.48	4.95:		
j		45.43740	60.48739								no		0.94	2.03			
l	17160	45.470193	60.413711			14.78	12.35	11.24	10.24	9.35	5.49	I	0.13:				
2	17132	45.452133	60.485071	15.97	14.88	11.04	10.18	9.27	8.42	5.04		I	0.16:				
c1	17020	45.380310	60.486940	9.90	7.25	5.27	4.45	3.27	2.56			I					
c2	17044	45.393692	60.486907			8.63	8.10	6.18	4.34			I					
c*													1503	1319	407.7	157.1	60.4
c**													1628	1508	476.5	189.3	77.2
c***													1731	1661	534.1	216.0	86.9

Table B.2. Characteristics of the YSOs: the bright-rim cloud BRC14 and vicinity

Name	#	χ^2	M_* (M_\odot)	T_* (K)	M_{disk} (M_\odot)	\dot{M}_{disk} ($M_\odot \text{ yr}^{-1}$)	\dot{M}_{env} ($M_\odot \text{ yr}^{-1}$)	L (L_\odot)	$A_V(\text{ext})$ (mag)	i (°)	Stage
a3	17125	149	1.4	4218	1.1e-2	1.0e-7	4.7e-4	21	1.96	49	I
b	17033	151 (−187)	4.2 (1.7–5.5)	4420	4.5e-2 (8.4e-4–2.2e-1)	5.3e-7 (5.6e-10–4.2e-6)	3.1e-4 (1.5e-4–1.0e-3)	133 (−60)	3.54 (−8.0)	49 (−18)	I
c	17157	51	1.3	4187	5.7e-4	4.9e-11	4.4e-4	20	8.0	49	I
e	17052	76 (−103)	5.1 (−1.2)	4772	3.9e-4 (−2.3e-2)	5.3e-10 (−5.2e-7)	1.0e-4 (0.8e-4–5.4e-4)	97 (−31)	5.51 (−8.0)	76 (−18)	I
2	17132	17 (−34)	2.8 (−6.4)	11428	4.2e-2 (7.2e-5–1.2e-1)	5.5e-7 (−1.5e-10)	0	79 (50–851)	8.0 (−1.96)	81 (−87)	II

Class I by KOE08. At least one of them (a3) is a 24 μm source. BRC14-a is likely associated with it. If true, only one model satisfies $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 . This model corresponds to an intermediate mass Stage I YSO. But the SED is not well constrained, perhaps because of the multiplicity of the central object.

- BRC14-b: A bright 24 μm source lies in this direction. It belongs to a small group of stars, observed at *Spitzer*-IRAC wavelengths; an object is clearly dominant at 8.0 μm and 24 μm . The SED is that of a Stage I YSO of intermediate mass. The parameters of the disk are not well constrained.

- BRC14-c: This source is faint and observed in the direction of a rather faint 24 μm source. This source was identified by KOE08 as a “deeply embedded protostar”. We confirm here that its SED is that of a Stage I YSO of intermediate mass.

- BRC14-d and -e: these sources, separated by some 5'', lie near the bright central IRS sources. They cannot be separated at SPIRE wavelengths. Only one 24 μm source is observed in the direction of BRC14-e; BRC14-d does not have any 24 μm counterpart and thus is a candidate Class 0 object. For BRC14-e, six models satisfy $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 . They point to a Stage I YSO of intermediate mass; but here again the parameters of the disk are not well constrained.

- BRC14-f: This source is observed in the direction of a Class II YSO according to KOE08. Because it lies near IRS1, its fluxes are difficult to measure. Its 24 μm counterpart has not been measured by KOE08, and our measurement at 100 μm is very uncertain. At *Spitzer*-IRAC wavelengths this YSO is the brightest source of a small group of stars. The SED is not well

constrained; the *Herschel* source is possibly not linked to the central YSO (but adjacent to it).

- BRC14-g, -h, -i: These sources lie near each other on the sky. BRC14-g and -h are possibly slightly extended. This region is complicated, as a small cluster is present in its direction. One IRAC Class II source is observed in the direction of BRC14-h (called h1 in Fig. B.3), which has no 24 μm counterpart. Two Class II sources are observed in the direction of BRC14-g (called g1 and g2 in Fig. B.3), and only g1 has a 24 μm counterpart. Another 24 μm source lies between -g and -h, that is not in KOE08 catalogue. It is difficult to associate -g and -h with any of these sources. For example, Fig. B.2 shows that if we associate BRC14-h to the Class II YSO h1 the SED is not well fitted at *Herschel* wavelengths. The situation is the same for BRC14-g. This and the fact that the *Herschel* sources are slightly extended suggest either that the cluster with its three Class II YSOs is embedded in (or lie adjacent to) the extended g and h condensations, or that it contains several Class 0 sources. BRC14-i has no IRAC counterpart, and has possibly a very faint 24 μm extended counterpart; it is a candidate Class 0 source.

- BRC14-j: This 100 μm source has no IRAC or 24 μm counterpart. It is a candidate Class 0 source.

Several other sources are Class I YSOs according to KOE08. YSOs #1 and #2 have very faint 100 μm counterparts, at the limit of detection (Fig. B.1). Their SEDs are not at all well constrained, as shown by Fig. B.2. However, the first 20 best fits for YSO #2 have all an envelope accretion rate equal to zero. Thus, YSO #2 is probably a Class II YSO, of intermediate mass; the parameters of its disk are not constrained. The SED of YSO #1

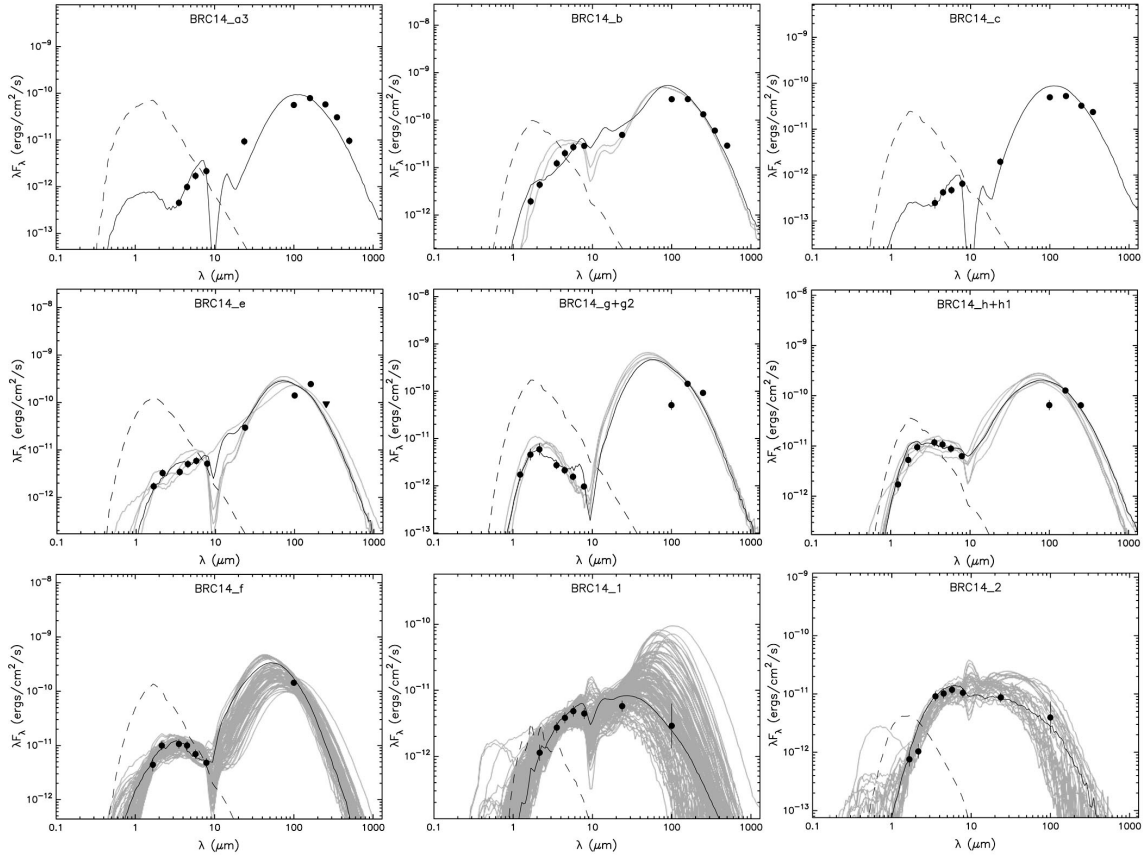


Fig. B.2. Spectral energy distributions of a few point sources in the BRC14 field. We used the SED fitting tool of ROB07.

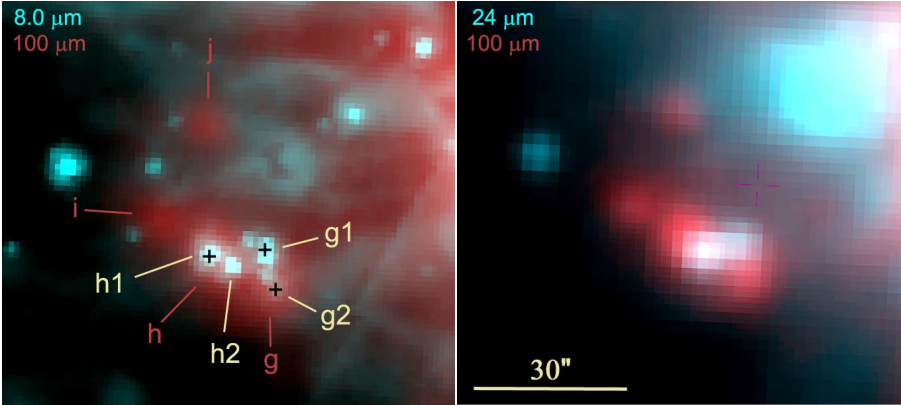


Fig. B.3. The BRC14-g, -h, -i, and -j 100 μm point sources and their associated mid-IR cluster. The black crosses show the position of the three Class II YSOs identified by KOE08.

resembles that of YSO #2; it is also probably a Class II YSO (in contrast to the finding of KOE08). The four other Class I YSOs (according to KOE08; #17122 close to BRC14-i, #17029, #17037, and #17076, close to the central IR sources) have no detectable or measurable 100 μm counterpart. We cannot estimate their evolutionary stage.

The cloud enclosed by BRC14 contains in its center two very bright *Herschel* sources, c1 and c2, observed in the direction of IRS1 and IRS2. They are discussed in the following section.

B.2. The dense condensation enclosed by BRC14

The molecular condensation enclosed by BRC14 is conspicuous in the *Herschel* images. It is the brightest and most massive condensation of the observed field (Fig. B.4). It belongs to the east-west filament extending toward Sh 201. It is symmetric with

respect to the direction of HD 18326. It contains the two bright IR sources IRS1 and IRS2.

The morphology of this condensation varies with the wavelength. At 100 μm it contains two slightly extended sources, observed in the direction of IRS1 and IRS2 (Fig. B.1); we call them c1 and c2 in Table B.1. The western one (IRS1) is the brightest (there is a factor of 3.5 in intensity between their peak emission values). At 160 μm , IRS2 becomes fainter; the ratio of the peak intensities is 2.85. At SPIRE wavelengths IRS2 is barely detected, IRS1 is dominant, but the peak emission is shifted toward the north-east; the column density peaks some 8'' (0.08 pc) north east of c1. The IR sources c1 and c2 are very bright 24 μm sources. They have been classified as Class I YSOs by KOE08, but no 24 μm measurements are given in the catalogue; c2 is the brightest at 24 μm , the contrary of what is observed at 100 μm and at longer wavelengths. This indicates that IRS2 is the hottest

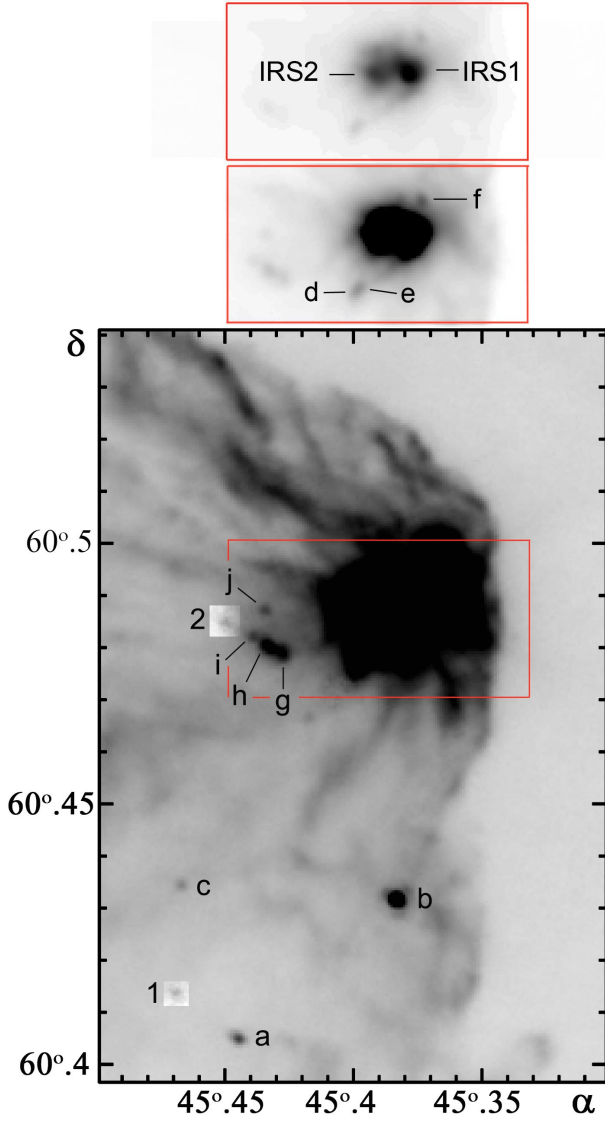


Fig. B.1. Identification of the YSOs in the BRC14 field. The grey-scale image shows the *Herschel*-PACS 100 μm emission. The three red boxes show the same zone, displayed with different intensity cuts to allow the identification of all the central sources.

of the two sources. This is expected as IRS2 is not a single source, but corresponds to dust heated by the cluster exciting a compact H II region (Deharveng et al. 1997).

This condensation has a peak column density $N(\text{H}_2) = 1.0 \times 10^{23} \text{ cm}^{-2}$, and a half-intensity diameter of 0.27 pc; this indicates a density $n(\text{H}_2) \sim 10^5 \text{ cm}^{-3}$. We have measured the flux of this condensation through circular apertures of radius 40'', 50'', and 60'', centered on the column density peak at $\alpha(2000) = 03^{\text{h}}01^{\text{m}}31^{\text{s}}79$, $\delta(2000) = +60^\circ 29' 21''$. We give these measurements in Table B.1, under the labels c^* , c^{**} , and c^{***} , respectively for the three apertures. The temperatures we derive from the *Herschel* fluxes are in the range 25.0 K to 26.4 K (the smallest aperture is the hottest), and the masses are in the range 170–250 M_\odot respectively for the smallest to the largest aperture (see table 5). This can be compared to the results obtained by Morgan et al. (2008) from SCUBA data; a temperature of 27 K, a diameter of 0.4 pc, and a mass of 94 M_\odot .

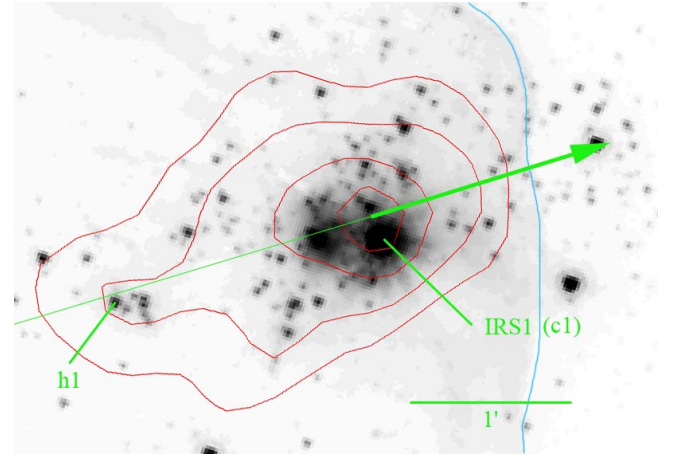


Fig. B.4. The bright-rimmed cloud BRC14. The underlying greyscale image is the *Spitzer* 3.6 μm emission showing the cluster AFGL4029 (IRS2) and the YSOs BRC14-c1 (IRS1; a Class I source according to KOE08)). The red contours correspond to column densities of 0.5, 1, 2.5, and $5 \times 10^{22} \text{ cm}^{-2}$. The peak column density is $\sim 1.0 \times 10^{23} \text{ cm}^{-2}$. The arrow points to the exciting star HD 18326. The blue line shows the position of the IF. North is up and east is left.

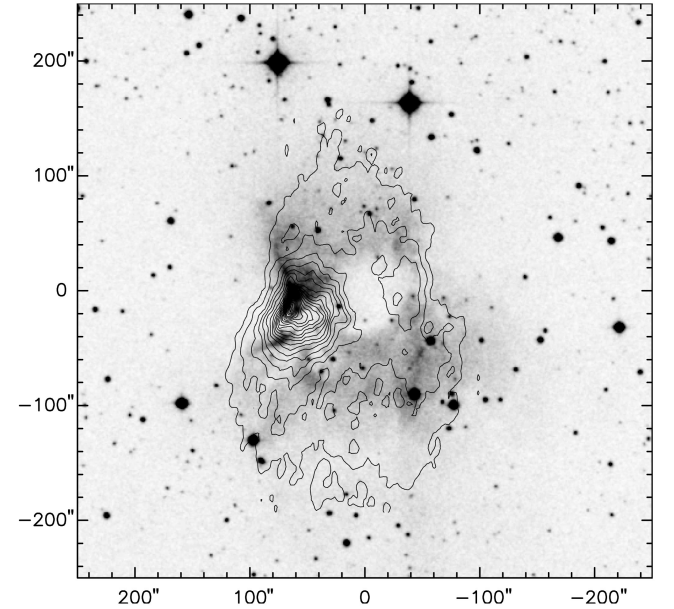


Fig. C.1. Emission of the ionized gas in Sh 201. The contours of the radio-continuum emission, as given by Ojha et al. (2004; their figure 6), are superimposed on a grey image of the H α emission (DSS-2 red image). The brightest emission comes from the dense ionized layer bordering a molecular condensation. North is up and east is left.

for integration in a 30'' aperture (if corrected for a distance of 2 kpc and an opacity at 850 μm of $1.25 \text{ cm}^2 \text{ g}^{-1}$).

Appendix C: The Sh 201 bipolar H II region and vicinity

C.1. Presentation of the region

Sh 201 is a small H II region located to the east of W5-E, about 7.6 pc away from the IF bordering BRC14. It is coincident with

the IRAS source IRAS 02593+6016, which has a luminosity of $\sim 1.1 \times 10^4 L_{\odot}$. It probably belongs to the same complex as W5-E since its velocity ($V(\text{RRLs}) = -34.6 \text{ km s}^{-1}$, Lockman 1989; $V(\text{H}\alpha) = -35.5 \text{ km s}^{-1}$, Fich 1990) is very similar to that of W5-E ($V(\text{H}158\alpha) = -37.0 \text{ km s}^{-1}$, Dieter 1967; $V(\text{H}\alpha) = -39.0 \text{ km s}^{-1}$; $V(\text{H}\alpha) = -34.3 \text{ km s}^{-1}$, Mampaso et al. 1987).

Sh 201 is a bipolar optical H II region of low excitation and low density, except for a bright H α knot on the east side (Mampaso et al. 1987 and Fig. C.1). These authors suggest that this bright optical structure results from the interaction of the ionized gas with an adjacent dense molecular cloud. Sh 201 is a thermal radio source with an integrated flux density of $1.2 \pm 0.2 \text{ Jy}$ at 6-cm (Felli et al. 1987 and references therein). At a distance of 2 kpc this flux density indicates an ionizing photons flux of $4.25 \times 10^{47} \text{ s}^{-1}$ (using equation 1 in Simpson & Rubin 1990). This ionized flux points to an ionizing star of spectral type O9–O9.5V, if single, according to the calibration of Martins et al. (2005). Felli et al. (1987), Fich (1993), and Ojha et al. (2004) all show high-resolution radio continuum maps of Sh 201. These maps, obtained at different frequencies, are very similar to one another; this confirms the thermal nature of the radio-continuum emission. The radio source has a bright arc shaped edge on the east side, and a smoothly decreasing surface brightness on the other side (Fig. C.1). All aforementioned authors interpret this arc-shaped structure as a result of the interaction between the H II region and an adjacent molecular cloud (on the east side). Felli et al. (1987) have shown that an O9 ZAMS star (the exciting star of the H II region) placed at a distance of 0.26 pc (for $d=2 \text{ kpc}$) of the border of the cloud is consistent with the available data. This model does not attempt to explain the bipolar nature of Sh 201. Fig. C.1 summarises the situation for the ionized gas: 1) the regions of radio and H α emission have a similar extent, except for the central absorption zone; 2) the arc-shaped bright radio feature corresponds to the bright H α knot; they are bordering a region of high extinction to the east of Sh 201; 3) a faint secondary radio component is present about 1 pc west of the brightest radio component (on the opposite side at the waist of the bipolar nebula).

Ojha et al. (2004) discuss the stellar content of Sh 201, based on near-IR observations (see also Carpenter et al. 1993). A cluster, containing more than a hundred stars, is present in the direction of IRAS 02593+6016; it contains YSOs identified by their near-IR excess. The brightest and most reddened source is, according to these authors, an O6–O8 type star. Two massive stars with spectral type probably earlier than B2 are also present, which display a near-IR excess. These stars are identified on Fig. C.2; they are respectively #2 (the exciting star), #1 and #3.

The molecular content of the region has been the subject of various studies. Martin and Barrett (1978) have detected and mapped the ^{13}CO emission of a flat molecular cloud, located at the waist of the bipolar nebula, that is probably responsible for the optical absorption observed there (see also figure 3 in Felli et al. 1987). This cloud extends east-west, with a size of more than $18 \text{ pc} \times 6 \text{ pc}$, and has a velocity similar to that of the H II region ($V(\text{CO}) = -39.1 \text{ km s}^{-1}$, Blair et al. 1975; $V(\text{CO}) = -40 \text{ km s}^{-1}$, Blitz and Fich 1982; $V(\text{CO}) = -38 \text{ km s}^{-1}$, Wootten and Habing 1985). Large-scale CO maps show the extent of this filament which goes from BRC14, through Sh 201, to eastern regions far away from Sh 201; its total east-west extent is $\sim 20 \text{ pc}$ (Carpenter et al. 2000; Karr and Martin 2003; Niwa et al. 2009, their cloud 7). This filament has two condensations on each side of Sh 201 (clumps 8 and 9 in Niwa et al. 2009). The CS map of Carpenter et al. (1993) shows that the east clump is by far

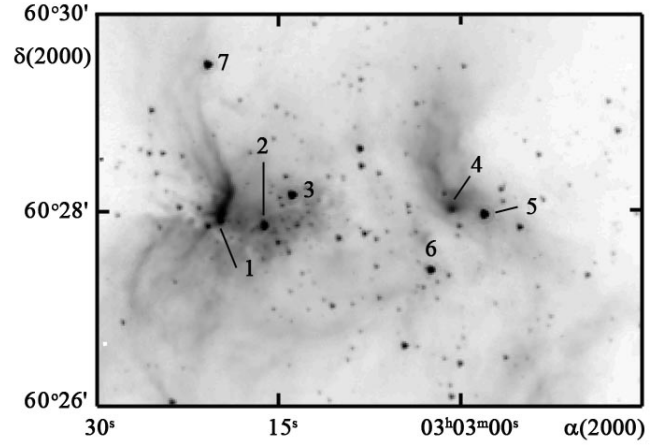


Fig. C.2. Young stellar objects and stars located at the waist of Sh 201. They are identified on the *Spitzer*-IRAC 4.5 μm image. Star #2 is the exciting star.

the brightest. It is also the densest of all the CO clumps located in the periphery of W5, with a density of $2.6 \times 10^4 \text{ cm}^{-3}$ (Niwa et al. 2009). It is the dense ionized layer bordering this clump that appears as the bright arc-shaped structure in the radio-continuum map and in H α .

The Sh 201 region has been observed by *Spitzer* (KOE08). The 8.0 μm IRAC image, dominated by the PAH emission bands from the PDR, gives the best illustration of the bipolar nature of Sh 201. Fig. C.3 shows the column density contour levels (obtained by *Herschel*) superimposed on the 8.0 μm image. It confirms: 1) the presence of a dense neutral filament extending east-west, with two bright condensations at the waist of the bipolar nebula and a hole of emission in the central region (occupied by ionized material); 2) the bipolar nature of Sh 201: two lobes extend perpendicular to the filament, up to 3.5 pc from the central exciting cluster.

Figure C.4 displays the distribution and temperature of the dust observed by *Herschel*. Red is 500 μm band that traces the emission of cold dust. As such, this channel shows the numerous cold filaments containing most of the mass, especially the east-west filament mentioned earlier that extends from BRC14 to Sh 201. Green is the 100 μm band, which traces the emission of warmer dust along the PDRs of the ionized regions. We see very clearly the two lobes perpendicular to the east-west filament, and a brighter emission at the waist of the bipolar nebula. Blue is the *Spitzer* 24 μm band, which traces even hotter dust. Its emission is located predominantly either in the PDRs or in the ionized region; we find 24 μm extended emission in the direction of the “interior” of W5-E and in the central region of Sh 201, close to the exciting cluster. Several cold dust condensations (red colour) are found along the east-west filament, one enclosed by BRC14, two on each sides of the waist of Sh 201 (the eastern one the brightest), and a small one at the extreme east, $\sim 5.5 \text{ pc}$ from Sh 201 (called filament #1 in Sect. 6.6 and Fig. 15). The eastern condensation at the waist of Sh 201 has a peak column density of $6.2 \times 10^{22} \text{ cm}^{-2}$, corresponding to a visual extinction of 66 mag. The mass of this condensation, integrated over an area limited by the contour of $0.6 \times 10^{22} \text{ cm}^{-2}$ (10% of the peak intensity) is $235 M_{\odot}$. This indicates a mean density of $3.5 \times 10^4 \text{ H}_2 \text{ cm}^{-3}$, in rather good agreement with the results of Niwa et al. (2009); they find from C^{18}O observations of their clump 9 (corresponding to our eastern condensation) a col-

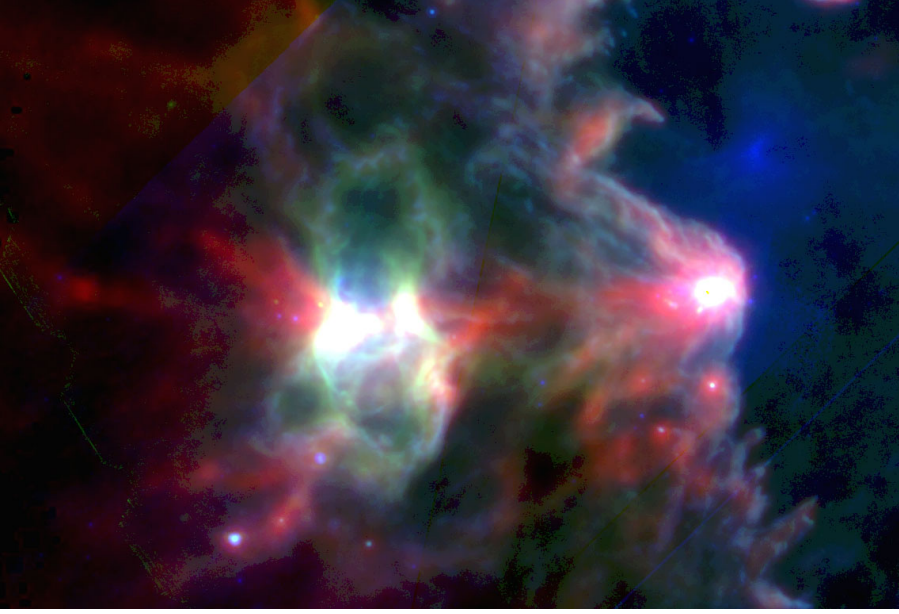


Fig. C.4. The vicinity of Sh 201 and BRC14. Composite colour image with the *Herschel*-SPIRE 500 μm image in red, the *Herschel*-PACS 100 μm image in green and the *Spitzer*-MIPS 24 μm image in blue (all in logarithmic units). Red is the emission of the cold dust; it shows the dense material in this region. Green traces the emission from warmer dust in the PDRs of the H II regions. Blue traces the emission from even hotter dust, located in the PDRs but also inside the ionized regions. YSOs appear as point sources at 24 μm and 100 μm . The field is $32.5' \times 21.8'$. North is up and east is left.

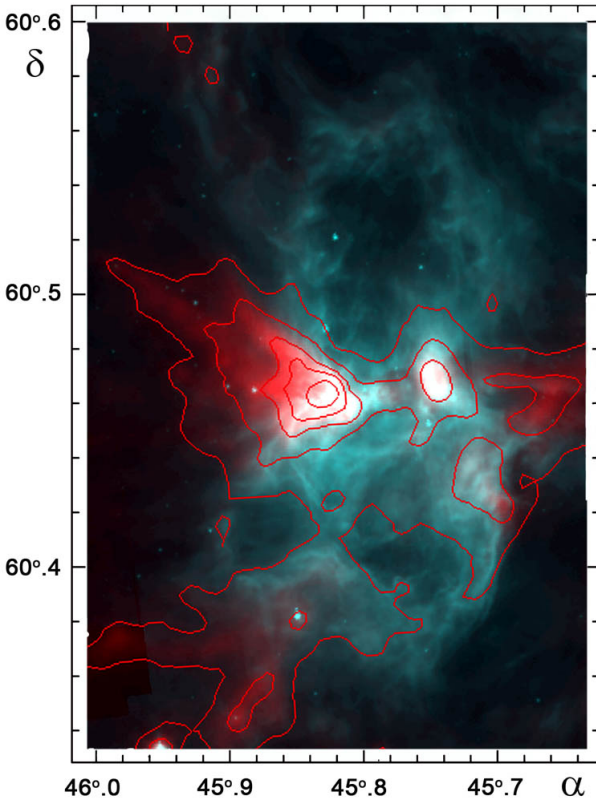


Fig. C.3. The bipolar nature of Sh 201. Red is for the column density map; the red contours are of column density; the peak column density is $6.2 \times 10^{22} \text{ cm}^{-2}$ (east condensation) and the contours levels are 0.1, 0.25, 0.5, 1, 2.5, and $5 \times 10^{22} \text{ cm}^{-2}$. Turquoise is for the *Spitzer* 8.0 μm image showing the PAH emission from the vicinity of the ionization front. The coordinates are J2000 in degrees.

umn density of $4.3 \times 10^{22} \text{ cm}^{-2}$, and a density of $2.6 \times 10^4 \text{ cm}^{-3}$. If we consider the contour at half-intensity as the integration limit, we measure a mass of $70 M_{\odot}$ for a size (diameter) of 0.25 pc. This indicates a core density of the order of $9 \times 10^4 \text{ H}_2 \text{ cm}^{-3}$, in

good agreement with the results of Carpenter et al. (1993) who showed with their CS (2-1) observations that this condensation has a dense core.

The western condensation at the waist of the bipolar nebula has a smaller column density of $1.1 \times 10^{22} \text{ cm}^{-2}$ at the emission peak. It is slightly elongated north-south along the ionization front, with an equivalent diameter of 0.36 pc at half intensity; thus a smaller density of the order of 10^4 cm^{-3} . Note that this condensation is not clump 8 in Niwa et al. (2009).

C.2. Young stellar objects in the vicinity of Sh 201

Eight point sources, identified in Fig. C.5, are detected at 100 μm around Sh 201. Five of these point sources have *Spitzer* counterparts in the catalogue of KOE08. Of these, four are Class I sources and one is Class II. Six of the point sources have 24 μm counterparts that are not included in KOE08 catalogue. We have measured the 24 μm fluxes of these six sources, as well as the *Herschel* fluxes. We give these data in Table C.1 (the values with an asterisk are our 24 μm measurements).

We fit the SEDs using the SED fitting-tool of ROB07, assuming an interstellar (external) visual extinction in the range 1.96 mag – 5 mag; this last value is the mean extinction found for the whole associated cluster by Carpenter et al. (1993). We list the parameters of the YSOs, as well as the external extinction for the best models, in Table C.2. We show a few SEDs in Fig. C.6, and comment on individual sources below:

- S201-a: this strong source is located in the direction of a bright filament, south of Sh 201, and far from it. We show in Sect. 6.4 that, based on velocity arguments, this source is not associated with the W5 region, but lies far in the background at about 5.45 kpc. (The parameters given in Table C.2 are obtained for this distance, allowing an external extinction free in the range 1.95 mag to 10.0 mag, and not taking into account the saturated 24 μm emission.) Twelve models satisfy $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 ; all point to a massive Stage I YSO. Based on its relatively high temperature (Sect. 6.3.1) and luminosity we suggest that this source contains a small cluster, and is similar to BRC13-a or BRC14-c1+c2.

- S201-b: this source is located to the south of Sh 201, in the direction of a low column density region. The *Spitzer*-IRAC

Table C.1. Photometry of point sources in the vicinity of Sh 201. The asterisks indicate our 24 μm measurements.

Name	#	RA(2000) (°)	Dec(2000) (°)	<i>J</i> mag	<i>H</i> mag	<i>K</i> mag	[3.6] mag	[4.5] mag	[5.8] mag	[8.0] mag	[24] mag	Class	S(100) Jy	S(160) Jy	S(250) Jy	S(350) Jy	S(500) Jy
a	17719	45.958974	60.336970	14.86	13.50	11.61	8.94	7.99	6.55	5.24	0.97*	I	30.00	40.00	9.32	3.91	2.48
b	17549	45.797180	60.335376	16.20	14.87	14.15	12.17	11.68	9.46	7.87	5.32*	I	1.64	2.85	1.62	0.66	
c	17660	45.903246	60.348649				14.07	12.22	10.77	9.85	5.76*	I	1.14	2.62	4.06	2.95:	2.31:
d	17609	45.857765	60.385978	11.67	10.06	8.83	7.26	6.70	6.20	5.25	1.81*	II	1.66	3.75	2.90:	1.65	1.27:
e	17644	45.889840	60.468891	15.39	13.30	11.82	9.69	8.88	8.17	7.24	3.81*	I	0.55	1.01	1.2:		
f		45.6981	60.4590								6.12*		1.03	3.16	5.19:		
g		45.8594	60.4770								no		5.34	15.44	17.24	11.02:	
h		45.9134	60.4785								no		0.79	1.44	1.21	1.39	
5	17484	45.741884	60.465859	13.75	11.84	10.40	9.01	8.23	7.36	5.95	1.86*	I					
6	17506	45.760247	60.456447	11.71	11.20	10.85	10.00	9.09	7.96	6.20	2.01*	I					
7	17595	45.836823	60.491218	14.23	13.05	11.85	9.72	8.83	7.89	6.15	2.46*	I					
s1		45.83351	60.46711										80				
s2		45.83082	60.46896										36				

Table C.2. Characteristics of the YSOs in the vicinity of Sh 201

Name	χ^2	M_* (M_\odot)	T_* (K)	\dot{M}_{disk} (M_\odot)	\dot{M}_{disk} ($M_\odot \text{ yr}^{-1}$)	\dot{M}_{env} ($M_\odot \text{ yr}^{-1}$)	L (L_\odot)	$A_V(\text{ext})$ (mag)	i (°)	Stage
a (5.45kpc)	74 (−110)	9.6 (5.1–17.8)	4450 (−9500)	1.1e-1 (3.3e-3–9.3e-1)	4.4e-6 (3.2e-9–1.7e-4)	1.6e-3 (1.9e-4–3.0e-3)	1540 (1060–9300)	9.5 (1.95–10.0)	18 (−63)	I
b	72 (−108)	0.7 (0.5–1.0)	3800	8.6e-4 (5.0e-4–6.9e-2)	5.8e-7 (4.2e-8–1.6e-6)	8.4e-5 (−1.2e-4)	14 (12–17)	2.1 (−4.9)	18	I
c	104 (−137)	1.4	4218	1.1e-2	1.1e-7	4.7e-4	21	1.96	49	I
d	121 (−160)	4.7 (2.9–5.7)	13170	1.6e-4 (−1.8e-3)	5.5e-9 (−6.4e-10)	0 (1.7e-7)	560 (−740)	3.53 (−5.0)	18 (−81)	II
e	55 (−88)	1.9 (0.2–4.7)	4317 (2970–15800)	1.4e-1 (−1.2e-5)	9.1e-8 (1.5e-10–6.4e-6)	1.7e-5 (−0)	34 (11–410)	3.62 (−5.0)	69 (18–87)	?
5	16 (−40)	0.4 (−5.4)	3512 (−17760)	3.9e-2 (−1.9e-4)	2.0e-5 (−7.4e-11)	3.0e-6 (−0)	42 (−730)	5.0 (−3.68)	18 (87)	II?
6	48 (−72)	7.4 (−3.6)	21340 (−5810)	3.5e-8 (−1.4e-3)	1.7e-13 (−3.2e-9)	0 (−1.9e-7)	2070 (−145)	4.0 (−1.96)	69 (32–87)	II?
7	18 (−42)	1.2 (−3.4)	4034 (−12800)	3.4e-4 (−3.5e-2)	1.0e-8 (−5.1e-10)	6.4e-6 (−0)	30 (−110)	3.17 (−5.0)	18 (−81)	II?

images show a small group of stars in this direction; one is dominant at 8.0 μm and at longer wavelengths. It is a rather low mass Stage I YSO.

- S201-c: the fit for this source is not very good, and only one model satisfies $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 . The model points to a Stage I YSO.

- S201-d: this source is possibly located at the intersection of two filaments, to the south of Sh 201. This location is also a region of interaction between these filaments and the southern lobe of Sh 201. Many models satisfy $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 ; they are all rather similar, but the fit is not good. S201-d is probably an intermediate mass Stage II YSO (the envelope accretion rate is low, null in some models), in agreement with KOE08. Fig. C.5 shows an extended structure, jet-like, which seems to originate from the YSO (it is only observed at 100 μm and 160 μm); CO or OH lines, present in the PACS bands, and observed in regions of jets and shocks (van Kempen et al. 2010) could perhaps explain this feature. This feature, however, is not observed at 4.5 μm , a *Spitzer* band which traces shocks (Cyganowski et al. 2008, 2009).

- S201-e: this source is located in the direction of the massive condensation on the east side, at the waist of Sh 201. Its SED is not constrained at all. The best model corresponds to a Stage I YSO, of intermediate mass. But numerous models satisfy $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 , which span a wide range of parameters. Thus, the nature of S201-e is very uncertain (from a low mass Stage I YSO to a high mass Stage II/III YSO).

The three other compact sources (f, g, h) have only *Herschel* measurements. Source f has a faint counterpart at all *Spitzer* wavelengths but it has not been measured by KOE08. Source g, which lacks a 24 μm counterpart and which is relatively bright at *Herschel* wavelengths, is probably a Class 0 source (Sect. 6.3.3). It is observed in the direction of the filaments bordering the massive condensation east of Sh 201. This location makes it a very good candidate for triggered star formation near the waist of

Sh 201. We can estimate the mass of the g and h sources using the *Herschel* fluxes (Table 5); these masses are respectively $\sim 46 M_\odot$ and $\sim 4 M_\odot$. We hypothesize that sources g and h are Class 0 sources.

The bright *Spitzer*-IRAC sources #1, #2, #3, and #4, located at the waist of Sh 201 and identified in Fig. C.2, have been classified as Class II sources by KOE08. According to Ojha et al. (2004), source #2 is the main exciting star of Sh 201, whereas sources #1 and #3 are early B stars. Source #4 is associated with bright extended emission at 24 μm , as is source #2; we suggest that it is also an early B star. Sources #5, #6, and #7 are Class I sources according to KOE08. They are bright 24 μm sources (no fluxes were given by KOE08; we measured these fluxes and give them with asterisks in Table C.1). None of these sources are detected at 100 μm . Their SEDs are not well-constrained. Most of the models for YSOs #5 and #7 indicate Stage II YSOs: their envelope accretion rates are low, even zero in some models, and their disks' masses are high (#5 and #7 only). But the mass and age of the central objects are not at all constrained. Four other sources have been classified as Class I in the field of Fig. C.2, which have no 100 μm counterparts. They are #17667, #17531, #17526, #17620. Except for #17667, they also have very faint or no 24 μm counterparts. Thus, their classification as Class I is doubtful.

Two other sources detected by *Herschel*, s1 and s2, lie at the eastern part of the waist of Sh 201, between the bright PAH-emitting filaments bordering the central ionized region and the massive molecular condensation. They are observed at 100 μm and 160 μm , but are not detected at shorter wavelengths and are not easily discernable at longer wavelengths, probably due to the lower resolution and the presence of the adjacent parental molecular condensation. In Fig. C.7 we show an unsharp-mask image at 100 μm of the field containing these sources. To create the unsharp-mask, we median filtered the PACS 100 μm image with a square window of 25'' and subtracted the filtered im-

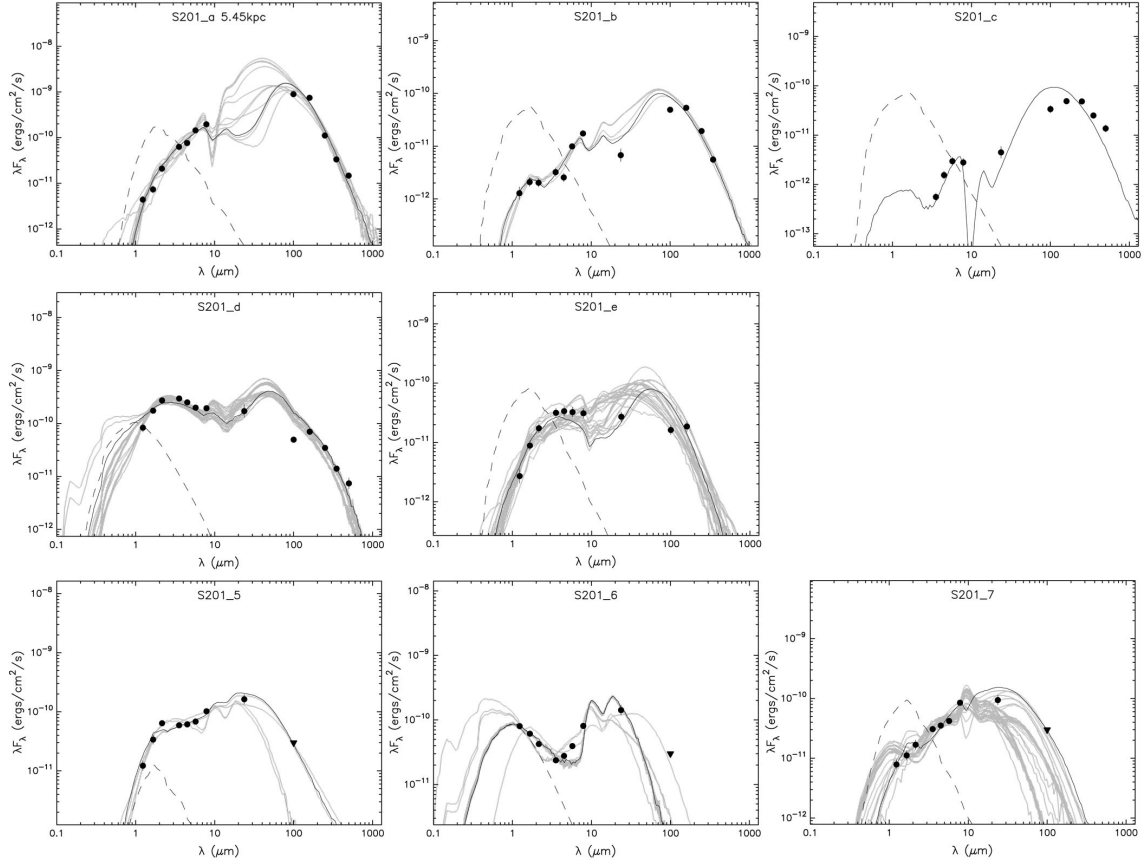


Fig. C.6. Spectral energy distributions of point sources in the Sh 201 field. We used the SED fitting-tool of ROB07.

age from the original one. Fig. C.7 (top) shows the two point sources, of which s1 is the brightest, and Fig. C.7 (middle) shows the 100 μm filaments in the PDR. The 100 μm flux of s1 and s2 are uncertain by a factor 2 because of their proximity to one another and because of the bright underlying filaments. s1 lies at $\alpha(2000) = 03^{\text{h}}03^{\text{m}}20.08^{\text{s}}$, $\delta(2000) = +60^{\circ}28'01''$ and has a 100 μm integrated flux ~ 80 Jy. s2 lies at $\alpha(2000) = 03^{\text{h}}03^{\text{m}}19.35^{\text{s}}$, $\delta(2000) = +60^{\circ}28'09''$ and has a 100 μm integrated flux ~ 36 Jy. Using the 24 μm detection limit ~ 7 mag, we obtain a spectral index in the range 24 μm –100 μm $\alpha \geq 7.2$ for s1 and ≥ 6.65 for s2. These very high indexes indicate that these sources are dominated by their envelopes. These sources are compact and their 100 μm fluxes are large. Thus, they are good candidates for high-mass Class 0 sources. Additionally, their location makes them good candidates for triggered star formation. An H_2O maser has been detected at 22 GHz by Blair et al. (1980) in the direction of star #1 (Fig. C.7); however, due to the resolution of these observations ($\sim 1.4''$), it is unclear if it is associated with #1, s1 or s2.

Appendix D: The region between W5-E and W5-W: W5-NE

Molecular material with a rather high column density of $\sim 5 \times 10^{21} \text{ cm}^{-2}$ is present between W5-E and W5-W (see Fig. 4). Two small H II regions are located here, surrounded by dense material. A bright rim is also present, turned towards W5-E and its exciting star. Several point sources, listed in Table D.1, are detected at 100 μm and at longer wavelengths.

D.1. The H II regions

At least two H II regions are present in the region located between W5-E and W5-W. In Fig. D.1 we show the $\text{H}\alpha$ emission of the ionized gas. The exciting star of the A H II region is unknown; we identify it with an optical star at the center of the optical and IR nebula that we call ex1 (see Fig. D.1 and Fig. D.2). Wilking et al. (1984) call this IR source IR2, and, using several indicators, propose a spectral type in the range B0–B2 ZAMS for the central star. The B H II region is excited by BD+60°596, a B1V star (Hiltner 1956) which lies at $\alpha(2000) = 02^{\text{h}}57^{\text{m}}08^{\text{s}}02$, $\delta(2000) = +60^{\circ}39'44''.1$ (a few other spectral types have been proposed for this star: B3III, Garmany & Stencel 1992; B0.5V–III, Zdanavicius & Zdanavicius 2001).

The NVSS survey at 20-cm shows radio-continuum emission from three regions, the A and B H II regions and the border of a bright rim (Fig. D.1; two components are listed in the NVSS catalogue for the B H II region). Fig. D.2 shows in green the *Spitzer*-IRAC 5.8 μm extended emission coming mostly from the PDR of the large W5-E and the PDRs surrounding the small H II regions. Fig. D.2 also shows in red the *Spitzer*-MIPS 24 μm extended emission due to the hot dust located inside the ionized regions; it allows to locate the exciting stars. A 24 μm extended emission appears north of the A H II region. It probably points to the presence of another early B star (there is an extension of the radio contours in this direction). An optical and IR star, that we call ex2, lies at the center of this extended emission.

The three stars (ex1, ex2, and BD+60°596) are in the catalogue of YSOs by KOE08. They are listed in Table D.1; ex2 and BD+60°596 are classified as Class III by KOE08; ex1 is classified as Class II. This classification is uncertain as these

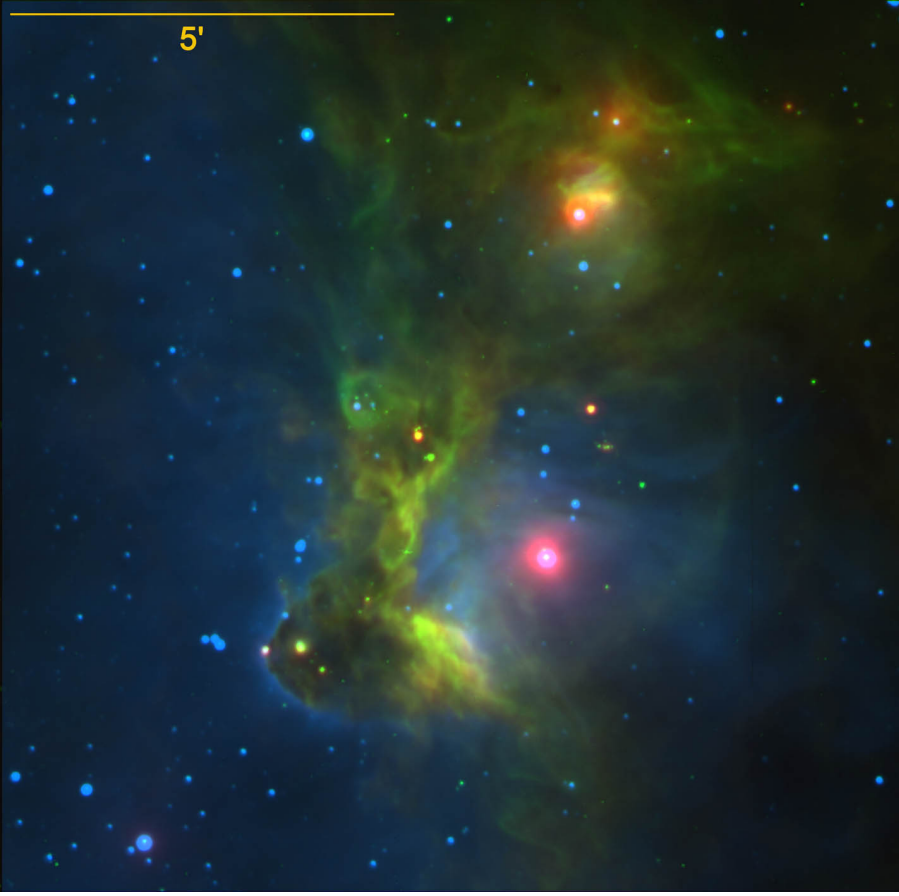


Fig. D.2. Colour composite image of the region located between W5-E and W5-W. Red is the *Spitzer*-MIPS emission from the hot dust at $24\ \mu\text{m}$, green is the *Spitzer*-IRAC emission at $5.8\ \mu\text{m}$ from the YSOs and the PDRs (PAH emission), blue is the $\text{H}\alpha$ emission of the ionized gas.

stars are at the center of bright $24\ \mu\text{m}$ extended emission. We can use the K magnitude and $J - K$ colour to estimate (a very coarse estimate) the spectral type of these three stars. Using a distance of 2 kpc, the 2MASS photometry of ex1 indicates a B1–B2 spectral type with a visual extinction ~ 3.4 mag. Using the NVSS flux density of 28.8 mJy for the A H II region, we estimate a ionizing photon flux $\log(N_{\text{Ly}\alpha}) = 45.9$ pointing to a spectral type B2V (Smith et al. 2002); if we consider the CGPS integrated flux of 0.13 Jy (Taylor et al. 2003) we obtain $\log(N_{\text{Ly}\alpha}) = 46.4$ ionizing photons per second, corresponding to a B1V spectral type (Smith et al. 2002). These results are in agreement with Wilking et al. (1984) spectral type. The 2MASS photometry of ex2 indicates a B1–B2 spectral type but with a higher visual extinction ~ 6.9 mag. The 2MASS photometry of BD+60°596 indicates a spectral type in the range B0V–B1V, with a low visual extinction ~ 1.4 mag. The integrated CGPS flux density of 0.31 Jy gives $\log(N_{\text{Ly}\alpha}) = 46.9$ pointing to a B0.5V spectral type (Smith et al. 2002); these results are in good agreement with the previous determinations.

D.2. The YSOs

Figure D.3 shows the *Herschel*-PACS $100\ \mu\text{m}$ image of the field. The colour image allows us to compare the location of the hot dust emitting at $24\ \mu\text{m}$ with the dust emitting at $100\ \mu\text{m}$. The bright extended $24\ \mu\text{m}$ emission comes from the inside of the H II regions and from the PDRs bordering them, whereas the extended $100\ \mu\text{m}$ emission comes only from the PDRs. The A and B H II regions are rather different. The B H II region is bordered by a dense PDR on its east side and seems to open on the west

side. The low extinction of its exciting star indicates that it lies in front of most of the material along this line of sight. The A H II region is more compact, but its influence seems to extend over a large region. The extinction of its exciting star is larger than that of BD+60°596, indicating that it is more embedded. A bright feature at $100\ \mu\text{m}$, bar-like, lies very close in projection to the ex1 exciting star.

Seven $100\ \mu\text{m}$ point sources are present in this field, identified in Fig. D.3, and listed in Table D.1. We fit some SEDs using the SED fitting-tool of ROB07 (assuming an external visual extinction of 1.96 mag). The obtained parameters are given in Table D.2. We comment on individual point sources below (the name IB refers to “in between”).

- IB-a: This source is the brightest of a group of possibly three $100\ \mu\text{m}$ sources. A detailed view of this region is given in Fig. D.4. It has a very faint $24\ \mu\text{m}$ counterpart, and thus is a candidate Class 0 source. Using only the *Herschel* fluxes we obtain for the envelope of this source a dust temperature of 13.7 K and a mass of $16.5\ M_{\odot}$ (Table 5). The Class I YSO a1 in KOE08 catalogue has a faint $100\ \mu\text{m}$ counterpart, too faint to be measured accurately. The western $100\ \mu\text{m}$ source has no $24\ \mu\text{m}$ counterpart at all; it is also a candidate Class 0 YSO.

- IB-b: This source is a Class I YSO according to KOE08. We show its SED in Fig. D.5. Three models satisfy $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 . The SED of this source points to a Stage I YSO, with a central object of about one solar mass, an envelope of $8.7\ M_{\odot}$ (Table 5), and a rather high accretion rate.

- IB-c: This source is a Class II YSO according to KOE08; however a far-IR counterpart appears to be present on all *Herschel* images. The SED is not well fit as shown by Fig. D.5; however it points to a Stage I YSO. We cannot exclude the exis-

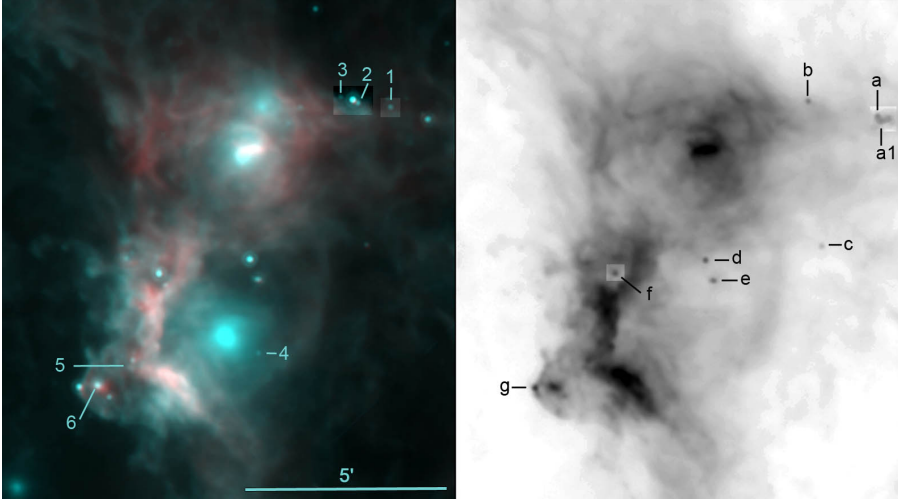


Fig. D.3. Colour composite image of the region between W5-E and W5-W. *Left:* Red is the *Herschel*-PACS emission at $100\ \mu\text{m}$, and turquoise is the *Spitzer*-MIPS emission at $24\ \mu\text{m}$. *Right:* The point sources detected at $100\ \mu\text{m}$ and discussed in the text are identified on a grey-scale image of the $100\ \mu\text{m}$ emission (given in logarithmic units).

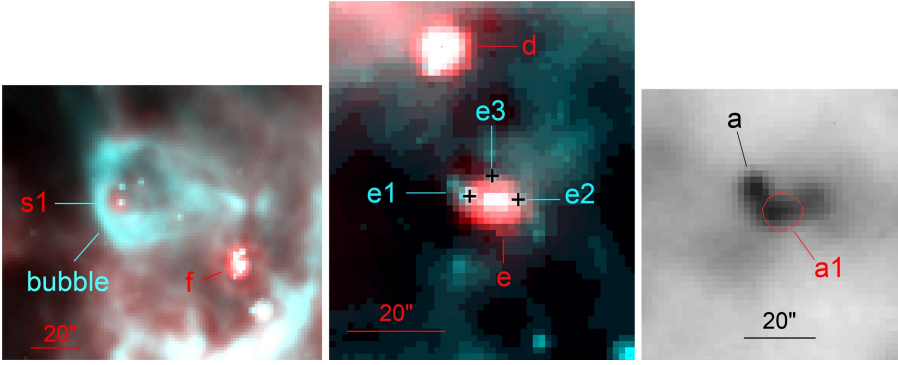


Fig. D.4. Features discussed in the text. *Left:* Colour composite image of a small bubble; red is the $24\ \mu\text{m}$ image, turquoise is the *Spitzer* $8.0\ \mu\text{m}$ image showing the PAH emission. *Middle:* YSO IB-e and its vicinity; red is the $100\ \mu\text{m}$ emission, turquoise is $8.0\ \mu\text{m}$ emission. *Right:* YSO-a; the underlying grey image is the $100\ \mu\text{m}$ emission, the red contour shows the position of the $24\ \mu\text{m}$ Class I source a1.

tence of two distinct sources: a Class II YSO adjacent to a condensation of $8.5\ M_{\odot}$ (Table 5).

- IB-d: This source is a Class I according to KOE08. We agree with this classification, but again the SED fit is not good. The parameters of the model give a Stage I source with a massive envelope ($20.3\ M_{\odot}$; Table 5), and a high envelope accretion rate. We did try the fits with the extinction as a free parameter; the fit is better (according to χ^2) with a higher external extinction (10.38 mag), but the fitted model is very similar to the previous ones.

- IB-e: This $100\ \mu\text{m}$ source is slightly elongated in the east-west direction. It is a complicated region as a small group of about half a dozen stars lies in its direction (see Fig. D.4). Only three YSOs (e1, e2, and e3) are in KOE08 catalogue. None of them are associated with the $100\ \mu\text{m}$ source, as shown by Fig. D.4. (The $24\ \mu\text{m}$ source resembles the $100\ \mu\text{m}$ source; it is elongated east-west). Two sources appear at $8.0\ \mu\text{m}$ in the direction of IB-e (which probably explains its elongation); they are not present in the catalogue of KOE08. Source e is probably associated with these two sources, which thus probably are Class I YSOs.

- IB-f: This source has been identified as a Class II YSO by KOE08. Two other faint stars lie nearby, but f is dominant at all *Spitzer* wavelengths. The SED is not well fitted by the SED fitting-tool of ROB07, as shown by Fig. D.5. This source is V* LW Cas, a variable star of uncertain nature (Kukarkin et al. 1971), not a cataclysmic variable (Downes & Shara 1993), not a symbiotic star (Henden & Munari 2008). It is probably evolved, and we do not know if it is associated or not with W5-E.

- IB-g: This source is a Class II YSO according to KOE08. It has a $100\ \mu\text{m}$ counterpart, but nothing is detected at longer

wavelengths. Numerous models have an acceptable fit to the SED. The model parameters point to an intermediate mass YSO, probably in Stage I; but the parameters of the disk are not well constrained.

The field of Fig. D.3 contains six other Class I YSOs according to KOE08; they are identified in this figure. YSOs a, a1, b, #1, #2, and #3 are part of a high column density ridge extending east-west; ex2 is possibly part of this ridge. This structure therefore appears to be an active star forming region, but its origin is unclear. YSOs #1 and #2 have a very faint counterpart at $100\ \mu\text{m}$ (too faint to be measured). YSO #4 is observed in the direction of the B H II region; it has a very faint counterpart at $100\ \mu\text{m}$. YSO #5 has no $100\ \mu\text{m}$ counterpart. It is faint at $24\ \mu\text{m}$ and therefore its classification by KOE08 as a Class I is uncertain. YSO #6 is peculiar. It has an extended and bright $100\ \mu\text{m}$ counterpart, located on the border of a bright condensation. A close examination of the *Spitzer* images shows that this source is also extended at all *Spitzer* wavelengths. This condensation, centered at $\alpha(2000)=02^{\text{h}}57^{\text{m}}33^{\text{s}}02$, $\delta(2000)=+60^{\circ}27'32''$ (position of the peak column density, $6.1 \times 10^{21}\ \text{cm}^{-2}$) has an integrated flux at $250\ \mu\text{m} \sim 10.5\ \text{Jy}$ in an aperture of radius $20''$. The temperature map indicates a mean temperature $\sim 23\ \text{K}$ for this extended structure and we estimate its mass to be $\sim 6\ M_{\odot}$.

Another peculiar structure is present in the field. It is a small bubble (diameter $\sim 35''$, thus $\sim 0.3\ \text{pc}$), well traced by the PAH emission at $8.0\ \mu\text{m}$ (see Fig. D.4). It contains three YSOs, classified as Class II by KOE08. One of them is especially bright at $24\ \mu\text{m}$ and possibly extended; it has no $100\ \mu\text{m}$ counterpart. We propose that this source, identified as s1 in Table D.1 and Fig. D.4, is a B star heating the nearby dust. (Its *K* magnitude

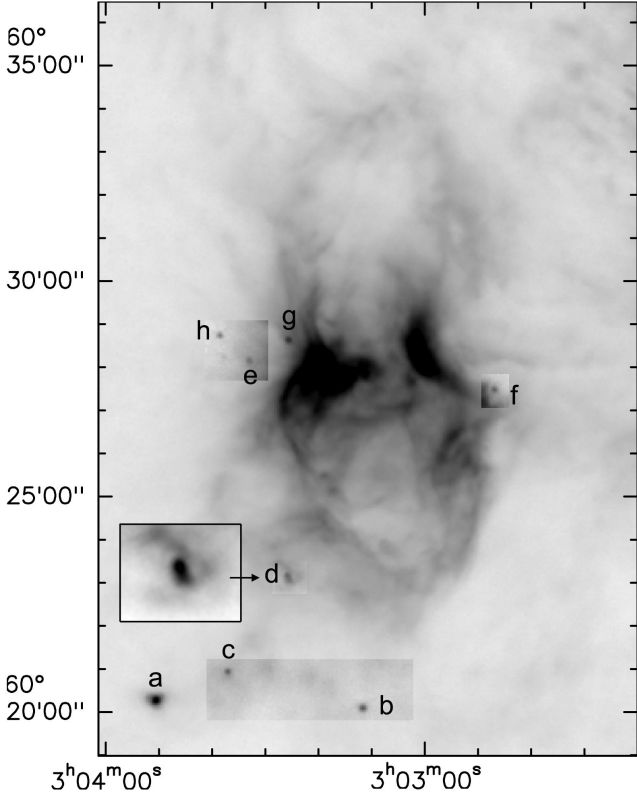


Fig. C.5. Young stellar objects detected at $100\ \mu\text{m}$ in the vicinity of Sh 201. They are identified on the *Herschel*-PACS $100\ \mu\text{m}$ image. (The contrast has been adjusted to better identify the sources.) The sources s1 and s2 are identified on Fig. C.7. The coordinates are J2000.

and $J - K$ colour could correspond to a B5–B7 star with an extinction $A_V \sim 5.6$ mag.)

Appendix E: Isolated point sources

There are some $100\ \mu\text{m}$ point sources in the W5-E field which lie outside the regions discussed previously. We give their measured *Herschel* fluxes in Table E.1. Four of these sources lie outside the coverage of the *Spitzer* images and therefore they have no measured *Spitzer* fluxes. Four other sources have faint *Spitzer* counterparts that have not been measured by KOE08. Another source (#4) is isolated, bright at $8\ \mu\text{m}$ and $24\ \mu\text{m}$, but curiously, is absent in the KOE08 catalogue. When possible we have measured the $24\ \mu\text{m}$ fluxes (these measurements have an asterisk in Table E.1). We attempt to classify them, using the spectral index of the SED between $24\ \mu\text{m}$ and $100\ \mu\text{m}$ (definition in Sect. 5). According to this metric, sources #1, #4, #5 (in agreement with KOE08 classification), #6, #7, #8, and #9 are Class I sources., sources #11 is a flat spectrum source, and source #12 is a Class II source (again in agreement with KOE08).

When possible we used the SED fitting tool of ROB07, assuming an interstellar visual extinction of 1.96 mag. The results are the following (see also Table E.2):

- Isolated-5: This source lies at the end of a small globule, not far from an IF and its associated PDR (Fig. E.1). It is possibly the head of a pillar separated from the PDR (similar to what is observed in the region of the pillars to the south-east of W5-E; see Sect. 6.2). The emission we detect at $100\ \mu\text{m}$ may also be the counterpart to the Class I YSO identified by KOE08. Five

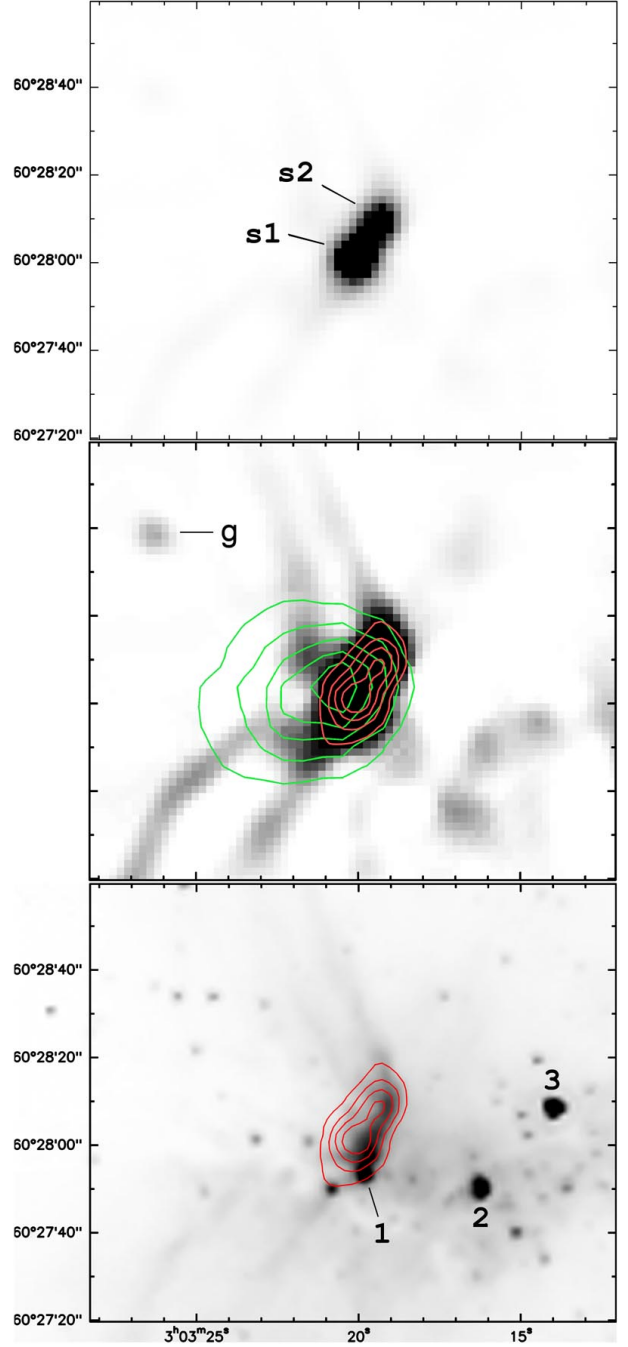


Fig. C.7. Young sources at the waist of Sh 201. *Top:* Unsharp-mask $100\ \mu\text{m}$ image of the eastern portion of the waist of Sh 201 showing the presence of two point sources. *Middle:* Same but deeper image with the column density contours in green showing the dense condensation east of Sh 201. The column density contour levels are 10^{22} , 3×10^{22} , 4×10^{22} , 5×10^{22} , and $6 \times 10^{22}\ \text{cm}^{-2}$; the red contours are for the $100\ \mu\text{m}$ emission (levels of 0.5, 1.0, 1.5, and $2.0\ \text{Jy/beam}$). *Bottom:* The $100\ \mu\text{m}$ contours are superimposed on the *Spitzer* $4.5\ \mu\text{m}$ image showing the stars discussed in the text. The coordinates are J2000.

models satisfy $\chi^2 - \chi^2$ (best) per data point ≤ 3 , all of which point to a low-mass Stage I source. But we cannot exclude that we are dealing with two distinct sources, a Class I YSO close to or inside a globule. Using only the *Herschel* fluxes and the modified blackbody model we estimate a mass of $1.7\ M_\odot$ for the globule's head or the YSO's envelope (Table 5).

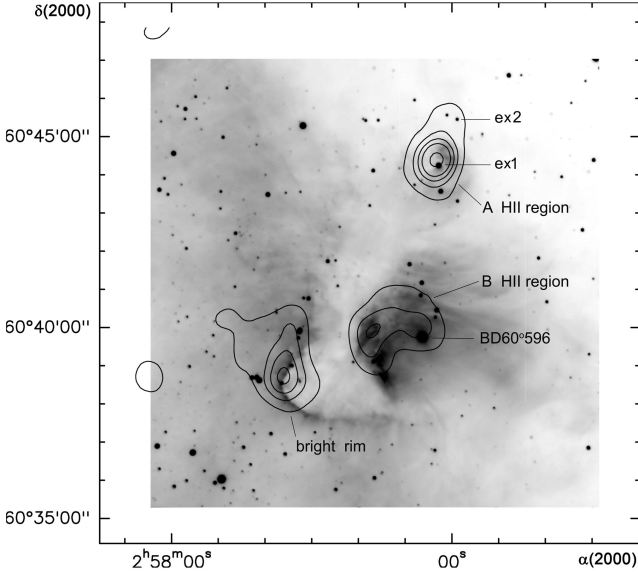


Fig. D.1. The ionized gas in the region between W5-E and W5-W. The NVSS radio-continuum contours at 20-cm (contour levels 0.002, 0.005, 0.0075, 0.01, and 0.015 Jy/beam) are superimposed on a grey-scale image of the $H\alpha$ emission. The $H\alpha$ image has been obtained at the 1.2-m telescope of the Observatoire de Haute Provence. Two H II regions are present in this field, excited by the stars ex1 and BD+60°596. Radio continuum emission is also observed at the border of a bright rim.

tain. A better fit is obtained if we let free the external extinction (up to 10 mag). This source is clearly a Stage II YSO (of intermediate mass), as all the models give an envelope accretion rate equal to zero and a massive disk.

- **Isolated-6:** As shown by Fig. E.2 the SED is not well-constrained. A large range of masses and ages are possible for the central source. However, the fact that the envelope accretion rate is never zero, possibly disagrees with KOE08 classification as a Class II.

- **Isolated-11:** According to KOE08 this source is a Class I YSO. It is a bright *Spitzer* and *Herschel* source, point-like at $100\ \mu\text{m}$ (as shown by Fig. E.1). Only one model satisfies $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 for the whole range of wavelengths. Its parameters, given in Table E.2, correspond to an intermediate mass Stage II/III source. Fig. E.2, however, shows that the fit is not satisfactory. When we use only the $100\ \mu\text{m}$ flux the fit is much better (SED of isolated-11*). In this case the SED corresponds to a Stage I YSO of about $1\ M_{\odot}$. As shown by Fig. E.1, the source lies near an IF, and is surrounded by a small bright rim. One possible explanation is that, at *Herschel*-SPIRE wavelengths we see the emission of the cold condensation enclosed by the bright rim, whereas at *Spitzer* and at $100\ \mu\text{m}$ we see the Class I YSO, possibly a Stage I YSO recently formed inside this condensation. Better spatial resolution in the FIR is necessary to confirm this interpretation.

- **Isolated-12:** Several very similar models satisfy $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 and therefore we give only the best one in Table E.2. All point to a Stage II source, in agreement with KOE08 classification. This is an intermediate mass YSO, with no accreting envelope.

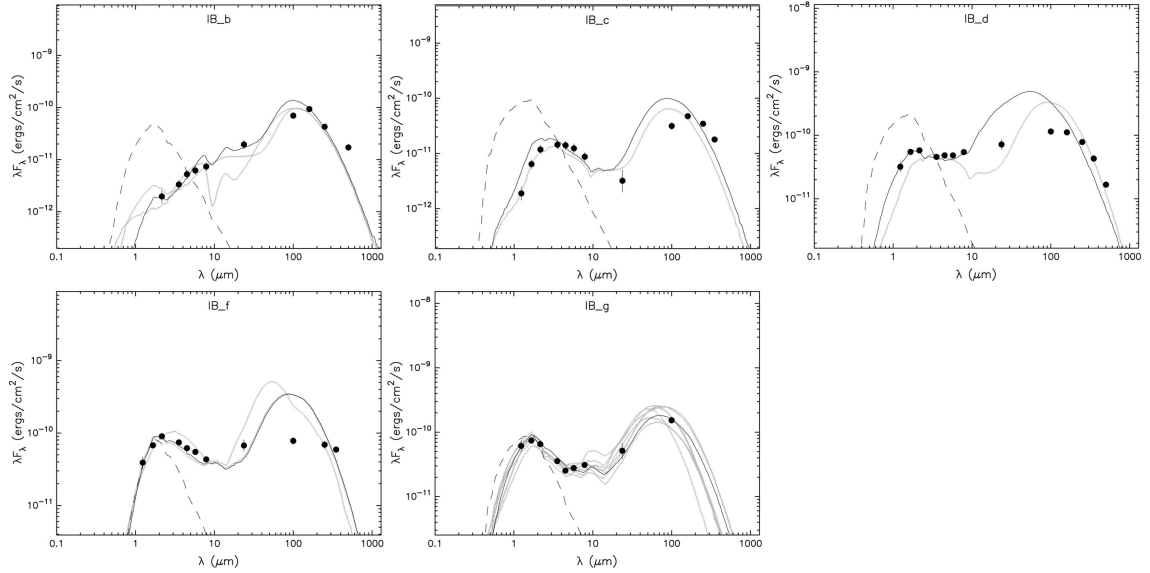
- **Isolated-14bis:** This source lies in the direction of PDR 2. It is classified as a Class II YSO by KOE08. It is bright at $24\ \mu\text{m}$ but its $100\ \mu\text{m}$ counterpart is rather faint (Fig. E.1). A Class I YSO lies nearby that we call isolated-14 (isolated-14 is very faint at $24\ \mu\text{m}$ and has no $100\ \mu\text{m}$ counterpart; thus, its classification is uncertain). For isolated-14bis several models satisfy $\chi^2 - \chi^2(\text{best})$ per data point ≤ 3 , which span a large range of parameters. The classification of isolated-14bis is therefore uncer-

Table D.1. Photometry of point sources: the region between W5-E and W5-W

Name	#	RA(2000) ($^{\circ}$)	Dec(2000) ($^{\circ}$)	<i>J</i> mag	<i>H</i> mag	<i>K</i> mag	[3.6] mag	[4.5] mag	[5.8] mag	[8.0] mag	[24] mag	Class	S(100) Jy	S(160) Jy	S(250) Jy	S(350) Jy	S(500) Jy
a		44.11650	60.75322										1.12	2.53:	3.42:	3.03:	
a1	12110	44.111756	60.751366	15.71	13.87	12.72	11.29	10.58	9.34	7.54	3.73	I					
b	12334	44.174880	60.759755			14.18	12.13	10.90	9.97	8.80	4.16	I	2.35	5.03	3.59:	2.89:	
c	12297	44.164765	60.700052	15.78	13.65	12.24	10.54	9.83	9.22	8.62	6.13	II	1.06	2.54	2.90	2.10:	
d	12648	44.263389	60.694426	12.71	11.32	10.51	9.29	8.49	7.74	6.63	2.75	I	3.85	5.97	6.58	5.06:	2.8:
e		44.25698	60.68601										3.21	9.76	6.70:	3.30:	2.0:
e1	12641	44.259853	60.686344			13.70	11.54	10.54	9.78	9.01		I					
e2	12620	44.254253	60.686155	15.17	13.60	12.72	11.49	10.90				II					
e3	12632	44.257101	60.687345			13.47	12.02	11.33	11.18	10.48		II					
f	12922	44.339923	60.688855	12.49	11.09	10.03	8.76	8.22	7.60	6.88	2.82	II	2.62	5.80	6.94:		
g	13165	44.407821	60.642469	12.00	10.99	10.38	9.56	9.19	8.34	7.24	3.11	II	5.15:				
ex1	12666	44.268587	60.736758	10.29	10.01	9.84	9.76	9.59	9.13	7.76	1.91	II					
ex2	12609	44.251992	60.756720	11.13	10.44	10.06	9.87	9.80	9.68	9.16		III					
BD+60°596	12717	44.283529	60.662250	8.80	8.72	8.72	8.79	8.74	8.51	7.42		III					
1	12234	44.143079	60.756463			14.63	12.23	11.10	10.23	9.45	5.94	I					
2	12313	44.170433	60.758392			14.89	12.49	11.28	10.48	10.05	5.98	I					
3	12381	44.185222	60.760892				14.56	13.27	11.98	11.29	7.08	I					
4	12631	44.256847	60.655770	17.40	16.31	14.47	11.84	10.80	9.95	9.01	5.64	I					
5	13019	44.366178	60.650644		15.68	13.95	11.90	11.16	10.18	9.12		I					
6	13118	44.392491	60.643065	12.57		11.22	10.48	10.24	7.77	5.82	2.98	I					
s1	13025	44.367277	60.695293	12.27	11.86	11.52	10.91	10.59	10.49	10.35	4.46	II					

Table D.2. Parameters of the YSOs in the region between W5-E and W5-W, obtained with the SED fitting tool of ROB07

Name	χ^2	M_* (M_{\odot})	T_* (K)	M_{disk} (M_{\odot})	\dot{M}_{disk} ($M_{\odot} \text{ yr}^{-1}$)	\dot{M}_{env} ($M_{\odot} \text{ yr}^{-1}$)	L (L_{\odot})	i ($^{\circ}$)	Stage
b	116 (−146)	0.9 (0.7–1.4)	3920	6.3e-3 (1.0e-3–1.1e-2)	4.9e-7 (−1.0e-7)	2.6e-4 (−4.7e-4)	19 (14–21)	18 (−41)	I
c	116 (−177)	1.9 (−1.5)	4300	2.1e-4 (−8.6e-4)	5.6e-10 (−2.0e-8)	2.4e-4 (−1.4e-4)	29 (−20)	49 (56)	I
d (1.96 mag)	273 (−289)	1.9 (3.3)	4150	2.6e-1 (−8.4e-1)	2.5e-6 (−6.4e-9)	1.2e-4 (−6.2e-4)	69 (−80)	18 (41)	I
d (10.38 mag)	199	3.2	4420	8.4e-4	6.4e-9	6.2e-4	80	31	I
g	16 (−43)	1.5 (−5.0)	4224	1.7e-2 (3.2e-4–1.3e-1)	1.5e-6 (4.5e-11–2.7e-6)	1.5e-4 (1.4e-5–5.9e-4)	30 (−120)	18 (−75)	I?

**Fig. D.5.** Spectral energy distributions of point sources in the region between W5-E and W5-W. We used the SED fitting tool of ROB07.

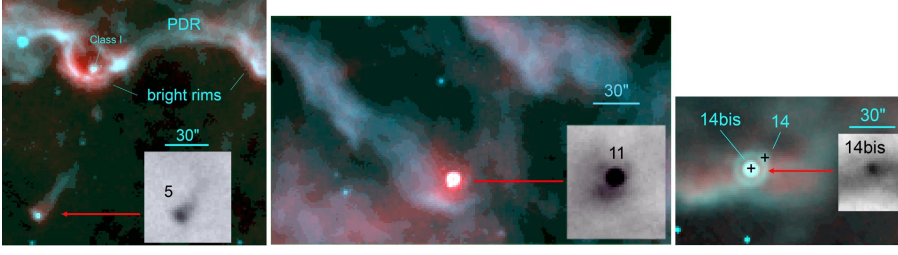


Fig. E.1. Colour composite images of three isolated sources discussed in the text. Red is the *Herschel* 100 μm emission and turquoise is for *Spitzer* 8.0 μm emission (isolated 5 and 11) or the 24 μm emission (isolated 14bis). The grey inserts show, at the same scale, the 100 μm source.

Table E.1. Photometry of isolated point sources detected at 100 μm

Name	#	RA(2000) ($^{\circ}$)	Dec(2000) ($^{\circ}$)	<i>J</i> mag	<i>H</i> mag	<i>K</i> mag	[3.6] mag	[4.5] mag	[5.8] mag	[8.0] mag	[24] mag	Class	S(100) Jy	S(160) Jy	S(250) Jy	S(350) Jy	S(500) Jy
1		44.96406	59.98754								7.00*		0.36	0.53:			
2		46.04535	60.59474									outside <i>Spitzer</i>	0.64	1.04	0.39:		
3		45.86957	60.64833									outside <i>Spitzer</i>	0.27				
4		45.72951	60.65243								4.23*		1.30	1.07	0.49:		
5	14148	44.65653	60.69825		15.85	14.03	12.08	11.13	10.19	9.18	5.44	I	0.57	1.33	0.91	0.58:	
6	17344	45.62620	60.69937		15.57	14.72	13.46	12.93	12.08	10.12	6.18*	II	0.32				
7		45.56116	60.70956								6.41*		0.44	0.53:			
8		45.23067	60.75611								5.86*		0.31	0.53:			
9		44.68234	60.78960								6.80*		0.84	0.99	0.59:		
10		46.30267	60.81229									outside <i>Spitzer</i>	0.68	1.01	0.99:		
11	12921	44.33975	60.88602	14.07	12.81	11.79	9.85	8.93	8.04	6.62	2.48	I	2.69	5.21	4.76	3.19	1.69:
12	14630	44.77136	60.90115	10.82	9.65	8.67	7.37	6.75	6.04	4.71	1.61	II	1.07	0.91	0.45:		
13		45.81908	60.96289									outside <i>Spitzer</i>	0.78	1.88	1.48	1.45:	
14	15939	45.06133	60.99216				13.19	12.32	11.60	10.62	7.15	I					
14bis	15965	45.06639	60.99005	14.73	12.94	11.59	9.90	9.13	8.61	7.56	4.10 (3.92*)	II	0.37				

Table E.2. Parameters of the isolated YSOs, obtained with the SED fitting tool of ROB07

Name	χ^2	M_* (M_{\odot})	T_* (K)	M_{disk} (M_{\odot})	\dot{M}_{disk} ($M_{\odot} \text{ yr}^{-1}$)	\dot{M}_{env} ($M_{\odot} \text{ yr}^{-1}$)	L (L_{\odot})	i ($^{\circ}$)	Stage
5	62 (−95)	0.7 (0.4–0.9)	3900	8.6e-4 (5.1e-4–6.6e-2)	1.9e-9 (−1.7e-7)	1.7e-4 (−2.3e-5)	9 (6–11)	31 (18–69)	I?
11	259	4.6	11300	2.9e-4	9.1e-11	1.5e-7	440	87	II/III?
11*	21 (−39)	0.8 (−1.4)	3931	6.6e-2 (3.4e-3–1.0e-1)	5.0e-6 (−3.4e-9)	4.7e-5 (1.5e-5–6.9e-5)	26 (21–35)	18	I
12	26 (−59)	4.9	16200	7.3e-2	9.9e-7	0	477	18	II
14bis ($A_V = 1.96$ mag)	25 (−52)	3.8 (1.9–7.9)	13940	1.3e-2 (2.0e-4–1.4e-1)	3.8e-8 (1.5e-10–1.1e-7)	0 (1.7e-5)	179 (34–2730)	81 (−69)	II
14bis ($A_V = 8.7$ mag)	2.7 (−29.7)	3.1 (−2.6)	12215 (−9400)	7.3e-2 (−2.7e-3)	9.0e-8 (7.0e-10–1.2e-7)	0	87 (−43)	75 (−18)	II

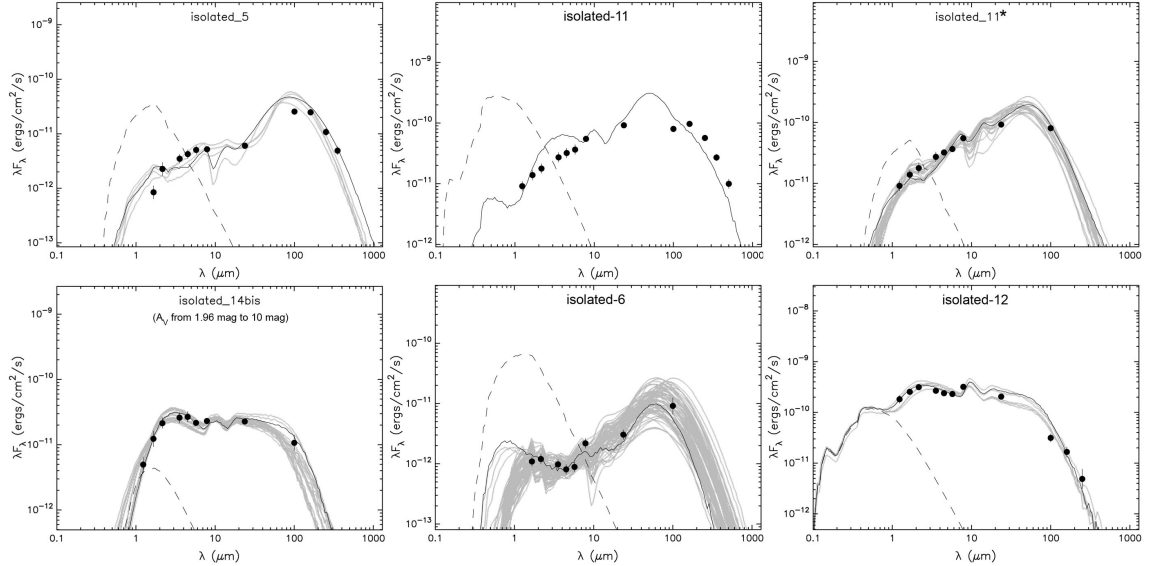


Fig. E.2. Spectral energy distributions of a few isolated point sources. We used the SED fitting-tool of ROB07. (In isolated-11* we use only the flux at 100 μm .)